

DISCOVERY OF 28 OPEN CLUSTERS WITH *GAIA* DR3

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RESUMEN

En este trabajo buscamos cúmulos abiertos en regiones de Formación Estelar en la Galaxia utilizando el algoritmo HDBSCAN aplicado a los datos astrométricos del *Gaia* Catálogo DR3. El resultado fue la identificación de 28 nuevos cúmulos abiertos, cuya existencia real está respaldada por la probabilidad de pertenencia determinada a partir de datos astrométricos y por la presencia de secuencias de cúmulos en sus diagramas de color y magnitud, lo que permitió un ajuste isócrono confiable. De los cúmulos abiertos identificados, 3 tienen edades menores de 50 Myr, 19 son de edad intermedia y 6 son cúmulos antiguos. Los cúmulos tienen radios aparentes que varían de 3 a 20 minutos de arco. Para todos los cúmulos, estimamos el movimiento propio medio, el paralaje medio y los parámetros fundamentales considerando las estrellas miembros de cada cúmulo. Uno de los cúmulos descubiertos está a una distancia de 10 kpc y tiene un $\log(\text{edad}) \approx 10,1$, lo que lo convierte en uno de los cúmulos abiertos catalogados más distantes y antiguos.

ABSTRACT

In this work we searched for open clusters in the fields of Star-Forming regions in the Galaxy using the HDBSCAN applied on the astrometric data from the *Gaia* DR3 catalog. It resulted in the identification of 28 new open clusters, of which the real existence is supported by the membership probability determined from astrometric data and by the presence of cluster sequences in their color-magnitude diagrams, which allowed for a reliable isochrone fit. Of the open clusters identified, 3 are younger than 50 Myr, 19 are of intermediate age, and 6 are old clusters. The clusters have apparent radii ranging from 3 to 20 arcmin. For all clusters, we estimate mean proper motion, mean parallax, and fundamental parameters considering the member stars for each cluster. One of the discovered clusters is at a distance of 10 kpc and has a $\log(\text{age}) \approx 10.1$, making it one of the most distant and oldest cataloged open clusters.

Key Words: open clusters and associations: general

1. INTRODUCTION

Open clusters (OCs) are key objects used to study the structure and dynamics of the Galaxy since their kinematics, distances, and ages can be determined with good precision from the properties of their member stars.

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Before the *Gaia* era, the most widely used catalogs of OCs and their fundamental parameters were the New Catalog of Open Clusters and Candidates (Dias et al. 2002) and the Milky Way Star Clusters (Kharchenko et al. 2003), both of which present about 3 thousand objects.

This scenario changed with the new generation of high precision all-sky astrometric *Gaia* catalogs which provided data to limiting magnitude about 21 offering an important opportunity to determine the parameters and members of the known open clusters as well as to discover new ones.

Recently, many works have used clustering algorithms, supervised or not, to search for new open clusters and to identify their member stars: He et al. (2022) used the pyUPMASK algorithm (Pera et al. 2021) with the K-means clustering method, Castro-Ginard et al. (2020), Castro-Ginard et al. (2019) and Hao et al. (2022) used DBSCAN (Ester et al. 1996), Cantat-Gaudin & Anders (2020) used UPMASK procedure (Krone-Martins & Moitinho 2014), Liu & Pang (2019) used the Friend of Friends method, Sim et al. (2019) used Gaussian mixture model and mean-shift algorithms, Jaehnig et al. (2021) used extreme deconvolution Gaussian mixture models and recently Kounkel & Covey (2019), Kounkel et al. (2020), Hunt & Reffert (2021) and Hunt & Reffert (2023) used HDBSCAN to search for open clusters.

As a direct consequence of the publication of the *Gaia* catalogs and the application of advanced machine learning techniques, currently about 8000 open clusters are known and cataloged (Hunt & Reffert 2023), representing an almost fourfold increase in the number of clusters originally cataloged in our database (Dias et al. 2002). However, according to Hunt & Reffert (2024), the Milky Way contains a total of one hundred thousand open clusters, of which only about 4% were discovered.

In this context, we have dedicated special attention to the task of a systematic search for previously unknown open clusters, and in this work we present the first results of a search for these objects in the fields from the catalog of star-forming regions in the Galaxy published by Avedisova (2002).

In the next section, we describe the procedures adopted in the search for clusters and the tools used in their analysis. Section 3 is dedicated to the presentation of the isochrone fitting procedure to determine distance and age, which indicate that the clusters are real. In Section 4 we discuss the results. Finally, in Sect. 5 we give some concluding remarks.

2. CLUSTER SEARCH AND ANALYSIS STRATEGY

In this work we searched for open clusters in the fields of one degree squared centered in the star formation regions provided in the catalog published by Avedisova (2002). Basically we applied the HDBSCAN (Hierarchical Density-Based Spatial Clustering of Applications with Noise) clustering algorithm (Campello et al. 2015, 2013) in two steps, firstly to find all the clusters in each field and subsequently to estimate the membership for stars in open clusters. These procedures are described below.

2.1. *Searching for clusters*

The HDBSCAN algorithm is an interesting method for searching for open clusters since it can identify clusters with different shapes and density, being efficient and relatively fast in dealing with large volumes of data, as is the case with the Gaia catalog. Some advantages of the method that make it more robust are its ability to deal with multidimensional data and it does not require the number of clusters to be initially defined.

As an evolution of the DBSCAN algorithm (Ester et al. 1996) it only requires the definition of the parameters `min_cluster_size` (minimal size to consider a cluster) and `min_samples` (primarily controls how tolerant the algorithm is towards noise). The other parameter is the `cluster_selection_method` used to select the clusters from the cluster tree hierarchy. While EOM (Excess of Mass) method produces clusters with large areas, the leaf method produces small homogeneous clusters.

