

ON THE REDUCTION OF THE CARTE DU CIEL CATALOGUES

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RESUMEN

Con miras a recuperar de la Carte du Ciel (CdC) posiciones útiles para la determinación de movimientos propios utilizando métodos de ajuste de bloques recién desarrollados, un área de la zona de Oxford de la CdC que consiste en 24 placas sobrepuestas y con época relativamente homogénea se reanalizó utilizando las medidas originales contenidas en el catálogo. Se obtuvieron posiciones para el equinoccio 1950.0 y época 1901.2 de 2911 estrellas hasta la magnitud límite del catálogo utilizando independientemente como sistema de referencia tanto el AGK3 como el FK4. Nuestros resultados demuestran que se pueden obtener posiciones precisas hasta $0''.3$ y que el método de ajuste de bloques empleado, es capaz de absorber términos de distorsión que son función lineal de las coordenadas. También se demuestra que no existen otros términos importantes de distorsión en el área escogida.

ABSTRACT

In the hope of retrieving useful positions for proper motion determinations from the Carte du Ciel (CdC) using recently developed block adjustment methods, an area from the Oxford Zone I of the CdC consisting of 24 overlapping plates of relatively homogeneous epoch was reanalyzed using the original measurements contained in the catalogue. Positions for equinox 1950.0 and epoch 1901.1 were obtained for 2911 stars down to the limiting magnitude of the catalogue using independently both the AGK3 and the FK4 catalogues as reference systems. Our results show that positions accurate to $0''.3$ may be obtained and that the block adjustment method used is capable of absorbing distortion terms which are a linear function of the coordinates. It is also shown that no other important distortion terms are present in the chosen plate material.

Key words: PROPER MOTIONS – STARS-CATALOGS

I. INTRODUCTION

The Carte du Ciel is the oldest practically complete astrophotographic coverage of the sky. The authors were fully aware of the necessity to complete coverage within as short a time span as possible. This, together with atmospheric refraction problems, led to a distribution of the work among a fairly large number of observatories. The consequence of this is a rather inhomogeneous material. Some zones, through subsequent analysis, have become known to be of surprisingly high accuracy, while others may possibly never serve their original purpose.

To make full use of the usable part of the Carte du Ciel two additional sources of data are needed: first, a modern epoch of position observation has to be produced. As has been shown by Della Prugna (1981) and by Cova (1981), not necessarily do the new observations have to be made with the same telescopes. A second and far more serious problem is the production of a reference system for the Carte du Ciel plates, unless one is satisfied with relative proper motions. The latter shall not be pursued any further in this paper.

Furthermore, a detailed analysis of the specific characteristics of the telescopes involved is required. Here we

refer to systematic errors such as field distortion (radial or asymmetric), color and magnitude dependent distortions, etc. For the telescopes responsible for the modern epoch this constitutes no major problem, since comparisons can be made with error-free data (see for instance Stock 1978). Concerning the Carte du Ciel material, use has to be made of the same old epoch plates since it cannot be assumed that the characteristics of the respective telescopes have remained unchanged for almost a century. As will be shown, the generous plate overlap provided by the Carte du Ciel plan permits an *a posteriori* determination of such errors based on the same plate material.

An extensive and detailed description of the status of the Carte du Ciel project is given by Eichhorn (1974). In some cases, for instance the Oxford zones, the original authors already give expressions which permit the conversion of the measured plane coordinates into equatorial coordinates. More recently, Günther and Kox (1972) reanalyzed one of the Oxford zones and give revised plate constants including color and magnitude terms. Also they refer the positions to the now generally accepted reference system defined by the FK4 Catalogue.

We proposed ourselves to elaborate a pilot project of

a reanalysis of a limited section of the Oxford I zone, using the original measurements contained in the catalogue and block adjustment principles. The project shall include an attempt to determine the field distortion of the Oxford astrograph. The AGK3 catalogue will be used as reference system, but it will also be attempted to use the FK4 as the only reference source.

II. THE TEST FIELD

Within a limited portion of the Oxford I catalogue presently available to us we selected as large a block of overlapping plates as possible, imposing the condition of a uniform epoch. The best choice that could be made is a field of 24 plates. Their numbers, epochs, and relative location as shown schematically in Figure 1.

The next step was to identify stars common to two or more plates. For this purpose the plate constants given in the catalogue were used, converting the original X and Y values into equatorial coordinates on a uniform system. These coordinates then were used to establish identities. Figure 2 shows the number of links which exist between the different plates.

