

## ABSTRACTS OF CONTRIBUTED PAPERS

## BOUNDARY LAYERS AROUND T TAURI STARS

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In the standard model of accretion disks there is a thin transition region between the Keplerian disk and the slowly rotating pre-main sequence star, the boundary layer, which is supposed to be responsible for the UV and optical excess in some T Tauri spectra.

We have modeled the boundary layer, assuming hydrostatic equilibrium in both the vertical and radial directions, and thermal equilibrium between turbulent viscous heating ( $\alpha$  prescription) and radiative cooling. Our input parameters are:  $M_*$ ,  $L_*$ ,  $T_*$ ,  $\alpha$ , and  $M$  we have solved an implicit equation for the temperature, and with this quantity we have computed the size of the region, the density and the optical depth.

With temperature and density, and introducing the inclination angle as a new parameter, we have calculated the boundary layer spectrum and compared our predicted colors  $U - B$ ,  $B - V$  and fluxes around the Balmer Jump, with published observational data from Rydgren et al. (1984, Pub.US Naval Obs., Vol 25, part 1) and Hartigan et al. (1991, ApJ, 82, 617). We take stellar spectra from Bruzual & Charlot (1993, ApJ, 405, 538), and opacities from Calvet et al. (1991, ApJ, 380, 617).

The predicted spectra and colors reproduce the observed features in UV and optical wavelengths.

ABOUT ONGOING LOW MASS  
STAR FORMATION IN NGC 6611E. de Lara, C. Chavarría-K., and  
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We present preliminary results of a CCD imagery study ( $RIH\alpha$ ) of selected parts of the stellar/nebular complex NGC 6611/M16. Our data show low mass star formation on the bright rims of this region, particularly those associated with the central part of M16 known as the "Queen Star".

ROSAT X-RAY SOURCES IN ORION  
STAR FORMING REGIONJ.M. Alcalá<sup>1</sup>, J. Krautter<sup>1</sup>, L. Terranegra<sup>2</sup>,  
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In the course of a systematic study of *ROSAT* X-ray sources in star forming regions (SFR) we have started spectroscopic observations in order to identify the optical counterparts of the *ROSAT* X-ray sources. The main purpose of this study is the search for weak line T Tauri stars (WTTS). WTTS are low-mass PMS stars which are found in SFR and which display the Li I  $\lambda 6707$  absorption line (which verify their PMS nature), but lack both strong  $H\alpha$  emission and strong IR excesses characteristic of classical T Tauri stars (CTTS). WTTS are believed to be PMS stars which lack dense surrounding matter, and therefore to be more easily detected by their strong solar-like coro-

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nal X-ray emission. Our spectroscopic observations revealed that, in fact, a significant fraction of the identified optical counterparts are WTTS.

Here we discuss a sub-sample of the new WTTS in Orion for which we have both spectroscopic and photometric data. We also discuss the spatial distribution and the evolutionary status of these stars, and present an analysis of the relationship between X-ray emission and the physical properties ( $M_*$ ,  $R_*$ , and  $L_*$ ) of these WTTS.

#### DUST IN PRE-MS SYSTEMS: GRAIN GROWTH?

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Dust grains associated with pre-main sequence low-mass (T Tauri) stars and intermediate-mass (Herbig Ae/Be) stars have been observed in submm continuum. Twin aims are to estimate total amounts of circumstellar material and to seek evidence for grain growth. Spectral fits to optically thin submm emission from DG Tauri, Haro 6-13 and DO Tauri suggest that circumstellar grain opacities ( $\kappa \propto \lambda^{-0.6}$ ) fall off much more slowly with increasing wavelength than is found for grains in the interstellar medium, a result also indicated by preliminary analysis of submm photometry of the Herbig Ae system HD 163296. Such low opacity indices might be indicative of the accumulation of grains. Absolute values of both the grain opacity and the total mass of material cannot be separated out using model fits alone; however, such fits can yield useful constraints on the appropriate opacity scaling when considered in combination with mass accretion rates inferred from integrated bolometric luminosities. In this way, the submm photometry of the T Tauri systems provides a lower limit to the mm/submm opacity of  $\kappa \geq 0.003 [1100/\lambda(\mu\text{m})]^{0.6} [\text{cm}^2 \text{g}^{-1}]$ , with corresponding disk masses being up to a few tenths of a solar mass. A rough estimate of the 2 mm opacity of the outer regions of the compact continuum source in HL Tauri suggests  $\kappa(1100 \mu\text{m}) \sim 0.03 [\text{cm}^2 \text{g}^{-1}]$ . It is interesting that such an opacity scaling would, if applied directly to the disks in DG Tauri, Haro 6-13 and DO Tauri, still imply disk masses which are in excess of the minimum-mass solar nebula. If the grains in HD 163296 are also distributed primarily in the form of a disk, then the implied total mass of gas+dust is  $\sim 0.19 M_\odot$ . The shallow submm opacity law for this Herbig Ae system indicates that, as for the T Tauri circumstellar environments, it is possible that grains might be comparatively evolved.

#### RADIO CONTINUUM, AMMONIA AND WATER MASER OBSERVATIONS OF BRIGHT UNASSOCIATED IRAS POINT SOURCES

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We present matching-beam 6 and 2-cm radio continuum observations made with the Very Large Array and ammonia and water maser observations made at the Haystack Observatory of 12 *IRAS* point sources selected from the survey of Scalise et al. (1989, A&A 221, 105) of bright, unassociated *IRAS* point sources. These sources have 60 or 100  $\mu\text{m}$  flux densities in excess of  $10^3 \text{ Jy}$  and have no previous reference in any of the 37 catalogs considered for association of *IRAS* sources with known sources. Six of the twelve sources have associated radio continuum, ammonia and water maser emission and all of them show at least one of these three emissions. In all sources detected, the ammonia is warm ( $T \sim 20 \text{ K}$ ) and suggests the association of dense molecular gas with embedded heating sources. It is argued that all sources in the sample could be associated with time-variable  $\text{H}_2\text{O}$  maser emission. The radio and far-infrared data appear to indicate that these sources are star-forming regions powered by a late O or early B-type star. Several of the sources of lower luminosity ( $\sim 5 \times 10^3 L_\odot$ ) appear to have ionizing photon fluxes in excess of those expected for a ZAMS star. Possible explanations for this discrepancy are discussed.

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#### KINEMATICS OF THE GALACTIC SUPERNOVA REMNANTS RCW 86, MSH 15-56, AND MSH 11-61

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We present radial velocity fields of the optical counterparts of three supernova remnants (SNRs): RCW 86, MSH 15-56 and MSH 11-61, obtained from scanning Fabry-Perot interferometer observations at H $\alpha$ . The kinematical distances, expansion velocities, ages, energies and phases of evolution are obtained.