However, the obtained results from HDBSCAN depends on the data as well as the setup defined to use the algorithm. We refer the reader to the work of Hunt & Reffert (2021) to search for open clusters for a complete description and use of the HDBSCAN algorithm, including tests with different setup and comparison of their performance with other methods.

In this work we applied the python code of the HDBSCAN² in an unsupervised fashion, to find all the clusters in each field, considering a five-dimensional astrometric space (position in the tangent plane, parallax, and proper motions). We followed the results of Hunt & Reffert (2021) and set the parameters `min_cluster_size` = 15, `min_samples` = 80 and `method` = leaf. We also scaled the data set to have a median of zero and a unit inter-quartile range using a RobustScaler object from scikit-learn (Pedregosa et al. 2011). Clusters with fewer than 15 stars were disregarded in our analyses.

In the Figure 1 we present the typical results for this procedure using the field 213.07 – 2.23 as an example, where four groups were found. In the plot, the member stars of each cluster as originally defined by the algorithm are presented in different colors. In this field three known clusters were found that were colored red for CWNU_1774, green for UBC_212 and cyan for FSR_1077 and one unknown detected open cluster (Dias_21) presented in blue color.

The list of detected clusters was cross-referenced with the most recent publications on new open clusters, as well as the catalog published by Hunt & Reffert (2023). In this process, we consider as new open cluster candidates those in which the central coordinates differed more than 5 arcmin, mean proper motion differed more than 2σ , and mean parallax differed more than 5σ (to consider large errors for more distant stars) of the known open clusters cataloged.

Subsequently, we used Aladin Sky Atlas (Bonnarel et al. 2000) which provides tools that help ensure that the member stars of these new candidates did not correspond to members of known clusters, since these new open clusters

²<https://hdbscan.readthedocs.io/en/latest/>

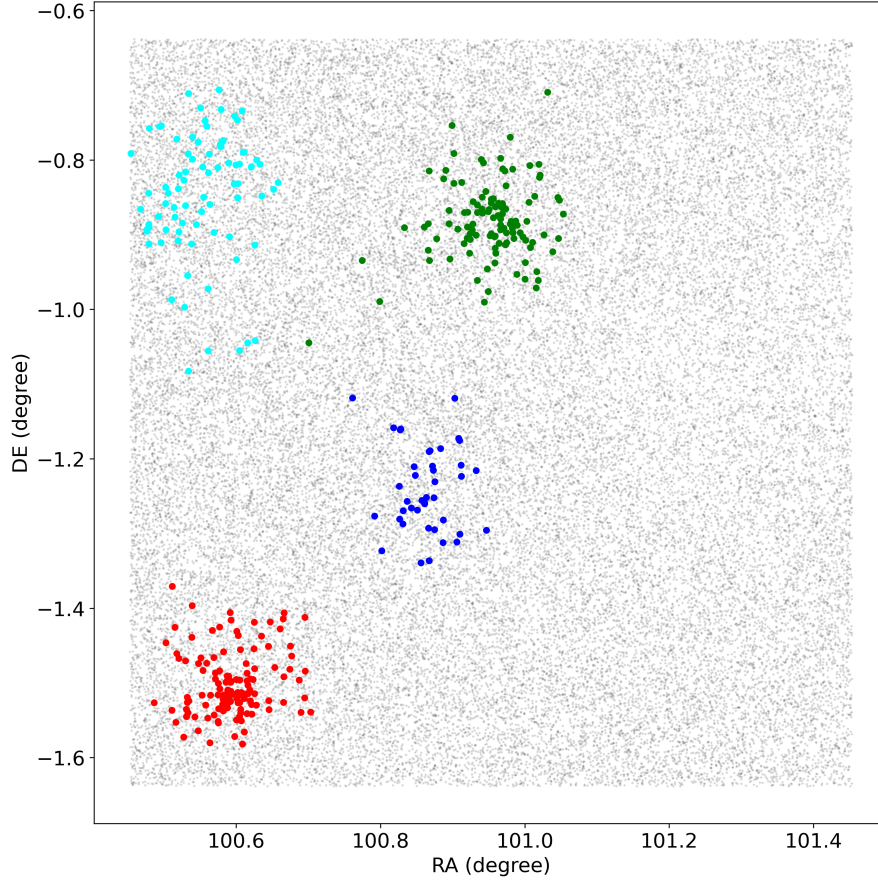


Fig. 1. The sky map with field of one degree centered in the field 213.07 – 2.23 is presented. In the plot, stars members as originally defined by the HDBSCAN algorithm, of each cluster are shown in different colors. In red for CWN1774, green for UBC212 and cyan for FSR1077. The stars in blue are members of Dias21, new detected open cluster.

can be in the same field (or close) to known open clusters, as is the case of Dias21 in the field of the three known clusters as presented in Figure 1.

The candidates listed as uncataloged open clusters had their color-magnitude diagram (CMD) checked by a visual inspection, and only those with a clear typical cluster sequence in the CMD were kept. We selected 28 unknown open cluster candidates and, for each of these cases, performed a more detailed analysis to determine the member stars to apply the isochrone fit and determine their distances and ages.

2.2. Memberships

A recent addition to the HDBSCAN library is the introduction of soft clustering³, which capitalizes on the smoothed density function provided by the

³https://hdbscan.readthedocs.io/en/latest/soft_clustering.html

condensed tree over the data points. In this approach, cluster labels are not directly assigned to points; rather, each point receives a probability assignment. However, the probability calculations still rely on specified parameters such as `min_cluster_size` and `min_samples`. Here we adopt this probability estimation as the membership of the stars.