III. COORDINATE TRANSFORMATION

Following a procedure proposed by Stock (1981) the measured coordinates X and Y are transformed to cartesian space coordinates u, v, w by

$$\tan r = \frac{(x^2 + y^2)^{1/2}}{F} \quad (1)$$

$$u = x \cos r, \quad (2)$$

$$v = y \cos r, \quad (3)$$

and

$$w = F \cos r. \quad (4)$$

Here it is assumed that tangential projection is applicable to the specific case. It is assumed that the origin of

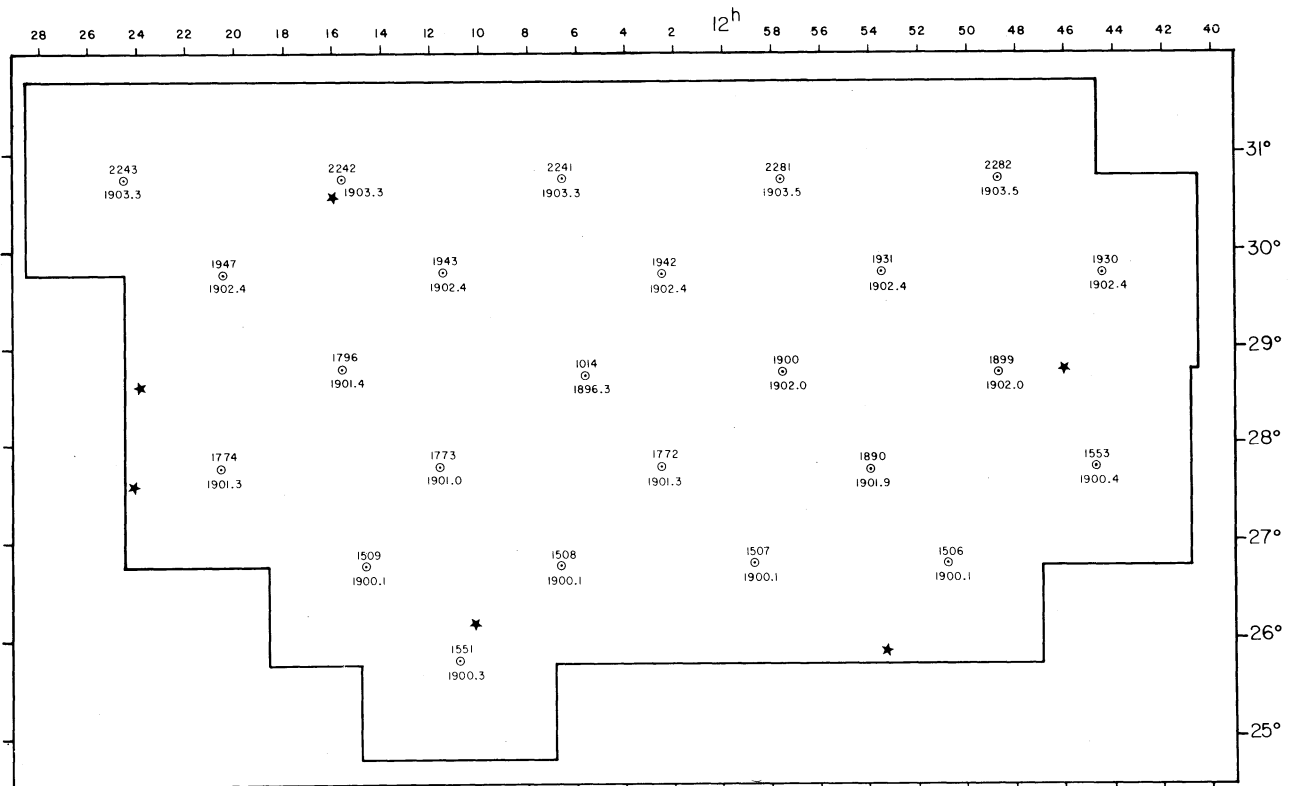


Fig. 1. Relative location, number, and epoch of the Oxford plates used in this paper. Asterisks indicate the location of reference stars contained in the FK4 supplements.

