

THE SCALE OF REDDENING FOR CLASSICAL CEPHEID VARIABLES

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RESUMEN

Se presenta un resumen de los enrojecimientos para 68 cefeidas y se usan estos datos para compararlos con otras escalas de enrojecimiento publicadas (obtenidas mediante índices independientes del enrojecimiento, fotometría en el sistema de seis colores de Lick, fotometría en el sistema de Strömgren, en el de Walraven, en el de Washington, en el Cape-*BVI*, en el del DDO y en el de Geneva), así como mediante estudios IRSB y espectroscopía de cefeidas, tanto antiguos como modernos. Los enrojecimientos obtenidos a partir de relaciones período-color son los menos confiables, los basados en excesos de color fotométricos varían en precisión: su exactitud depende del método y de la muestra usada para calibrar. Se ilustra la exactitud y precisión de las escalas de enrojecimiento de cefeidas publicadas, y se construye un nuevo sistema de enrojecimientos estandarizados para 198 variables, con una incertidumbre de ± 0.028 en E_{B-V} , y con una precisión de menos de ± 0.01 para muchas de ellas. Los excesos de color se usan para estudiar la manera como el color intrínseco depende del período de pulsación. Los resultados contradicen las ideas prevalecientes: la dispersión del color (o de la temperatura efectiva) en la franja de inestabilidad de las cefeidas no cambia con el período de pulsación.

ABSTRACT

Field reddenings are summarized for 68 Cepheids and used for comparison with other published reddening scales: from reddening-independent indices, photometry on the Lick six-color system, Strömgren system, Walraven system, Washington system, Cape *BVI* system, DDO system, and Geneva system, IRSB studies, and Cepheid spectroscopy, old and new. Reddenings from period-color relations are least reliable; photometric color excesses vary in precision, accuracy depending on methodology and calibration sample. The accuracy and precision of published Cepheid reddening scales are revealed, leading to a new system of standardized reddenings for 198 variables, of average uncertainty ± 0.028 in E_{B-V} and precision less than ± 0.01 for many. The color excesses are used to map the dependence of intrinsic color on pulsation period, the results contradicting ideas about the period dependence of dispersion in Cepheid effective temperatures.

Key Words: ISM: dust, extinction — stars: fundamental parameters — stars: variables: Cepheids

1. INTRODUCTION

Classical Cepheids are pulsating yellow supergiants that make ideal standard candles for distance estimates because of the fairly tight relationship that exists between their luminosities and periods of pulsation — the period-luminosity relation, also referred to as the Leavitt law when expressed as absolute magnitude versus period. The general properties of Cepheids consequently make them of consid-

erable importance for the reliability of the distance scale. That, in turn, has inspired a quest, still ongoing, to delineate the Cepheid instability strip as accurately as possible (see Turner 2012b). As noted by Turner, MacLellan & Henden (2011), an empirical description of the instability strip has been difficult to establish because of complications arising from Galactic reddening and extinction; extinction alone represents an overlooked obstacle because of

the prescriptive fashion by which it is often applied (see Turner 2014). But the creation of an accurate scale of Cepheid reddenings (color excesses) represents an important first step.

Historically, the color excesses of Cepheids have been established using field reddenings derived from photometric and spectroscopic observations of their nearby early-type companions, using photometric or spectrophotometric observations of the variables in calibrated reddening-independent systems, from comparison of their cyclical spectroscopic and photometric variations with derived intrinsic colors for supergiants of comparable spectral types, and from simple period-color relations tied to a few standard Cepheids calibrated by the previous methods. Each method is fraught with difficulties arising from uncertainties in the methods of establishing intrinsic colors for luminous stars of all types, as well as dereddening them, given that few are close enough in the Galactic plane to be unaffected by the contaminating effects of nearby dust clouds.

The last technique, however, suffers from a more fundamental problem: the intrinsic width of the instability strip in effective temperature, combined with the fact that stars of identical mean radius (or pulsation period) lie diagonally across the strip over a range of mean effective temperatures, implies that stars of identical period can have quite different intrinsic colors, in conflict with the basic premise of a period-color relation. Coincidentally, many well-studied bright Cepheids have large light amplitudes, and such variables are usually found in the more constricted region of effective temperature lying close to, and slightly warmer than, the center line of the instability strip (see Turner, Abdel-Sabour Abdel-Latif & Berdnikov 2006a). Adoption of period-color relations to deredden Cepheids can therefore yield reasonable color excesses in some cases, namely Cepheids of large amplitude, although the scatter inherent to the methodology must propagate into uncertainties for individual stars.

The accuracy and precision of reddenings for early-type stars have gradually improved over the past 60 years through developments in dereddening techniques and more reliable intrinsic colors. In parallel to such advances, Kovtyukh et al. (2008) have increased the precision of spectroscopic parameters for Cepheids and yellow supergiants through use of as many temperature-sensitive and gravity-sensitive line ratios as permitted by the dimensions of their observational spectral windows. Their technique eliminates any dependence on adopted reddening

relations by matching stellar atmosphere models to Cepheid spectra over their cycles to infer effective temperatures, and hence intrinsic colors, for comparison with observed colors. Older versions of Cepheid dereddening procedures from spectra were generally of much lower precision.