Since the results provided by the HDBSCAN algorithm depend on the parameters we applied the code in a different way to ensure an unbiased membership determination for stars in open clusters. In this approach we use a Monte Carlo method to sample the space of the clustering algorithm tuning parameters to capture variations in membership assignment. The clustering algorithm is run 50 times, each time with varying tuning parameters, resulting in a probability estimation distribution for individual stars. Averaging probabilities from the N runs yields a final membership probability for each star. Furthermore, the variance of individual probability estimations is computed and stars that deviate by more than 100% are assigned a probability of zero. This approach eliminates the need for initial assumptions about the HDBSCAN tuning parameters, thus improving the accuracy of the cluster analysis.

The 28 new open clusters candidates were reanalyzed through the HDBSCAN algorithm in this modified fashion, applied to *Gaia* DR3 astrometric data of the stars located in the area delimited by the estimated central coordinates and two times the apparent radius estimated considering the stars selected from the HDBSCAN applied as described in the Section 2.1.

3. DISTANCES AND AGES

It is well known that the stars of a real open cluster align along a specific feature, the cluster sequence, in a CMD. This feature is most evident when only stars are included with a sufficiently high stellar membership probability, e.g. as determined by the method described above. Likewise, the stars that form this feature should exhibit an over-density in a 3D plot with proper motion and parallax data, since it is assumed that they occupy a limited volume in space and have similar velocities. In this scenario, a precise isochrone fit to the stars with membership probability greater than 0.51 is a strong indicator for the presence of a real open cluster and allows for determining its age, extinction, and metallicity. In addition, it yields an estimate of the distance of the cluster independent of the parallax measurements.

Consequently, the next step in our analysis was to use *Gaia* DR3 G_{BP} and G_{RP} magnitudes to perform an isochrone fit considering the astrometric membership probabilities of the stars derived by the method outlined in Section 2.2.

In previous works, we commented on the subjectivity of visual fitting procedures and the dependence of the results on the knowledge of the stellar membership in the analyzed field. Generally, visual isochrone fits do not objectively correspond to the mathematical best fit, for example, using maximum likelihood methods, possibly introducing a bias in the results. For this reason,

we applied the cross-entropy (CE) method to fit theoretical isochrones, as detailed in Monteiro et al. (2017). This approach was previously successfully applied to *Gaia* DR2 data in Dias et al. (2018) and Monteiro & Dias (2019) and to hundreds of clusters using the *Gaia* eDR3 data (Dias et al. 2021).

The CE method is an iterative statistical procedure where in each iteration the initial sample of the fit parameters is randomly generated using predefined criteria. The code then selects the 10% best fits based on the calculated weighted likelihood values, taking into account the probabilities of astrometric membership. The likelihood function is used to define the objective function, which is then minimized with respect to the parameters. In our code, a prior in A_V is adopted as a normal distribution with μ and variance for each cluster taken from the 3D extinction map produced by Capitanio et al. (2017). The optimization algorithm then minimizes with respect to the parameters. More details can be found in Monteiro et al. (2010) and Monteiro et al. (2017).

Basically, a synthetic cluster is generated using the Padova PARSEC version 1.2S database of stellar evolutionary tracks and isochrones (Bressan et al. 2012), which uses the *Gaia* filter band passes of Evans et al. (2018) and is scaled to the solar metal content with $Z_{\odot} = 0.0152$. The code scans the following parameter space limits:

- age: from $\log(\text{age}) = 6.60$ to $\log(\text{age}) = 10.15$;
- distance: from 1 to 25000 parsec;
- A_V : from 0.0 to 5.0 mag;
- $[Fe/H]$: from -0.90 to 0.70 dex

Our code accounts for reddening by including extinction in the generated synthetic cluster, which is then compared to the data through the likelihood function. For each star of the synthetic cluster, we use the A_{λ}/A_V relations of Monteiro et al. (2020) updated to use the *Gaia* DR3 filter band passes, to obtain the observed photometry for the generated synthetic cluster. In this way, each generated star is reddened according to its color and then compared with the observational data through the likelihood in the optimization process.

4. RESULTS

In this study, we searched for open clusters in 2960 fields of the star formation catalog of Avedisova (2002) using HDBSCAN algorithm applied on *Gaia* DR3 astrometric data with the tuning parameters described in Section 2.1 that allowed us to find 1620 previously cataloged open clusters and 3535 new candidates. Our analysis, including visual check of the features in the CMD, allowed us to identify 28 new open clusters.

We would like to note that this work did not exhaust the possibilities of searching for open clusters in these fields, since changing the tuning parameters of the HDBSCAN may favor finding clusters with different characteristics

from the ones we found. We also chose to provide only the list with the best candidates to avoid introducing noise in the catalogs of known open clusters.

In the Figure 2 we show the typical CMD with the membership probabilities of the stars and the optimal isochrone fit determined by the CE optimization algorithm, indicating the reality of the clusters. In the plots, all the stars in the investigated field are presented in light gray, and the member stars are represented by a point color, with the redder color referring to a higher membership probability. The CMDs show the expected characteristics of a field with an open cluster, exhibiting a sequence typical of the presence of an open cluster relatively separate from the sequence of field stars. Note that, as expected, and found by the HDBSCAN in each case, there is a clump of high membership probability stars in the vector proper motion diagram and position in the sky (in this work positions in the tangent plane), with more stars with higher membership closer to the center of the field. The CMDs of all the discovered clusters are shown in Figures A1, A2 and A3.