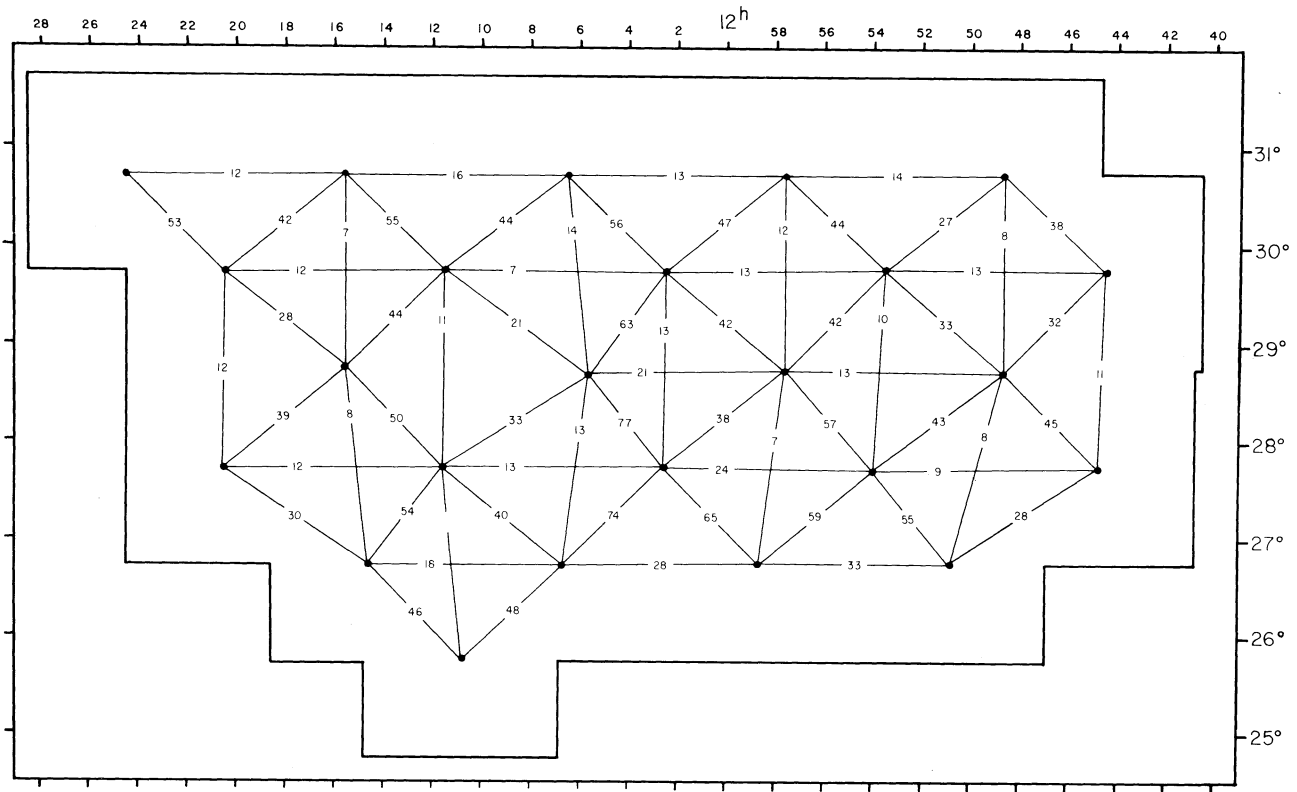


Fig. 2. Schematic representation of stars common to neighboring plates.

the X and Y coordinates coincides with the tangential point. These coordinates are transformed into standard three-dimensional coordinates ξ, η, ζ by

$$\xi = a_{11} u + a_{12} v + a_{13} w \quad (5)$$

$$\eta = a_{21} u + a_{22} v + a_{23} w \quad (6)$$

$$\zeta = a_{31} u + a_{32} v + a_{33} w \quad (7)$$

The latter are related to the equatorial coordinates α and δ by

$$\xi = F \sin \alpha \cos \delta \quad (8)$$

$$\eta = F \cos \alpha \cos \delta \quad (9)$$

$$\zeta = F \sin \delta \quad (10)$$

The matrix a_{ij} in equations (5) to (7) is an orthogonal rotation matrix.¹¹ The values of the coefficients are determined from stars for which X and Y as well as α and δ are known. If a plate by plate reduction is carried out, a minimum of three reference stars per plate is required. When overlapping plates covering a larger field are available, block adjustment principles may be used. In such a case, a minimum of three reference stars in the entire field is required. Least square methods are applied when the number of equations exceeds the number of unknowns. The equations to be used are

$$a_{m11} u_{mi} + a_{m12} v_{mi} + a_{m13} w_{mi} = \xi_i + \epsilon_{\xi i} \quad (11)$$

$$a_{m21} u_{mi} + a_{m22} v_{mi} + a_{m23} w_{mi} = \eta_i + \epsilon_{\eta i} \quad (12)$$

$$a_{m31} u_{mi} + a_{m32} v_{mi} + a_{m33} w_{mi} = \zeta_i + \epsilon_{\zeta i} \quad (13)$$

in the case of a reference star with the number i measured on the plate with the number m , the values ξ_i, η_i, ζ_i are calculated from equations (8) to (10). The coefficients

a_{mij} are the rotation matrix corresponding to plate m . The errors ϵ_{ξ_i} , ϵ_{η_i} , ϵ_{ζ_i} are residuals accumulated by errors in the measured coordinates, the coefficients, and the catalogue positions.

For an object i measured on plates m and n the equations

$$a_{m11} u_{mi} + a_{m12} v_{mi} + a_{m13} w_{mi} - a_{n11} u_{ni} - a_{n12} v_{ni} - a_{n13} w_{ni} = \epsilon_{\xi_i} \quad (14)$$

$$a_{m21} u_{mi} + a_{m22} v_{mi} + a_{m23} w_{mi} - a_{n21} u_{ni} - a_{n22} v_{ni} - a_{n23} w_{ni} = \epsilon_{\eta_i} \quad (15)$$

$$a_{m31} u_{mi} + a_{m32} v_{mi} + a_{m33} w_{mi} - a_{n31} u_{ni} - a_{n32} v_{ni} - a_{n33} w_{ni} = \epsilon_{\zeta_i} \quad (16)$$

may be used.