An oft-cited source of Cepheid reddenings was established 25 years ago by Fernie (1990a) and used to study the Cepheid instability strip (Fernie 1990b). The intrinsic relations adopted for that compilation were calibrated using a variety of different procedures noted above (Fernie 1987), resulting in zero-point offsets with respect to other reddening calibrations of that era. Turner (1995), for example, noted the existence of systematic offsets in reddening for some Cepheids, notably long period variables.

The importance of accurate Cepheid reddenings has been pointed out previously (Turner & Burke 2002; Turner 2010, 2012b). A Cepheid's mean radius is closely related to its pulsation period according to independent studies using the Baade-Wesselink method and its variants (Laney & Stobie 1993; Gieren et al. 1998; Arellano Ferro & Rosenzweig 2000; Turner & Burke 2002; Turner et al. 2010), while its average effective temperature can be established fairly reliably from its unreddened broad band $B-V$ color (Gray 1992; Turner & Burke 2002), as also argued by Fry & Carney (1999). Combining the two values then yields a Cepheid's mean luminosity according to the standard relationship, $\langle L \rangle = 4\pi \langle R \rangle^2 \sigma \langle T_{\text{eff}} \rangle^4$, where angular brackets denote averages over the pulsation cycle. The Cepheid period-luminosity relation can therefore be constructed from basic information on the pulsation period and reddening of Cepheids in combination with observed $B-V$ colors (Turner 2012b). The distribution of Cepheids in period-color space can also be established in similar fashion (see Turner 2001, 2010) to provide an empirical description of the Cepheid instability strip.

The gradual growth in field reddenings for Galactic Cepheids (e.g., Laney & Caldwell 2007) has made it possible to make detailed comparisons of existing reddening compilations in order to examine the accuracy and precision of each. The present study was therefore initiated in order to compile all available field reddenings for Cepheids, to redo any where reanalysis was warranted, and to use the resulting system of reddenings to examine in detail a selection of existing reddening compilations, including those often cited for the extragalactic Cepheid calibration.

2. THE REDDENING SYSTEM

The reddening system established here is based upon the growing collection of space reddenings for Cepheids that has developed over the past thirty-five years, mostly as by-products of photometric studies of open clusters and associations containing Cepheids as potential calibrators for the period-luminosity relation. The collection is not complete (several studies are still in progress), but is large enough for the present purpose.

Reddenings were rederived for a selection of Cepheids studied previously, namely those of Feltz & McNamara (1976), Turner (1980a), Turner et al. (2005), and Turner (2013). The analyses made use of Strömrgren system data transformed to the UBV system as described by Turner (1990) — see, as an example, Turner (2012a), — existing data, and BV data supplemented by new CCD observations from the AAVSO (American Association of Variable Star Observers) Photometric All-Sky Survey (APASS).

Thirty years ago it was common practice to adopt intrinsic UBV colors for A-type dwarfs from compilations such as those of Johnson (1966) and FitzGerald (1970). Later studies (Collins & Smith 1985; Peacock & Smith 1987) revealed that such colors are likely to be biased by a mix of stars having different rotational velocities and projection angles. True zero-age zero-rotation main-sequence (ZAZRMS) stars of spectral types A and F have colors that differ systematically from those of Johnson (1966) and FitzGerald (1970) according to Peacock & Smith (1987), although the standardization for the latter study was mistakenly tied to A0 stars with colors affected by a zero-point offset. True ZAZRMS colors for dwarfs were therefore recompiled by smoothing the colors for OB and KM stars from Johnson (1966) and FitzGerald (1970), renormalizing the Peacock & Smith (1987) Johnson and Strömrgren system intrinsic color predictions for non-rotating AF dwarfs to A0 stars with $[(B - V)_0, (U - B)_0] = [0.00, 0.00]$, and adding the colors for unreddened Pleiades GK dwarfs (roughly solar metallicity) from the study by Turner (1977b).

The resulting UBV colors were shown previously to match the observed colors of OB and AF stars in uniformly-reddened open clusters (Turner 1996). Prior to that, they were tested and confirmed using observed colors for large samples of BAF-type stars from the *Photoelectric Catalogue* of Blanco et al. (1970). Only stars considered to have UBV photometry properly tied to the Johnson system were used for testing. The resulting intrinsic ZAZRMS relation is shown in Figures 2 & 3 of Turner (1996)