From a quantitative perspective, these cases exhibited typical errors in the average astrometric parameters, as well as distances and ages, compared to the values reported in the literature, particularly those published in Dias et al. (2021), where we used the same isochrone fitting technique. As expected, the clusters with greater uncertainty in the distance and age parameters present a larger spread in the turn-off and main sequence.

Table 1 summarizes the positions, sizes, and astrometric parameters of these 28 discovered open clusters determined using stars with a membership probability greater than 0.50. The mean radial velocities are provided for 14 clusters using the radial velocity data given in *Gaia* DR3 catalog. Table 2 provides the results of the fundamental parameters (distance, age, $[\text{Fe}/\text{H}]$ and A_V) obtained by the isochrone fit given by the mean and standard deviation of the results of ten runs of the fitting procedure, as detailed in last section. In addition to the cluster parameters, all membership probabilities and individual stellar data are also made electronically available.

5. DISCUSSIONS

The superposition of known open clusters in their projections on the sky can occur with some frequency, especially in regions rich in stars and clusters. However, depending on the situation, the identification of new clusters can be obscured by a denser open cluster in the field, as for example discussed by Hunt & Reffert (2023) for the case of HSC_2382 in the field of IC_2662, as well as new star clusters discovered in the field of the open cluster NGC_5999 by Ferreira et al. (2019) and around NGC_6649 by Gao (2024).

In this work, five discovered clusters are located in the field of known open clusters. Our investigation of the cross-match with the clusters published in Hunt's catalog, as well as the mean astrometric parameters and distance and age estimated from the isochrone fit, ensure that this is not a rediscovery of the same known open cluster. The new clusters in the same field of known clusters are the following: Dias_22 is about 13 arcmin from Haffner_6, Dias_23

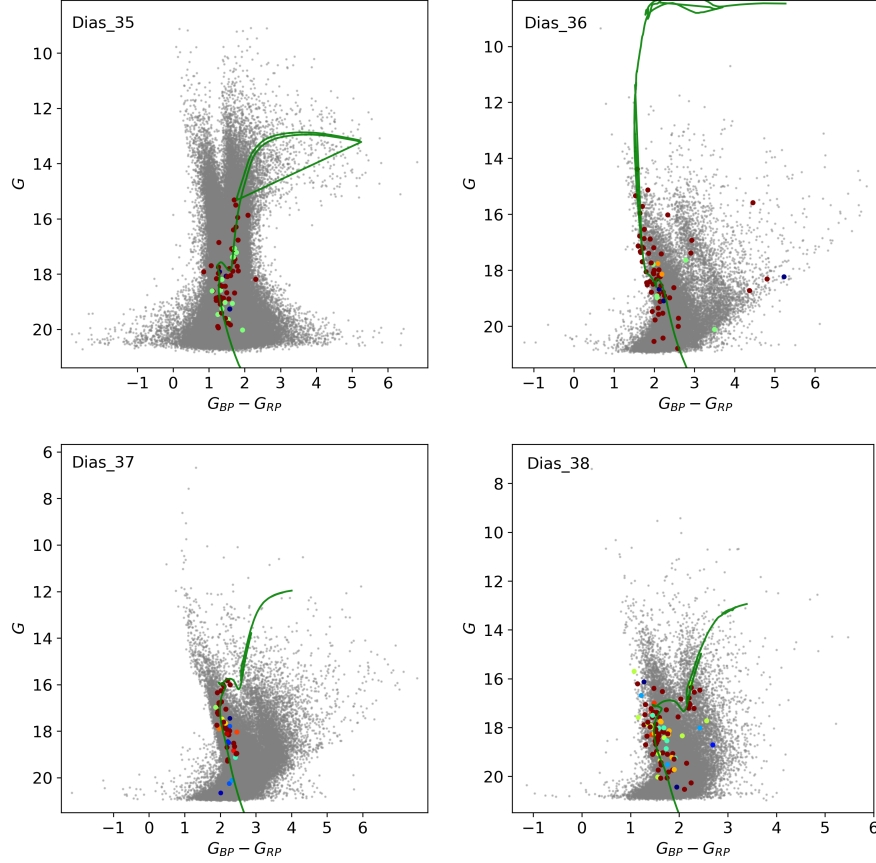


Fig. 2. Color-magnitude diagram with isochrone fitted for some of the discovered clusters studied. The color is proportional to the astrometric membership, where more red is higher membership. The gray points correspond to field stars. The CMDs of all clusters in the sample are shown in Appendix 8.

is 12 arcmin from UBC_640, Dias_31 is 19 arcmin from Ruprecht_171, Dias_32 is 16 arcmin from Ruprecht_143 and Dias_34 is situated about 8 arcmin from the center of HSC_320.