Evidently, the distribution of the residuals calculated in the equations (11) to (16) is not necessarily Gaussian, even if the error distribution of the measured data or the catalogue positions are of a Gaussian type. Even so, experience shows that very good approximations for the coefficients a_{mij} can be obtained by imposing

$$\Sigma \epsilon_{\xi_i}^2 = \text{Minimum} \quad ,$$

$$\Sigma \epsilon_{\eta_i}^2 = \text{Minimum} \quad ,$$

$$\Sigma \epsilon_{\zeta_i}^2 = \text{Minimum} \quad .$$

In any case the procedure which we have used leads to an algorithm which is very easy to handle.

The sums of the residuals must include all objects and all plates for which either set of the above equations is applicable. For each object weight factors may be applied to either equations (11) to (13) or equations (14) to (16).

IV. FIELD DISTORTION

When the measured plate coordinates X and Y contain distortion components, then equations (1) to (3) cannot be applied directly. The coordinates have to be corrected first to an undistorted system. However, as was mentioned by Stock (1981), if orthogonality is not imposed on the matrices, then the system worked out in Section III is capable of absorbing distortion terms which are linear functions of the coordinates. This means that, excepting large zenith distances, no correction for differential refraction is required. The same is true for other minor effects, such

as for example the aberration of light due to the earth's motion. The principal effect we are concerned with here is due to the optical system. We can distinguish several types of distortions. Some are related to the location of the images, others to their structure. We can also distinguish distortion of radial symmetry which is independent of the position angle on the plate from unsymmetric types which show a position-angle dependence.

Considering the type of optics of the Oxford refractor we may at first instance expect what is usually termed as a "pin cushion" or "barrel" deformation of the field. We shall at this point neglect effects due to either coma or to chromatic aberrations. Thus, the displacement ΔR of an image in the direction of the radius R (distance from tangential point) takes the form

$$\Delta R(R) = \sum_{i=1}^N b_i R^{2i} \quad (17)$$

with

$$R = (X^2 + Y^2)^{1/2} \quad (18)$$

For reasons of symmetry only even powers of R are used. The extent of the expansion, i.e., the value of N , has to be determined empirically.

Let us consider the case of two partially overlapping plates, with the measured coordinates X_1, Y_1 and X_2, Y_2 respectively. If no distortion is present, the three dimensional coordinates ξ_1, η_1, ζ_1 , and ξ_2, η_2, ζ_2 can be calculated from equations (1) to (4). Between these two sets of coordinates we expect the relation

$$\xi_1 = a_{11} \xi_2 + a_{12} \eta_2 + a_{13} \zeta_2 \quad , \quad (19)$$

$$\eta_1 = a_{21} \xi_2 + a_{22} \eta_2 + a_{23} \zeta_2 \quad , \quad (20)$$

$$\zeta_1 = a_{31} \xi_2 + a_{32} \eta_2 + a_{33} \zeta_2 \quad , \quad (21)$$

Again the matrix a_{ij} is an orthogonal matrix. This time, however, its diagonal elements will be near unity, and the remaining ones near zero.

For the field size under consideration $\xi \sim X$ and $\eta \sim Y$, such that we may use

$$\rho = (\xi^2 + \eta^2)^{1/2} \quad (22)$$

and

$$\Delta \rho(\rho) = \sum_{i=1}^N b_i \rho^{2i} \quad (23)$$

instead of equations (17) and (18).

The components of the distortion in the ξ - and η - direction are

$$\Delta\xi = \frac{\xi}{\rho} \Delta\rho \quad (24)$$

and

$$\Delta\eta = \frac{\eta}{\rho} \Delta\rho \quad (25)$$

These terms have to be introduced into equations (19) and (20). Considering that these correction terms may be expected to be very small, we shall neglect their products with the matrix elements which are small, and find

$$\begin{aligned} \xi_1 \left(1 + \frac{\Delta\rho_1}{\rho_1}\right) &= a_{11} \xi_2 \left(1 + \frac{\Delta\rho_2}{\rho_2}\right) \\ &+ a_{12} \eta_2 + a_{13} \zeta_2 \quad , \end{aligned} \quad (26)$$

$$\begin{aligned} \eta_1 \left(1 + \frac{\Delta\rho_1}{\rho_1}\right) &= a_{21} \xi_2 + a_{22} \eta_2 \\ \left(1 + \frac{\Delta\rho_2}{\rho_2}\right) &+ a_{23} \zeta_2 \quad . \end{aligned} \quad (27)$$

Each of these two equations contains $3 + N$ unknowns, of which the b_k -terms are common to both. For a least-square determination of the latter from plate pairs with more than $3 + N$ stars in common we can either solve the ξ - and η terms separately and check whether both lead to the same b_k coefficients, or we can combine them into one single set of $6 + N$ equations with an equal number of unknowns, thus imposing equality of the b_k -terms.