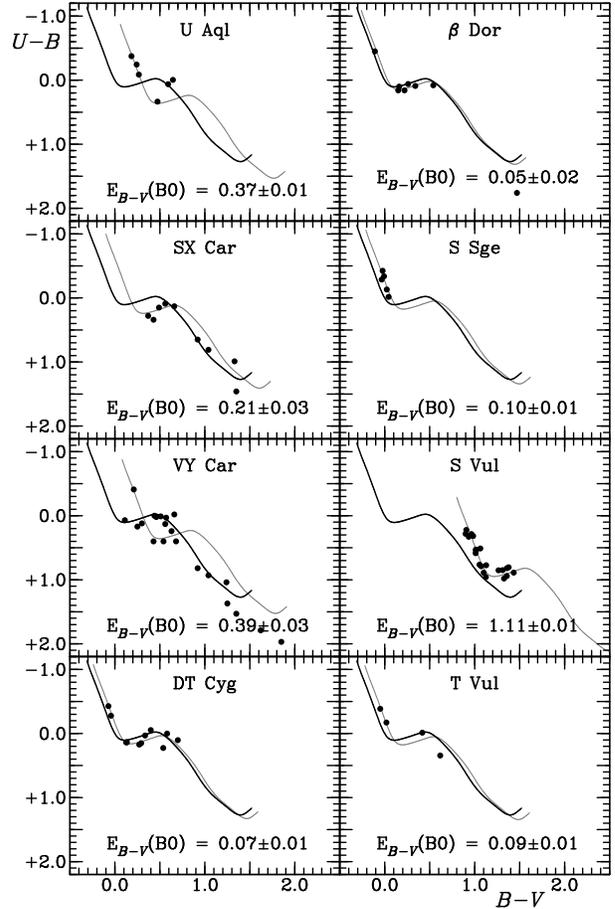


Fig. 1. Existing and transformed UBV data for stars surrounding eight Cepheids. The intrinsic ZAZRMS color-color relation is shown (in black) in each case, along with the same relation reddened by the average B0-star color excess (indicated in each case and depicted by the gray lines) for stars associated with each Cepheid.

and here in Figure 1, but for the benefit of potential users is also tabulated in Table 1. Figure 1 depicts existing and transformed UBV data used to establish new space reddenings for the eight Cepheids noted above.

An overlooked feature of interstellar reddening is that no single relationship accurately describes it over all regions of the Galaxy (Wampler 1961, 1962; Mathis 1990; Turner 1989, 1994; Zagury 2012, 2013; Turner, Majaess & Balam 2014). Regional variations in the reddening law display an inherent dependence upon direction viewed through the Galaxy, so the reddenings derived in the color-color diagrams of Figure 1 were inferred using reddening laws, *i.e.*, E_{U-B}/E_{B-V} , appropriate for each field. The corresponding color excesses for the Cepheids were also

TABLE 1
INTRINSIC UBV COLORS FOR ZERO-AGE ZERO-ROTATION DWARFS

$(B-V)_0$	$(U-B)_0$								
-0.33	-1.205	+0.04	+0.073	+0.41	-0.015	+0.78	+0.369	+1.15	+1.041
-0.32	-1.171	+0.05	+0.083	+0.42	-0.018	+0.79	+0.388	+1.16	+1.052
-0.31	-1.133	+0.06	+0.091	+0.43	-0.020	+0.80	+0.406	+1.17	+1.064
-0.30	-1.089	+0.07	+0.097	+0.44	-0.021	+0.81	+0.425	+1.18	+1.075
-0.29	-1.047	+0.08	+0.101	+0.45	-0.022	+0.82	+0.444	+1.19	+1.086
-0.28	-1.008	+0.09	+0.104	+0.46	-0.022	+0.83	+0.464	+1.20	+1.097
-0.27	-0.969	+0.10	+0.105	+0.47	-0.020	+0.84	+0.485	+1.21	+1.109
-0.26	-0.932	+0.11	+0.105	+0.48	-0.018	+0.85	+0.506	+1.22	+1.120
-0.25	-0.895	+0.12	+0.104	+0.49	-0.014	+0.86	+0.527	+1.23	+1.132
-0.24	-0.859	+0.13	+0.103	+0.50	-0.008	+0.87	+0.549	+1.24	+1.143
-0.23	-0.823	+0.14	+0.100	+0.51	-0.002	+0.88	+0.570	+1.25	+1.154
-0.22	-0.787	+0.15	+0.098	+0.52	+0.007	+0.89	+0.592	+1.26	+1.166
-0.21	-0.751	+0.16	+0.095	+0.53	+0.016	+0.90	+0.615	+1.27	+1.177
-0.20	-0.715	+0.17	+0.091	+0.54	+0.023	+0.91	+0.637	+1.28	+1.189
-0.19	-0.678	+0.18	+0.088	+0.55	+0.033	+0.92	+0.660	+1.29	+1.200
-0.18	-0.641	+0.19	+0.084	+0.56	+0.044	+0.93	+0.682	+1.30	+1.211
-0.17	-0.604	+0.20	+0.080	+0.57	+0.055	+0.94	+0.704	+1.31	+1.221
-0.16	-0.566	+0.21	+0.076	+0.58	+0.066	+0.95	+0.725	+1.32	+1.231
-0.15	-0.528	+0.22	+0.072	+0.59	+0.078	+0.96	+0.745	+1.33	+1.240
-0.14	-0.489	+0.23	+0.068	+0.60	+0.091	+0.97	+0.766	+1.34	+1.248
-0.13	-0.451	+0.24	+0.064	+0.61	+0.104	+0.98	+0.786	+1.35	+1.255
-0.12	-0.412	+0.25	+0.059	+0.62	+0.118	+0.99	+0.805	+1.36	+1.261
-0.11	-0.373	+0.26	+0.055	+0.63	+0.132	+1.00	+0.824	+1.37	+1.266
-0.10	-0.334	+0.27	+0.050	+0.64	+0.146	+1.01	+0.842	+1.38	+1.269
-0.09	-0.295	+0.28	+0.046	+0.65	+0.160	+1.02	+0.860	+1.39	+1.271
-0.08	-0.257	+0.29	+0.041	+0.66	+0.175	+1.03	+0.877	+1.40	+1.271
-0.07	-0.220	+0.30	+0.037	+0.67	+0.190	+1.04	+0.894	+1.41	+1.270
-0.06	-0.181	+0.31	+0.032	+0.68	+0.205	+1.05	+0.910	+1.42	+1.266
-0.05	-0.147	+0.32	+0.027	+0.69	+0.221	+1.06	+0.925	+1.43	+1.261
-0.04	-0.113	+0.33	+0.022	+0.70	+0.236	+1.07	+0.940	+1.44	+1.254
-0.03	-0.081	+0.34	+0.017	+0.71	+0.252	+1.08	+0.954	+1.45	+1.246
-0.02	-0.051	+0.35	+0.013	+0.72	+0.268	+1.09	+0.967	+1.46	+1.236
-0.01	-0.022	+0.36	+0.008	+0.73	+0.284	+1.10	+0.980	+1.47	+1.224
+0.00	+0.000	+0.37	+0.003	+0.74	+0.301	+1.11	+0.993	+1.48	+1.213
+0.01	+0.026	+0.38	-0.002	+0.75	+0.317	+1.12	+1.006	+1.49	+1.203
+0.02	+0.044	+0.39	-0.006	+0.76	+0.334	+1.13	+1.018	+1.50	+1.191
+0.03	+0.060	+0.40	-0.010	+0.77	+0.352	+1.14	+1.029	+1.51	+1.178
								+1.52	+1.166