Dias_38 and Dias_39 are two OCs found in the same investigated field with angular separation of 20 arcmin. In the Figure 3 the sky map and the CMD of the investigated field 98.51 ± 3.31 of one degree square are given with the 98 and 44 member stars of OC Dias_38 and Dias_39 superimposed in green and blue, respectively. As can be seen in Table 1 there is no significant difference in the astrometric parameters, but according to the results of our isochrone fit, over-plotted in the CMD in the Figure 3, the clusters have about the same age, however Dias_38 is about 1500 pc more distant than Dias_39. Unfortunately, we did not find members with determined radial velocity that would be useful for determining the Galactic orbits of the objects. This could be an interesting case for checking the possibility that the two are binary

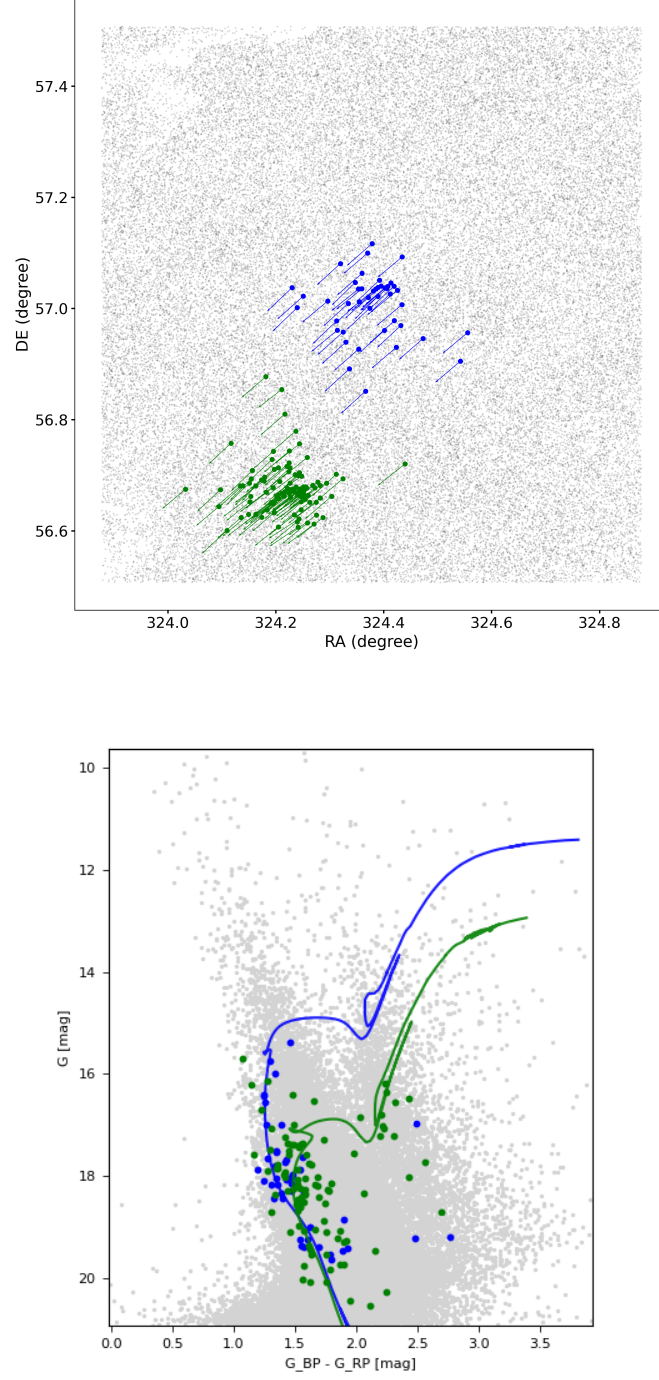


Fig. 3. Sky map and color-magnitude diagram of the stars in the field $98.51 + 3.31$ with the member stars of the clusters Dias_38 and Dias_39 plotted in green and blue, respectively. The length of the vectors are presented in arbitrary units. The isochrones in the CMD are those with results given in the Table 2.

TABLE 1

MEAN ASTROMETRIC PARAMETERS FOR 28 DISCOVERED OPEN CLUSTERS.