For the empirical determination of the b_k -coefficients we selected all plate pairs which had forty or more stars in common, a total of 24 pairs. The effectiveness of the method outlined above was demonstrated accidentally using these plate pairs. Due to an error in the computer program an improper projection geometry was applied to the plates. This error was detected on the basis of the large and consistent coefficients which the test produced, and which corresponded exactly to the difference between the applied and the correct projection. After correcting the error the coefficients became of an erratic nature, amounting to insignificant values if averaged over the 24 plate pairs. The matrix coefficients, on the other hand, deviate significantly and consistently from an orthogonal matrix. This means that the type of field

distortion present in the Oxford Refractor, due to the optics, telescope alignment, atmospheric refraction, screw errors of the measuring machine, etc., are effectively absorbed by the matrix if orthogonality is not imposed.

V. REDUCTION WITH THE AGK3 SYSTEM

A total of 335 stars were found to be in common with the AGK3 catalogue, with an average of 25 stars per plate. For each catalogue star the position was calculated for the epoch of 1901.2, the latter being the average epoch of the 24 Oxford plates. The block reduction was then carried out using unit weight for all links and for all reference objects.

From the position obtained for stars occurring on two or more plates their mean errors can be calculated. The mean rms errors ϵ_α and ϵ_δ for a single image are

$$\epsilon_\alpha = 0.9031 \quad \text{and} \quad \epsilon_\delta = 0''.22.$$

These correspond to just a little more than half of the last digit of the coordinates in the original Oxford Catalogue. This does not leave much room for color- or magnitude-dependent systematic errors. The latter type of error cannot be absorbed by the matrix elements.

VI. REDUCTION WITH THE FK4 CATALOGUE

Six stars were found in common with the supplements of the FK4 (1963), none with the main catalogue. As may be appreciated in Figure 2, these stars are not "strategically" located in the field covered by the plates under consideration. Again several solutions were calculated with weights for the reference stars ranging from 1.0 to 100.0. Only minor differences were found between the different solutions. However, they produce rather large systematic differences with the AGK3 in the regions away from the reference stars. These differences amount to several seconds of arc at one edge of the field. Evidently more reference stars, or at least a more favorable distribution is needed. We plan to add plates from neighbouring Carte du Ciel zones in order to increase the number of fundamental reference stars.

VII. THE PROPER MOTIONS

Proper motions were calculated by comparing the positions of the AGK2 catalogue with those obtained from the Oxford data. For the latter, naturally, we used the solution based on the AGK3 as reference source. A mean epoch difference of 28.9 years was adopted. Comparisons of these proper motions, i.e., AGK2-Oxford, with the proper motions AGK3-AGK2 is shown in Figure 3. Only stars appearing on two or more Oxford plates were used for the graphs. In this form practically all stars along the edge of the field were eliminated. The remaining stars were plotted in Figure 4. The notorious increase in the

scatter is attributed to two sources, namely, (1) the Oxford position has less weight, and (2) most of the stars are located outside of the area where the block adjustment is effective. Thus, positions and proper motions, as given in Table 1 and Table 2, are of full weight only when the position is located within the limits.

$$11^{\text{h}} 45^{\text{m}} < \alpha < 12^{\text{h}} 20^{\text{m}}, \quad 26^{\circ} 45' < \delta < 30^{\circ} 45'$$

VIII. THE POSITION CATALOGUE

The positions given in Table 1 are for a mean epoch of 1901.2, for the equinox 1950.0. The columns contain the following data:

- column 1: Running number
- columns 2-4: Right ascension
- column 5: Mean error of right ascension

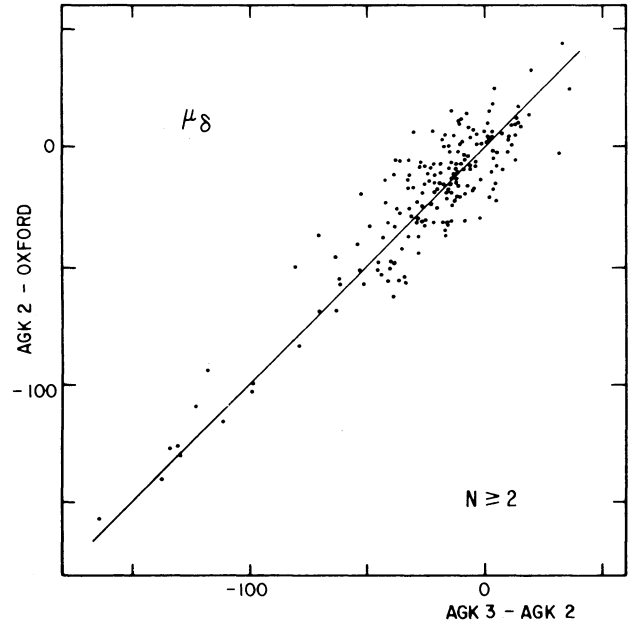
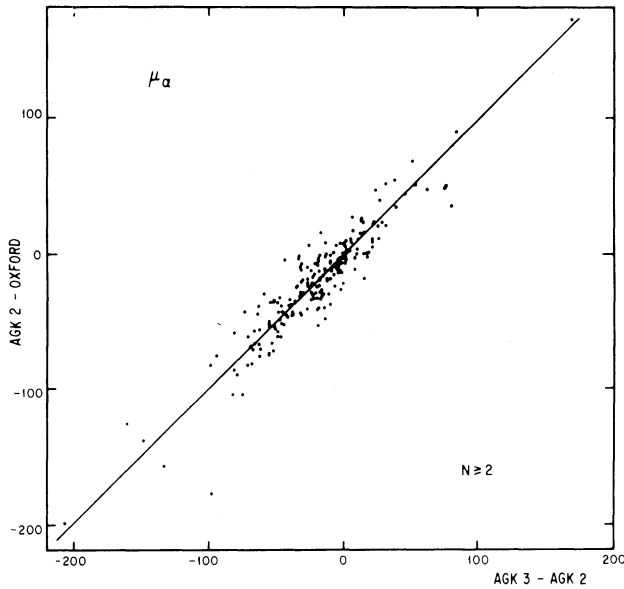