adjusted from their averaged equivalent B0-star reddenings using the transformation of Fernie (1963). Such corrections are fairly straightforward given that the mean colors of each Cepheid are known, but do depend upon the accuracy of the adopted Fernie (1963) relationship.

Another misconception is that reddening increases linearly with distance (e.g., Laney & Caldwell 2007), whereas in reality reddening increases discretely whenever the line of sight crosses an interstellar (dust) cloud causing extinction (see Turner et al. 2011). In some directions, such as a few degrees away from the Galactic equator, all reddened stars may be affected by light passing through only one relatively nearby dust cloud, no matter how distant they lie beyond the cloud. That makes the derivation of space reddening for many Cepheids relatively uncomplicated. They need not belong to a cluster or association, as long as there are some early-type stars in the same line of sight.

A problem arose with the reddening for the Cepheids QZ Nor and V340 Nor, associated with the open cluster NGC 6067. Majaess et al. (2013a) adopted reddenings of $E_{J-H} = 0.13$ for the cluster and Cepheids based upon a spectroscopic reddening of $E_{B-V}(B0) = 0.35 \pm 0.04$ tied to spectra and *UBV* photometry by Thackeray, Wesselink & Harding (1962) for NGC 6067. A re-examination of the *UBV* colors for cluster stars from Thackeray et al. (1962) and Pedreros (1984) in conjunction with a likely reddening law for the field (see Turner 1976c; Turner et al. 2014) revealed that the cluster displays differential reddening. Color excesses for cluster stars vary from $E_{B-V} = 0.30$ to $E_{B-V} = 0.48$, with most stars reddened by 0.36–0.42. The E_{J-H} reddening adopted by Majaess et al. (2013a) corresponds to $E_{B-V} = 0.42$ with the relations of Turner (2011), in contrast to the mean spectroscopic reddening. Since V340 Nor lies in the nuclear region of NGC 6067 while QZ Nor is a coronal member, the field reddening for both Cepheids must be examined separately.

The available data for stars lying within 2 arc-min of each Cepheid were therefore reanalyzed using the methodology described previously (see also Turner 1996). The observations were primarily the *UBV* data of Thackeray et al. (1962) and Pedreros (1984), *BV* data from APASS, and *JHK_s* data (Cutri et al. 2003) from the 2MASS survey (Skrutskie et al. 2006). Stars near V340 Nor displayed a reddening of $E_{B-V} = 0.39 \pm 0.02$ (transformed from $E_{J-H} = 0.115 \pm 0.005$), which was adopted for the Cepheid space reddening given in Table 2. There

were fewer stars available to test the space reddening near QZ Nor, but they did appear to be consistent with a color excess identical to that for V340 Nor. That reddening was therefore adopted for the space reddening of QZ Nor in Table 2.

The sample of 68 Cepheids of known space reddening is summarized in Table 2, where the sources used to establish the color excesses are indicated in the Table notes. The reddening of each Cepheid is based upon studies of the surrounding field in which the exact form of the reddening law in the field was (usually) established beforehand. The space reddening sample is clearly quite extensive, although a few of the Cepheids are of uncertain use given questions about their population types (see Turner et al. 2011).