field	name	RA.ICRS [deg]	DE.ICRS [deg]	R [arcmin]	N	pmRA [mas yr ⁻¹]	e_pmRA [mas yr ⁻¹]	pmDE [mas yr ⁻¹]	e_pmDE [mas yr ⁻¹]	Plx [mas]	e_Plx [mas]	RV [km s ⁻¹]	e_RV [km s ⁻¹]	NRV
121.37+3.47	Dias_12	9.2993	66.3245	12.681	79	-1.540	0.227	-0.006	0.163	0.159	0.087	-91.057	2.089	2
123.93+1.43	Dias_13	15.0612	64.3680	9.861	54	-1.603	0.228	-0.114	0.114	0.162	0.061	...	...	0
131.86+1.33	Dias_14	33.1042	62.5949	19.458	141	-0.921	0.238	0.092	0.302	0.213	0.056	-60.781	0.830	2
137.69+3.88	Dias_15	45.9315	62.9139	19.751	85	-0.089	0.138	0.056	0.209	0.188	0.050	-33.649	2.538	2
149.59+0.90	Dias_16	61.9305	53.5426	14.053	52	0.116	0.244	0.326	0.121	0.161	0.043	...	...	0
151.63+3.36	Dias_17	66.5584	53.4130	3.394	31	0.108	0.174	-1.044	0.136	0.227	0.070	...	...	0
180.73-2.51	Dias_18	84.2710	26.5927	6.774	43	0.340	0.138	-0.306	0.186	0.220	0.076	-11.718	2.078	2
184.89-1.74	Dias_19	87.7687	23.6001	8.012	49	0.398	0.141	-0.632	0.274	0.197	0.091	...	...	0
203.71+1.18	Dias_20	99.5408	8.2739	17.297	44	-0.167	0.146	-0.377	0.163	0.209	0.049	39.631	1.380	5
213.07-2.23	Dias_21	100.8634	-1.2482	6.837	44	-0.423	0.137	0.534	0.100	0.178	0.051	...	...	0
227.75-0.15	Dias_22	110.0193	-13.3799	11.545	40	-0.647	0.083	0.751	0.097	0.226	0.035	...	...	0
245.93+1.16	Dias_23	120.7232	-28.0161	6.765	35	-1.815	0.130	2.482	0.200	0.192	0.035	47.955	2.839	2
251.88-0.47	Dias_24	122.8776	-34.0747	8.613	45	-2.482	0.092	3.008	0.109	0.167	0.026	96.460	3.521	5
251.52+2.00	Dias_25	124.5981	-32.1535	17.975	66	-2.293	0.222	2.760	0.359	0.160	0.023	76.386	3.494	5
273.89-1.58	Dias_26	140.2921	-52.4072	7.759	40	-3.598	0.222	2.928	0.135	0.115	0.049	...	...	0
292.34-4.88	Dias_27	166.1039	-65.0980	16.645	256	-5.232	0.466	2.062	0.249	0.137	0.057	32.308	1.785	7
293.34-0.86	Dias_28	171.5147	-62.1815	4.021	50	-6.135	0.243	1.207	0.162	0.280	0.040	...	...	0
318.06-0.46	Dias_29	224.3485	-59.3583	14.043	147	-4.748	0.724	-3.380	0.332	0.344	0.054	-27.172	1.566	10
354.99+5.17	Dias_30	258.0942	-30.3556	5.505	32	-3.817	0.287	-5.791	0.301	0.084	0.055	2.528	2.009	3
16.28-2.81	Dias_31	277.6833	-16.1074	8.520	31	-1.179	0.193	-3.047	0.241	0.259	0.056	...	...	0
20.29-1.14	Dias_32	278.2984	-11.9047	8.905	148	-0.304	0.242	-1.854	0.291	0.261	0.079	5.205	5.335	4
22.76-0.25	Dias_33	278.6223	-9.3315	7.998	28	-0.834	0.170	-2.938	0.425	0.327	0.023	-4.218	3.239	2
32.74-0.08	Dias_34	283.1168	-0.4266	4.022	41	-1.626	0.106	-4.295	0.087	0.217	0.032	51.251	1.468	6
26.61-3.98	Dias_35	283.6507	-7.5504	16.807	64	-2.782	0.225	-5.609	0.313	0.147	0.038	...	...	0
54.15+0.17	Dias_36	292.4729	19.0487	3.783	66	-2.595	0.157	-5.523	0.181	0.211	0.065	...	...	0
60.57-0.19	Dias_37	296.4799	24.1809	3.782	44	-1.709	0.102	-4.280	0.116	0.267	0.080	...	...	0
98.51+3.31	Dias_38	324.2194	56.6787	11.981	98	-2.526	0.255	-2.044	0.120	0.103	0.044	...	...	0
98.51+3.31	Dias_39	324.3774	57.0073	9.374	44	-2.466	0.101	-2.185	0.084	0.129	0.035	...	...	0

clusters. We understand that these cases need further confirmation.

Finally, we point out that Dias_30 is the older and more distant open cluster candidate of the sample and is also among the oldest and most distant in the list of known open clusters. The CMD, presented in Figure A2, shows that the members selected by the HDBSCAN algorithm turn away from the main sequence towards the giant branch. The absence of a defined main sequence due to possible members beyond the limiting magnitude of the *Gaia* DR3 catalog results in poor parameter estimates by our code. However, we opted to maintain this new candidate in our list, since it is at the limit of detection of the *Gaia* DR3 data.

6. CONCLUSIONS

This paper presents the first results of the ongoing project to search for open clusters using a clustering algorithm applied to *Gaia* data. We report 28 new open clusters of our Galaxy discovered by applying the HDBSCAN algorithm on *Gaia* DR3 data in fields of star formation from the catalog published by Avedisova (2002).

We determined the astrometric membership probability of the stars in each field using the HDBSCAN algorithm applied in a variational fashion, which

TABLE 2

RESULTS OF FUNDAMENTAL PARAMETERS OBTAINED FROM THE ISOCHRONE FIT.

field	cluster	dist	e_dist	logt(age)	e.log(age)	Av	e_Av	FeH	e_FeH
		[pc]	[pc]			[mag]	[mag]	[dex]	[dex]
121.37+3.47	Dias_12	6069	1506	8.939	0.405	4.010	0.395	-0.181	0.063
123.93+1.43	Dias_13	3120	1180	8.770	0.731	3.064	0.349	-0.333	0.101
131.86+1.33	Dias_14	3546	690	8.357	0.494	2.407	0.223	-0.208	0.042
137.69+3.88	Dias_15	6403	1310	9.027	0.218	2.743	0.283	-0.181	0.130
149.59+0.90	Dias_16	3834	760	8.805	0.312	2.616	0.183	-0.361	0.116
151.63+3.36	Dias_17	3137	615	8.225	0.763	2.480	0.220	-0.295	0.034
180.73-2.51	Dias_18	3733	413	8.659	0.238	2.403	0.241	-0.263	0.101
184.89-1.74	Dias_19	5193	557	8.787	0.723	2.545	0.164	-0.269	0.137
203.71+1.18	Dias_20	4678	761	8.882	0.197	2.197	0.187	-0.304	0.097
213.07-2.23	Dias_21	4295	225	7.730	0.209	1.973	0.058	-0.227	0.088
227.75-0.15	Dias_22	3504	236	8.327	0.430	1.614	0.118	-0.206	0.074
245.93+1.16	Dias_23	5563	545	9.210	0.292	1.246	0.305	-0.183	0.053
251.88-0.47	Dias_24	5769	705	8.179	0.863	1.676	0.225	-0.194	0.121
251.52+2.00	Dias_25	4626	286	7.885	0.416	1.356	0.096	-0.262	0.056
273.89-1.58	Dias_26	5246	1649	8.519	0.456	2.777	0.504	-0.149	0.231
292.34-4.88	Dias_27	6346	878	9.476	0.228	2.072	0.201	-0.088	0.066
293.34-0.86	Dias_28	3489	696	6.969	0.444	2.889	0.139	0.083	0.079
318.06-0.46	Dias_29	2457	373	8.070	0.754	4.064	0.111	0.156	0.031
354.99+5.17	Dias_30	10691	3453	10.022	0.216	2.056	0.506	0.195	0.061
16.28-2.81	Dias_31	3966	788	9.052	0.216	2.838	0.177	0.217	0.078
20.29-1.14	Dias_32	4783	240	8.823	0.055	4.050	0.101	0.332	0.062
22.76-0.25	Dias_33	3537	455	8.667	0.569	3.956	0.121	0.342	0.106
32.74-0.08	Dias_34	3077	655	8.414	0.259	4.039	0.101	0.409	0.131
26.61-3.98	Dias_35	5539	1074	9.577	0.049	1.425	0.063	0.219	0.074
54.15+0.17	Dias_36	3893	555	7.007	0.689	3.633	0.106	0.102	0.069
60.57-0.19	Dias_37	2903	774	8.951	0.761	3.709	0.404	0.074	0.049
98.51+3.31	Dias_38	6826	1583	8.962	0.101	2.897	0.225	-0.178	0.111
98.51+3.31	Dias_39	5181	1027	8.611	0.621	2.782	0.313	-0.246	0.036