Fig. 3. Comparisons of proper motions AGK3-AGK2 versus AGK2-Oxford for stars which occur on two or more Oxford plates.

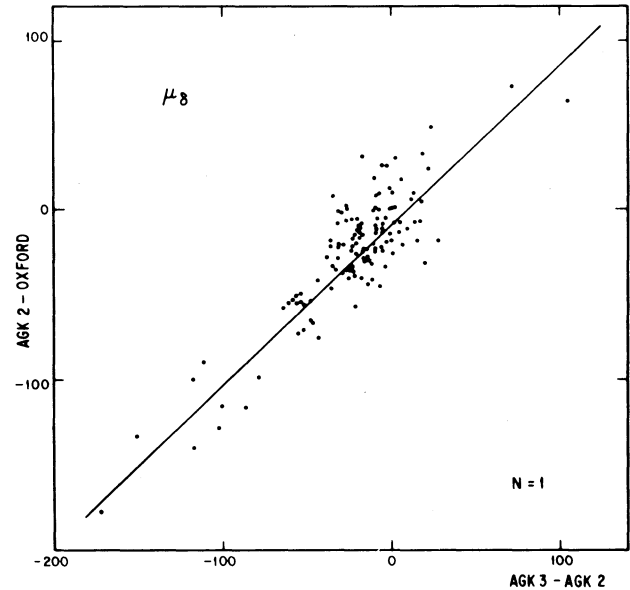
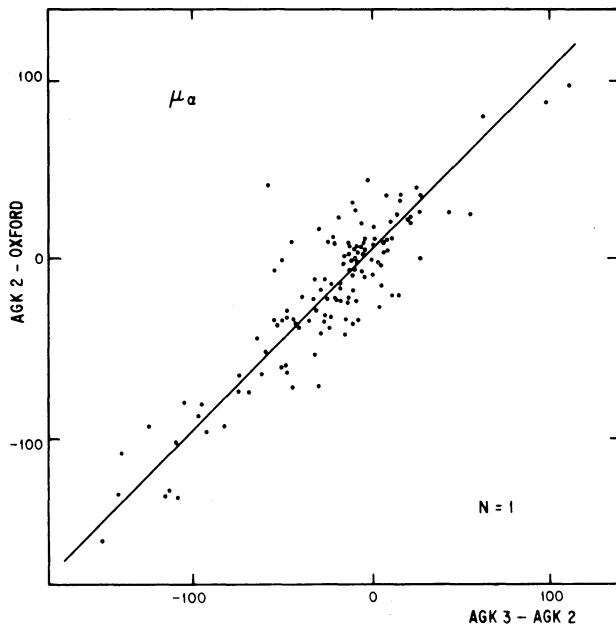


Fig. 4. Comparisons of proper motions AGK3-AGK2 versus AGK2-Oxford for stars which occur on only one Oxford plate.

- columns 6- 8: Declination
- column 9: Mean error of declination
- column 10: Magnitude as given in the Oxford Catalogue
- column 11: Number of Oxford plates on which the stars occurs
- column 12: 1 indicates that the star appears in the AGK3; 2 indicates that it appears in the FK4-Supplement.

IX. THE CATALOGUE OF PROPER MOTIONS

The proper motions AGK2–Oxford and AGK3–AGK2 are given in Table 2. The columns contain:

- column 1: Running number of Table 1
- column 2: Right ascension
- column 3: Declination
- column 4: Magnitude as given in the Oxford Catalogue
- column 5: Proper motion in right ascension in 0".001/year from AGK2–Oxford
- column 6: Proper motion in declination in 0".001/year from AGK2–Oxford
- column 7: Proper motion in right ascension in 0".001/year from AGK3–AGK2
- column 8: Proper motion in declination in 0".001/year from AGK3–AGK2.

X. THE CROSS REFERENCE CATALOGUE

For future reference and for other users we give in Table 3 the running number of Table 1 and the corresponding Oxford plate numbers and star numbers. It is important to note that we have used only the last four digits of the star numbers given in the Oxford catalogue.