3. TESTING THE SAMPLE

The Cepheid space reddenings of Table 2 were first compared with the color excesses of Turner, Leonard & English (1987), which were derived using reddening-independent *KHG* data from Feltz & McNamara (1980) in conjunction with a small set of space reddenings by Turner (1984), where some of the latter differed slightly from the values presented here. There are 16 stars in common to the two data samples, the comparison being depicted in Figure 2 (top). The photometric *KHG* index of Brigham Young University is essentially independent of interstellar reddening (see discussion by Turner et al. 1987), and *KHG* reddenings depend upon Cepheid space reddenings only for calibration purposes. They therefore make a useful first test of the reddening compilation being established here.

A combination of least squares and non-parametric straight line fits to the data of Figure 2 (top) gave the formal results summarized in Table 3 and depicted in the figure. Despite small differences in the reddenings of the calibrating objects for the original *KHG* reddenings (Turner et al. 1987), they appear to agree closely enough with the Table 2 data to be considered a reasonable match. Thereafter, the color excesses from Turner et al. (1987) were used as published to augment the system of reddenings in Table 2. The resulting standard system of 91 objects contains 52 stars from Table 2, average values for 16 stars in common to the two systems, and 23 additional Cepheids from Turner et al. (1987). The system of standard reddenings denoted here as “Std” refers to the augmented system in the other comparisons made in Figure 2.

The reddening system of Kovtyukh et al. (2008) was expected to fall very close to the space and *KHG* reddening scales because of the manner in

TABLE 2
SPACE REDDENINGS FOR CEPHEIDS

Cepheid	Field	E_{B-V}	Source	Cepheid	Field	E_{B-V}	Source
T Ant	Anon Ant OB	0.30 ± 0.01	24	AC Mon	Field	0.55 ± 0.01	34
U Aql	Field	0.35 ± 0.01	2, 44	CV Mon	van den Bergh 1	0.75 ± 0.02	22
FF Aql	Field	0.25 ± 0.01	28, 41	S Nor	NGC 6087	0.17 ± 0.01	12
RT Aur	Field	0.06 ± 0.03	28	TW Nor	Lyngä 6	1.29 ± 0.10	35
EW Aur	Field	0.58 ± 0.03	34	QZ Nor	NGC 6067	0.39 ± 0.02	42, 44
OX Cam	Tombaugh 5	0.75 ± 0.07	29	V340 Nor	NGC 6067	0.39 ± 0.02	42, 44
XZ CMa	Tombaugh 1	0.31 ± 0.01	9	UY Per	King 4	0.89 ± 0.05	33
YZ CMa	Field	0.56 ± 0.03	34	HZ Per	Field	1.36 ± 0.04	34
CN CMa	Field	0.63 ± 0.02	34	OT Per	Field	1.39 ± 0.08	34
ℓ Car	Field	0.17 ± 0.02	28	AQ Pup	Turner 14	0.47 ± 0.07	39
U Car	Anon Car OB	0.30 ± 0.03	14	BD Pup	Field	0.67 ± 0.02	34
SX Car	Ruprecht 91	0.20 ± 0.03	25, 44	BE Pup	Field	0.64 ± 0.02	34
VY Car	Ruprecht 91	0.36 ± 0.03	25, 44	LR Pup	Field	0.42 ± 0.01	34
WZ Car	Collinder 236	0.27 ± 0.01	32	V620 Pup	Turner 13	0.64 ± 0.02	39
GT Car	Teutsch 106	0.66 ± 0.02	17	S Sge	Field	0.09 ± 0.01	2, 44
SU Cas	Alessi 90	0.33 ± 0.02	38	GY Sge	Sge OB1	1.15 ± 0.02	6
CE Cas	NGC 7790	0.49 ± 0.05	43	U Sgr	M25	0.49 ± 0.03	43
CF Cas	NGC 7790	0.49 ± 0.05	43	W Sgr	Field	0.12 ± 0.01	28
CG Cas	Berkeley 58	0.64 ± 0.02	30	X Sgr	Field	0.19 ± 0.01	28
DL Cas	NGC 129	0.51 ± 0.01	16	Y Sgr	Field	0.22 ± 0.01	28
FO Cas	Field	0.76 ± 0.05	34	VY Sgr	Field	1.24 ± 0.04	34
IO Cas	Field	0.59 ± 0.02	34	WZ Sgr	Turner 2	0.56 ± 0.01	17
V Cen	NGC 5662b	0.28 ± 0.01	8	AY Sgr	Field	0.94 ± 0.02	34
V810 Cen	Stock 14	0.25 ± 0.01	7	BB Sgr	Collinder 394	0.30 ± 0.02	10, 11
δ Cep	Cep OB6	0.07 ± 0.01	23, 36	V1882 Sgr	Field	0.64 ± 0.01	34
X Cyg	Ruprecht 173/5	0.25 ± 0.02	21	RU Sct	Trumpler 35	0.95 ± 0.02	4
SU Cyg	Turner 9	0.15 ± 0.01	19, 20	V367 Sct	NGC 6649	1.27 ± 0.02	5
DT Cyg	Field	0.07 ± 0.01	2, 44	SZ Tau	NGC 1647	0.29 ± 0.01	9, 15
V1726 Cyg	Platais 1	0.33 ± 0.02	18, 27	SW Vel	Vel OB1c	0.35 ± 0.02	17
β Dor	Field	0.05 ± 0.02	3, 28, 44	CS Vel	Ruprecht 79	0.75 ± 0.01	13
ζ Gem	Field	0.02 ± 0.01	28, 37	S Vul	Turner 1	1.01 ± 0.01	44
HD 18391	TKMLM 1	1.02 ± 0.09	31	T Vul	Field	0.11 ± 0.02	2, 28, 44
T Mon	Mon OB 2	0.23 ± 0.04	1	SV Vul	Vul OB1	0.45 ± 0.01	9
UY Mon	Field	0.11 ± 0.01	34	α UMi	Harrington 1	0.02 ± 0.01	26, 40