eliminates the need for initial assumptions about the tuning parameters of the method.

For all clusters, we determined mean proper motions and parallaxes as well as the fundamental parameters (d , $\log(t)$, A_V) and $[\text{Fe}/\text{H}]$), using our non-subjective multidimensional global optimization tool to fit theoretical isochrones to *Gaia* DR3 photometric data.

We were able to obtain the mean radial velocity for 14 open clusters using radial velocity data published in *Gaia* DR3 catalog.

Surprisingly, we found Dias_30 among the oldest and most distant in the

list of the known open clusters in the Galaxy, making it an interesting case to be investigated further with deep data such as Euclid, LSST and ELT.

Finally, the results of this work indicate that there is still work to be done on open cluster discovery. Although the need for automated searches is evident, the use of clustering algorithms has not yet been exhausted as, depending on the technique, OCs with different characteristics are detected. Even HDBSCAN deserves to be tested with different tuning parameters in the search, mainly due to its efficiency and speed.

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8. APPENDIX

Color-magnitude diagram with isochrone fitted for the discovered clusters.

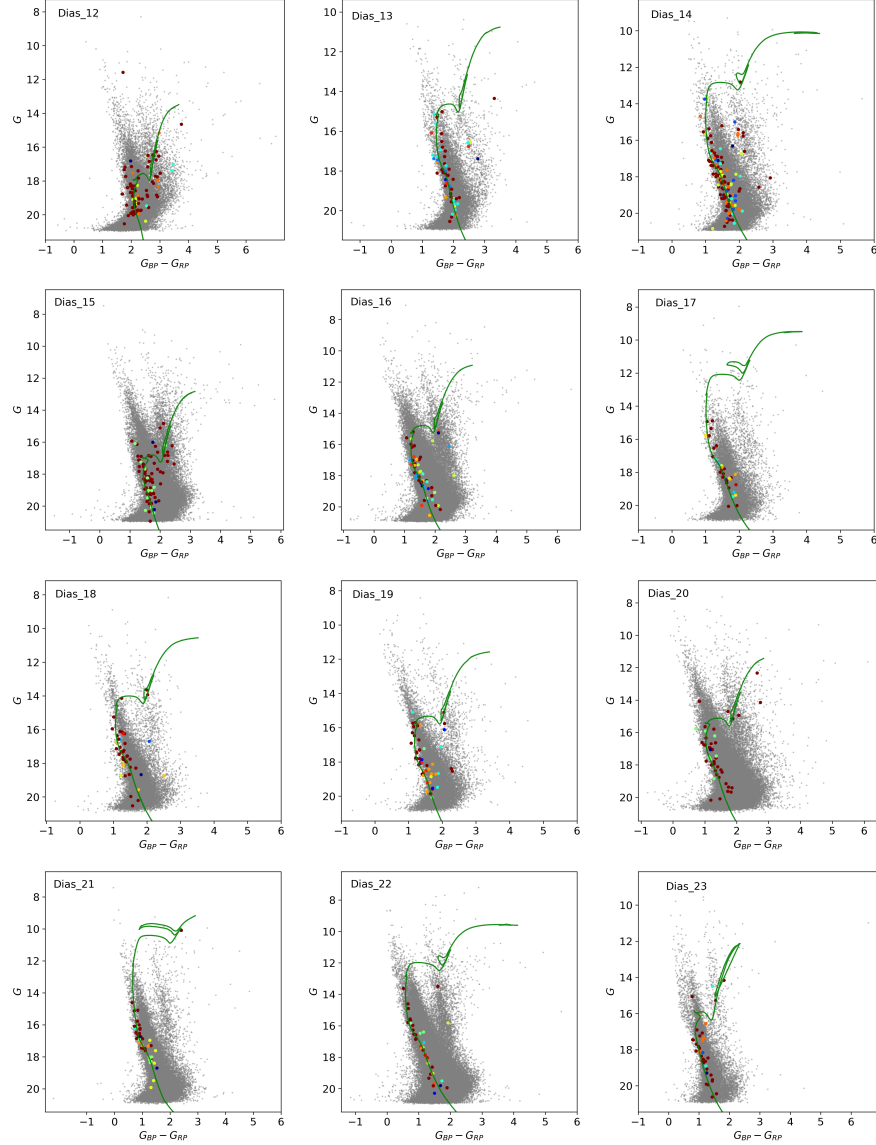


Fig. A1. Color-magnitude diagram with isochrone fitted for discovered clusters studied. The color is proportional to the astrometric membership, where more red is higher membership. The gray points correspond to field stars.