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TABLE I (CONTINUED)

Table with columns: No, alpha, epsilon_alpha, delta, epsilon_delta, mp, N, A, No, alpha, epsilon_alpha, delta, epsilon_delta, mp, N, A, No, alpha, epsilon_alpha, delta, epsilon_delta, mp, N, A. The table contains multiple rows of data representing various astronomical observations.

REDUCTION OF CDC CATALOGUES

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REDUCTION OF CDC CATALOGUES

Table with multiple columns containing numerical data and identifiers, organized in a grid-like structure for catalog entries.

TABLE 2

THE PROPER MOTION CATALOGUE

Table with columns for N0, alpha, delta, mp, AGK2, AGK3, and their respective offsets. It lists star data in two columns, with 804 rows on the left and 807 rows on the right.

TABLE 2 (CONTINUED)

Table with columns for N0, alpha, delta, mp, AGK2, AGK3, N0, alpha, delta, mp, AGK2, AGK3. The table contains multiple columns of numerical data for various entries, organized in pairs of columns.

TABLE 3
THE CROSS REFERENCE CATALOGUE

Table with 14 columns: No, Oxford, No, Oxford, No, Oxford, No, Oxford, No, Oxford, No, Oxford, No, Oxford. It lists cross-references between different astronomical catalogs.

REDUCTION OF CDC CATALOGUES

TABLE 3 (CONTINUED)

Table with 8 columns: No, Oxford, No, Oxford, No, Oxford, No, Oxford, No, Oxford, No, Oxford, No, Oxford. Rows contain numerical data representing reduced CDC catalogues.

TABLE 3 (CONTINUED)