Source: (1) Turner (1976b), (2) Feltz & McNamara (1976), (3) Turner (1980a), (4) Turner (1980b), (5) Turner (1981), (6) Forbes (1982), (7) Turner (1982a), (8) Turner (1982b), (9) Turner (1983), (10) Turner (1984), (11) Turner & Pedreros (1985), (12) Turner (1986), (13) Walker (1987), (14) Turner (1988), (15) Turner (1992), (16) Turner et al. (1992), (17) Turner et al. (1993), (18) Turner et al. (1994), (19) Turner et al. (1997), (20) Turner et al. (1998a), (21) Turner (1998), (22) Turner et al. (1998b), (23) Benedict et al. (2002), (24) Turner & Berdnikov (2003), (25) Turner et al. (2005), (26) Turner (2005), (27) Turner et al. (2006b), (28) Benedict et al. (2007), (29) Majaess et al. (2008), (30) Turner et al. (2008), (31) Turner et al. (2009a), (32) Turner et al. (2009b), (33) Turner et al. (2010), (34) Turner et al. (2011), (35) Majaess et al. (2011), (36) Majaess et al. (2012a), (37) Majaess et al. (2012b), (38) Turner et al. (2012a), (39) Turner et al. (2012b), (40) Turner et al. (2013a), (41) Turner et al. (2013b), (42) Majaess et al. (2013a), (43) Majaess et al. (2013b), (44) This paper.

which the color excesses were derived relative to a stellar atmosphere model-based effective temperature scale. But differences were noted previously for those reddenings with respect to the empirical color-effective temperature scale of Gray (1992) by Turner et al. (2013a,b), possibly arising from the surface gravity term included in the derivation of colors from effective temperatures by Kovtyukh et al. (2008). A comparison of all overlapping reddenings in Figure 2 (middle) indicates a small zero-point offset and a reddening-dependent term between the two scales, as noted by the fit in Table 3.

The Kovtyukh et al. (2008) color excesses are readily tied to the system of space and *KHG* reddenings using the relationship of Table 3, which was the procedure used here to augment the standard system of reddenings. There is moderate scatter in the reddening comparison, however, and it is not clear in which system it originates. About a third of the reddenings in the Kovtyukh et al. (2008) sample are tied to only 1 or 2 spectra, and that applies to the most deviant data in Figure 2 (middle). The remaining two thirds of the sample is tied to Cepheids with anywhere from 3 to 26 spectra, so lack of spectroscopic sampling cannot be the origin of the scatter. The adjusted reddenings from Kovtyukh et al. (2008) were therefore assimilated into the standard system with some care, particularly for deviant Cepheids in the comparison sample.

The reddening system of Andrievsky et al. (2012) was derived in analogous fashion to that of Kovtyukh et al. (2008), so it was expected to yield similar results, given its stellar atmosphere model-based effective temperature scale. All stars in their compilation were analyzed previously for reddening by Kovtyukh et al. (2008). The comparison in Figure 2 (bottom) indicates both a zero-point offset and a larger reddening dependence for the Andrievsky et al. (2012) reddenings. Presumably both arise because the Andrievsky et al. (2012) reddening system was calibrated using an intrinsic color relation tied to both T_{eff} and $\log g$, as in Kovtyukh et al. (2008), but normalized relative to the reddening scale of Storm et al. (2011). In order to agree with the present reddening scale, it is necessary to apply a zero-point offset of +0.05 and a scaling factor of 0.89 to the Andrievsky et al. (2012) color excesses, specifically by the relationship given in Table 3. The scatter in the latter is slightly smaller than for the Kovtyukh et al. (2008) reddenings, most likely because they are restricted to Cepheids well-sampled spectroscopically.