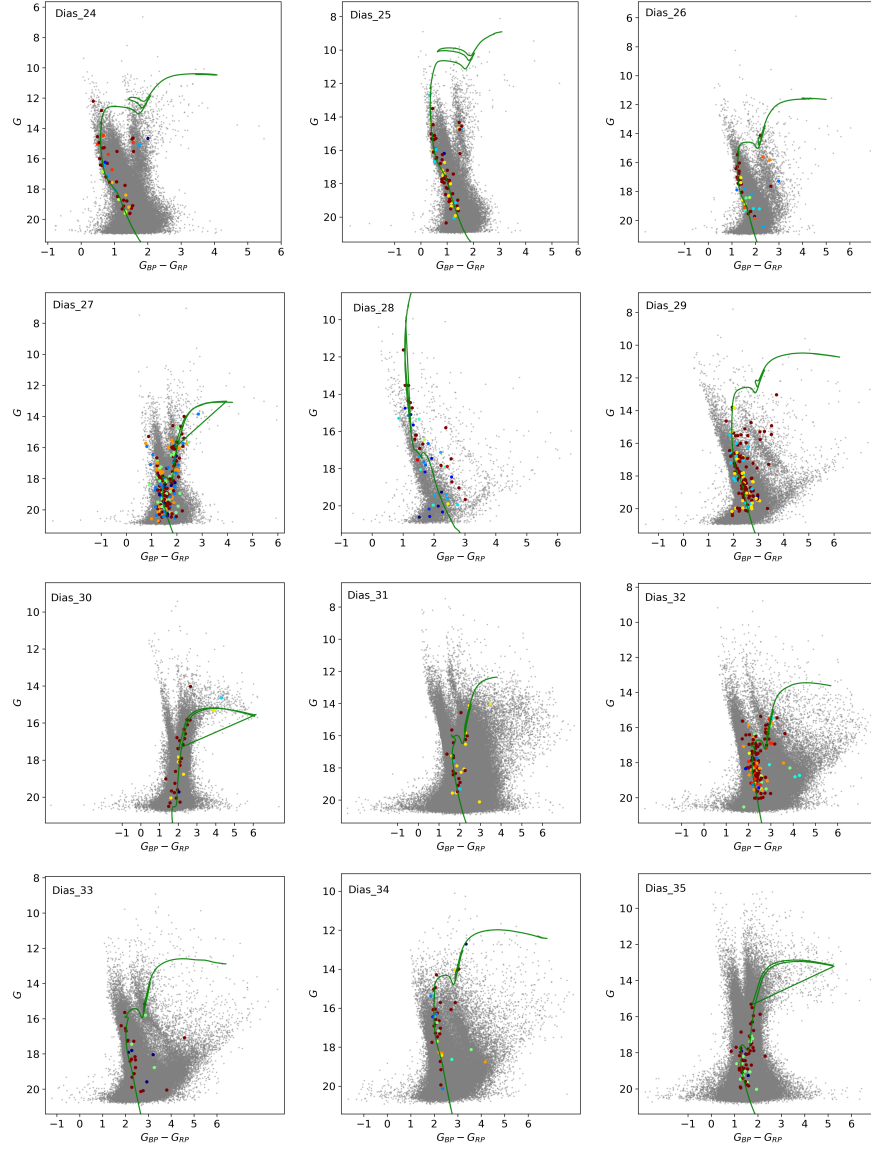


Fig. A2. Same as A1.

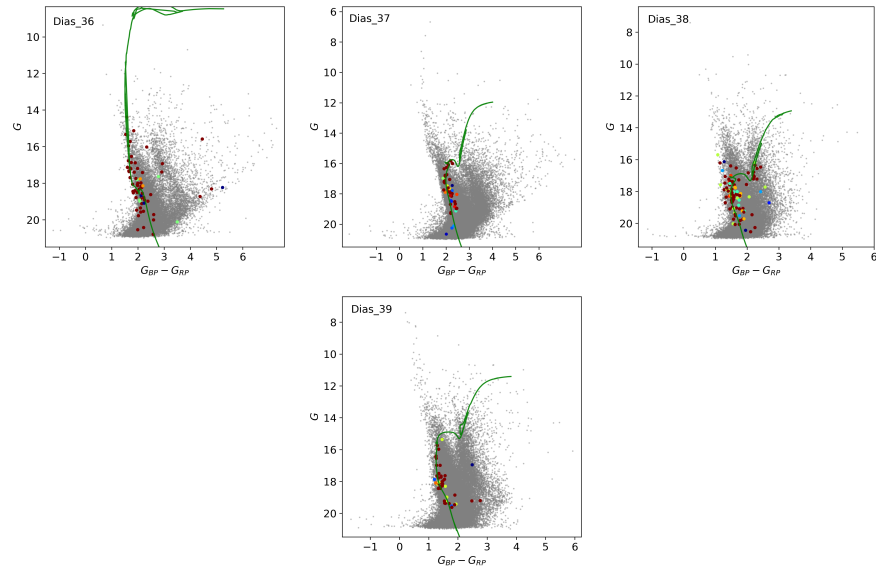


Fig. A3. Same as A1.

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