composition shows He and N enrichment, typical of PNe with massive progenitors. We found a nebular expansion velocity of 28 km s^{-1}. From these parameters we derived a distance of 1.0 ± 0.2 kpc and we estimate a stellar luminosity of $L/L_{\odot} \cong 20$ and a radius of $R/R_{\odot} \cong 0.01$. From the theoretical evolutionary tracks in the H-R diagram (Shaw & Kaler 1989, ApJS, 69, 495) we obtain a mass of about 1 $M_{\odot}$ and $\log g = 8.4$ for this star. These parameters correspond to one of the most evolved central stars of PNe indicating that this object is already in the white dwarf cooling sequence. Only a few objects have been reported with these characteristics, among them NGC 7293, A 21 and A 31. These rare objects are very important in the study of the link between white dwarfs and their precursors.

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LUMINOSITY FUNCTION OF THE STARS IN THE GALACTIC BULGE

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We construct the observed Luminosity Function (LF) of a sample of stars in the low obscuration region known as Baade's Window (BW). This LF is shifted by ~1.5 mag towards brighter bolometric magnitudes with respect to that published by Frogel & Whitford (1987, ApJ, 320, 199). This has been interpreted as produced by observations of slightly different stellar populations in the same area due to the way the stars have been selected. The observed LF may be separated into two Gaussian components; they are respectively associated with the bulge and disc contributions. The component which we claim represents the disc contribution in BW may be reproduced reasonably well by projection and integration of the M giants' LF in the solar neighborhood. This fact lends support to our associating the wide component of the observed LF with a disc-like stellar population. Values for the characteristic radii for the bulge (~1.1 kpc) and for the disc (~2.3 kpc) are derived.

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A CODE TO COMPUTE STELLAR MASSES ONTO THE H-R DIAGRAM

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We have written a code to determine the initial an[d] present mass of a star located on the H-R diagram The code is based on an interpolation procedure between the two adjacent evolutionary tracks where th[e] star is located according to its luminosity and temperature (M_V and color index). With just a fe[w] modifications the code may be used with any set o[f] evolutionary tracks, be they computed with or without mass loss. Considering that several evolutionar[y] tracks may pass by the same locus on the H-R diagram, the stellar mass obtained from the star's position may be rather ambiguous. Our code takes int[o] account all possible solutions.

RECENT ADVANCES IN THE DEVELOPMENT OF THE UNAM SCANNING FABRY-PEROT INTERFEROMETER (PUMA) FOR THE STUD[Y] OF INTERSTELLAR MEDIUM

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"PUMA" is an instrument for the study of the kinematics of the gas in cosmic nebulae. It is being developed for the 2-m F/7.9 Ritchey-Chretien telescope o[f] the National Astronomical Observatory (OAN-SPM) at San Pedro Mártir.

The system is an instrument consisting of a focal reducer coupled to a Queensgate scanning Fabry-Perot interferometer (SFPI). The images or interference patterns of a sky region obtained at different wavelengths will be focused by an objective o[n] an optoelectronic detector, a CCD or a Mepsicron depending on the spectral range used.

It has a set of interference filters, a calibratio[n] system and field diaphragms to isolate regions of interest. The SFPI can be moved out of the optica[l] path in order to acquire direct images, or to allow the integration of another optical elements.

"PUMA" will provide: a) Direct imagery with o[r] without interference filters. b) Interferograms at Hα (λ6563 A), [N II](λ6589 A), [O III](λ5007 A), an[d] [S II](λ6717 A).

The selection of optical materials was based o[n] a design made at Marseilles Observatory to achiev[e] an apochromatic system, covering the desired wid[e]