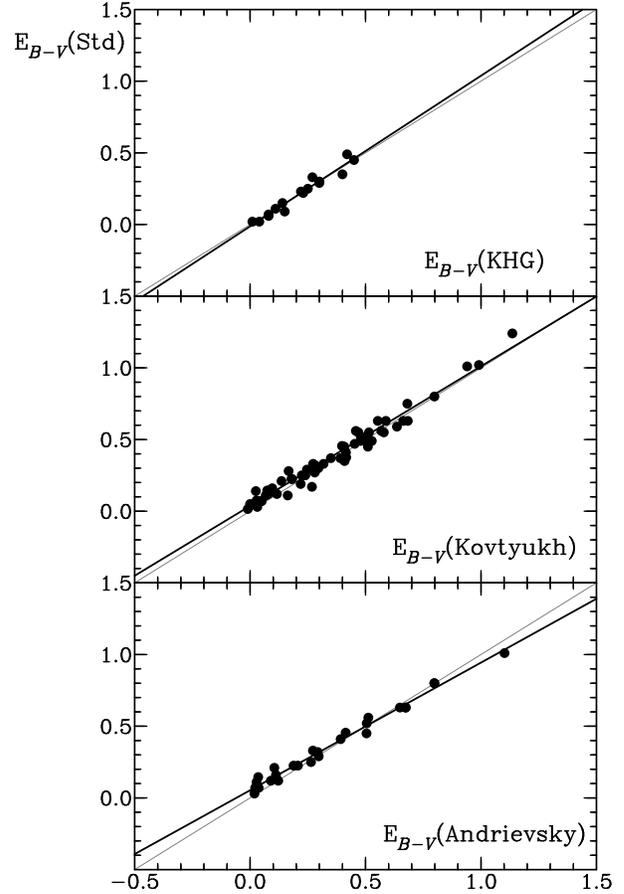


Fig. 2. A comparison of the adopted standard color excesses (see text) with reddening compilations by Turner et al. (1987, top), Kovtyukh et al. (2008, middle), and Andrievsky et al. (2012, bottom). The y -axis values are those of Table 1 supplemented by *KHG* reddenings in subsequent comparisons. Gray lines represent results expected for coincident reddening scales, and black lines represent fits to the data.

The availability of two or more independent estimates of reddening for each Cepheid formed the rationale for creating a more extensive standard system of 198 Cepheids by averaging the adjusted color excesses of Andrievsky et al. (2012) with the space reddenings, *KHG* reddenings, and adjusted Kovtyukh et al. (2008) reddenings.

The reddening system of the Storm et al. (2011) study was established by applying the infrared surface-brightness (IRSB) method to their Cepheid sample, with the Cardelli et al. (1989) extinction law used to adjust the observed star brightnesses in different wavelength bands. A comparison of the derived reddenings from that system with the standard system developed here is depicted in Figure 3 (top),

and a regression fit is tabulated in Table 3. The zero-point for the Storm et al. (2011) system agrees closely with the present standard (Std) reddening system, but there is moderate scatter and a reddening-dependent trend opposite to that for the Andrievsky et al. (2012) system. The effect may be related to the manner of making extinction corrections, given that the Cardelli et al. (1989) extinction law has recently been shown to conflict with actual extinction parameters in many Galactic fields (Zagury 2013; Turner et al. 2014). The large scatter for some Cepheids can be explained by the fact that they are southern hemisphere objects for which only Kovtyukh et al. (2008) reddenings are available for comparison, often Cepheids sampled by Kovtyukh et al. (2008) with only a single spectrum at one phase of the light curve.

An initial melding with the present system of the Storm et al. (2011) reddenings, modified according to the relationship in Table 3, created noticeable discrepancies for three stars. Two in particular, WZ Car and S Vul, had solid field reddenings in Table 2, while the reddening for VZ Pup differed by $\sim 0^m.2$ from the reddening derived by Kovtyukh et al. (2008). Initially it was decided to include only the adjusted reddenings of Storm et al. (2011) for the 57 least deviant objects in the extended standard system, and subsequent analyses were carried out accordingly. It was later noticed that the Kovtyukh et al. (2008) reddening for VZ Pup created an anomalous intrinsic color for the variable, whereas that from the adjusted Storm et al. (2011) reddening did not. VZ Pup is a 23^d Cepheid, and the long-period pulsators, in particular, evolve the fastest, display the most rapid rates of period change (Turner et al. 2006a), and undergo the largest random period fluctuations (Turner et al. 2009c). Given that the Kovtyukh et al. (2008) reddening for the star is tied to only one spectrum, there is a possibility that the derived light curve phase for that spectrum is erroneous, thereby affecting the reddening in systematic fashion. None of the other Cepheid reddening comparisons involve VZ Pup in significant fashion, so the later decision to exclude the transformed Kovtyukh et al. (2008) reddening for the star in favor of the transformed Storm et al. (2011) reddening does not impact the analyses.