| No | Oxford | No | Oxford | No | Oxford | No | Oxford | No | Oxford | No | Oxford | No | Oxford |
|------|-----------|------|-----------|------|-----------|------|-----------|------|-----------|------|-----------|------|-----------|
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| 2578 | 2243 8411 | 2625 | 1774 4524 | 2673 | 1947 8174 | 2718 | 1947 8200 | | 2243 8363 | 2808 | 1947 8169 | 2857 | 2243 8322 |
| 2579 | 2243 8418 | 2626 | 1947 8215 | 2674 | 1774 4642 | | 2243 8319 | 2761 | 2243 8384 | 2809 | 1947 8260 | 2858 | 2243 8420 |
| 2580 | 1947 8193 | | 2243 8334 | 2675 | 1947 8113 | 2719 | 1947 8254 | 2762 | 1947 8183 | | 2243 8386 | 2859 | 2243 8364 |
| | 2243 8309 | 2627 | 1947 8239 | | 1774 4769 | | 2243 8383 | 2763 | 2243 8488 | 2810 | 1774 4669 | 2860 | 2243 8365 |
| 2581 | 1947 8106 | | 2243 8369 | 2676 | 2243 8464 | 2720 | 1774 4551 | 2764 | 2243 8498 | 2811 | 2243 8467 | 2861 | 2243 8328 |
| 2582 | 1774 4762 | 2628 | 1774 4525 | 2677 | 1947 8266 | 2721 | 2243 8426 | 2765 | 1947 8121 | 2812 | 1947 8188 | 2862 | 2243 8417 |
| 2583 | 1774 4647 | 2629 | 1947 8185 | | 2243 8488 | 2722 | 1774 4664 | 2766 | 2243 8385 | | 2243 8303 | 2863 | 2243 8377 |
| 2584 | 1947 8272 | | 2243 8310 | 2678 | 1774 4583 | 2723 | 1947 8224 | 2767 | 1774 4594 | 2813 | 1774 4689 | 2864 | 2243 8356 |
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| 2585 | 1774 4586 | 2631 | 1774 4561 | 2680 | 2243 8432 | 2724 | 1774 4736 | | 2243 8301 | 2815 | 1774 4530 | 2866 | 2243 8357 |
| 2586 | 2243 8425 | 2632 | 1774 4742 | 2681 | 1774 4527 | 2725 | 1774 4509 | 2769 | 1774 4592 | 2816 | 2243 8479 | 2867 | 2243 8391 |
| 2587 | 1947 8237 | 2633 | 1774 4686 | 2682 | 1774 4598 | 2726 | 2243 8458 | 2770 | 2243 8419 | 2817 | 1947 8220 | 2868 | 2243 8315 |
| | 2243 8368 | 2634 | 1774 4575 | 2683 | 1774 4577 | 2727 | 1947 8167 | 2771 | 1774 4614 | | 2243 8340 | 2869 | 2243 8305 |
| 2588 | 1774 4724 | 2635 | 1774 4662 | 2684 | 1947 8145 | 2728 | 2243 8487 | 2772 | 1947 8115 | 2818 | 2243 8469 | 2870 | 2243 8428 |
| 2589 | 1774 4756 | 2636 | 1774 4663 | 2685 | 1947 8114 | 2729 | 1947 8217 | | 1774 4771 | 2819 | 1947 8132 | 2871 | 2243 8399 |
| 2590 | 1774 4523 | 2637 | 1774 4540 | | 1774 4770 | | 2243 8336 | 2773 | 1774 4579 | 2820 | 1947 8206 | 2872 | 2243 8373 |
| 2591 | 1774 4640 | 2638 | 1947 8223 | 2686 | 1947 8212 | 2730 | 1947 8255 | 2774 | 1774 4679 | | 2243 8327 | 2873 | 2243 8304 |
| 2592 | 1947 8163 | | 2243 8354 | | 2243 8335 | | 2243 8375 | 2775 | 1774 4542 | 2821 | 1947 8261 | 2874 | 2243 8306 |
| 2593 | 1947 8154 | 2639 | 2243 8438 | 2687 | 1774 4687 | 2731 | 1774 4655 | 2776 | 2243 8478 | | 2243 8387 | 2875 | 2243 8333 |
| 2594 | 2243 8326 | 2640 | 1774 4649 | 2688 | 1774 4546 | 2732 | 1947 8130 | 2777 | 1774 4666 | 2822 | 2243 8341 | 2876 | 2243 8349 |
| 2595 | 1947 8134 | 2641 | 1774 4650 | 2689 | 1947 8253 | 2733 | 1774 4744 | 2778 | 1774 4593 | 2823 | 2243 8342 | 2877 | 2243 8448 |
| 2596 | 1947 8135 | 2642 | 1774 4541 | | 2243 8381 | 2734 | 1774 4643 | 2779 | 1774 4580 | 2824 | 2243 8480 | 2878 | 2243 8461 |
| 2597 | 1947 8208 | 2643 | 1947 8137 | 2690 | 1947 8175 | 2735 | 1947 8210 | 2780 | 1947 8116 | 2825 | 2243 8372 | 2879 | 2243 8329 |
| | 2243 8330 | 2644 | 1774 4651 | 2691 | 2243 8414 | | 2243 8332 | | 1774 4772 | 2826 | 2243 8388 | 2880 | 2243 8482 |
| 2598 | 2243 8473 | 2645 | 2243 8475 | 2692 | 1947 8155 | 2736 | 1774 4728 | 2781 | 1947 8157 | 2827 | 2243 8398 | 2881 | 2243 8358 |
| 2599 | 1947 8238 | 2646 | 1947 8156 | | 2243 8382 | 2737 | 1774 4745 | 2782 | 2243 8439 | 2828 | 2243 8489 | 2882 | 2243 8441 |
| | 2243 8362 | 2647 | 1947 8164 | 2693 | 2243 8477 | 2738 | 1774 4746 | 2783 | 1774 4729 | 2829 | 2243 8459 | 2883 | 2243 8462 |
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| 2601 | 1947 8273 | 2649 | 1774 4576 | 2695 | 1774 4711 | 2740 | 1774 4566 | 2785 | 2243 8466 | 2831 | 2243 8440 | 2885 | 2243 8436 |
| | 2243 8402 | 2650 | 1947 8125 | 2696 | 1774 4678 | 2741 | 2243 8337 | 2786 | 1947 8122 | 2832 | 2243 8347 | 2886 | 2243 8374 |
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| 2604 | 1774 4539 | 2653 | 1774 4641 | 2698 | 2243 8446 | 2743 | 1947 8131 | 2789 | 1774 4607 | 2835 | 2243 8355 | 2889 | 2243 8491 |
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| 2613 | 1774 4612 | 2663 | 1774 4603 | 2708 | 1774 4735 | 2752 | 2243 8497 | 2797 | 1947 8177 | 2845 | 2243 8499 | 2899 | 2243 8378 |
| 2614 | 2243 8454 | 2664 | 1947 8182 | 2709 | 1947 8144 | 2753 | 1947 8218 | 2798 | 1947 8178 | 2846 | 2243 8389 | 2900 | 2243 8450 |
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| 2619 | 1774 4630 | 2668 | 2243 8476 | 2714 | 1774 4617 | 2757 | 1947 8225 | 2803 | 1947 8267 | 2851 | 2243 8304 | 2905 | 2243 8379 |
| 2620 | 2243 8455 | 2669 | 1774 4702 | 2715 | 1947 8246 | | 2243 8345 | | 2243 8397 | 2852 | 2243 8490 | 2906 | 2243 8307 |
| 2621 | 1774 4574 | 2670 | 1774 4526 | | 2243 8371 | 2758 | 1947 8194 | 2804 | 1774 4737 | 2853 | 2243 8376 | 2907 | 2243 8451 |
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| 2623 | 2243 8486 | | 2243 8331 | | 2243 8318 | 2759 | 1774 4591 | 2806 | 1947 8184 | 2855 | 2243 8313 | 2909 | 2243 8359 |
| | | | | | | | | | | | | 2910 | 2243 8422 |
| | | | | | | | | | | | | 2911 | 2243 8408 |

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