The close agreement of the Table 2 reddenings with results obtained using the *KHG* index (Turner et al. 1987), the reddening-independent parameter developed at Brigham Young University for measuring the strength of the stellar G band relative to the Ca II K line and *Hδ*, raises the question of whether

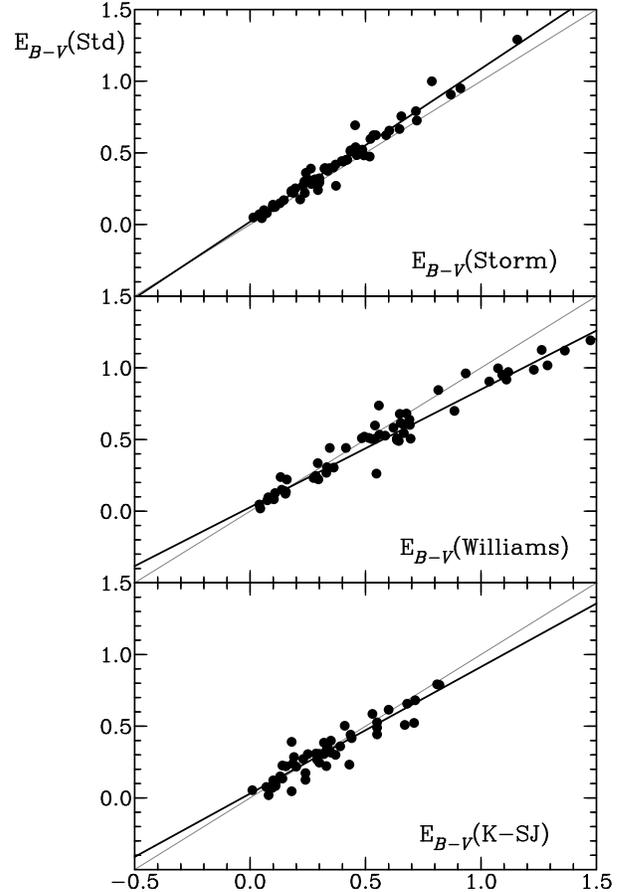


Fig. 3. A similar comparison to that of Figure 2 for x -axis color excesses from Storm et al. (2011, top), Williams (1966, middle), and an average of the reddenings from Kraft (1960) and Spencer Jones (1989, bottom), with y -axis values tied to the extended 200-star standard system. Black and gray lines are as in Figure 2.

similar results apply to reddenings derived in other systems employing a reddening-free parameter. The reddening system of Williams (1966), for example, was calibrated using cluster Cepheids and the break in the continuum across the *G* band in Cepheid spectra, in similar fashion to the method of establishing T_{eff} from *KHG* photometry. The Strömrgren system reddenings, E_{b-y} , of Williams (1966) were converted to Johnson system reddenings, E_{B-V} , using the 0.73 factor adopted by Turner et al. (1987), and are compared with the present results in Figure 3 (middle), as well as in Table 3. The scatter in the reddening comparison is large, making the Williams (1966) color excesses of little value for averaging into our standard system. The scaling factor of 0.82 may relate to the manner in which the system was cal-

TABLE 3
REDDENING COMPARISONS

X -coordinate	a (zero-point)	b (slope)	n
$E_{B-V}(KHG)$	-0.013 ± 0.009	1.050 ± 0.049	16
$E_{B-V}(\text{Kovtyukh})$	$+0.037 \pm 0.006$	0.974 ± 0.019	58
$E_{B-V}(\text{Andrievsky})$	$+0.053 \pm 0.008$	0.890 ± 0.016	26
$E_{B-V}(\text{Storm})$	$+0.020 \pm 0.006$	1.067 ± 0.019	64
$E_{B-V}(\text{Williams})$	$+0.028 \pm 0.009$	0.822 ± 0.020	56
$E_{B-V}(\text{K-SJ})$	$+0.030 \pm 0.013$	0.884 ± 0.038	48
$E_{B-V}(\text{Mianes})$	$+0.013 \pm 0.026$	0.992 ± 0.063	19
$E_{B-V}(\text{PBo})$	$+0.006 \pm 0.009$	0.995 ± 0.028	40
$E_{B-V}(\text{PBe})$	$+0.013 \pm 0.009$	1.091 ± 0.034	43
$E_{B-V}(\text{FM80})$	$+0.019 \pm 0.006$	0.861 ± 0.017	37
$E_{B-V}(\text{Fe87})$	$+0.045 \pm 0.007$	0.887 ± 0.024	38
$E_{B-V}(\text{Gray})$	$+0.050 \pm 0.004$	0.851 ± 0.015	40
$E_{B-V}(\text{TY})$	$+0.020 \pm 0.009$	0.901 ± 0.019	105
$E_{B-V}(\text{Harris})$	$+0.082 \pm 0.011$	0.878 ± 0.018	70
$E_{B-V}(\text{Dean})$	$+0.014 \pm 0.013$	1.007 ± 0.038	20
$E_{B-V}(\text{Pel})$	$+0.077 \pm 0.008$	0.876 ± 0.021	62
$E_{B-V}(\text{Bersier})$	$+0.045 \pm 0.008$	0.849 ± 0.030	29
$E_{B-V}(\text{Kiss})$	$+0.066 \pm 0.016$	0.738 ± 0.057	18
$E_{B-V}(\text{Ferne67})$	-0.047 ± 0.012	1.056 ± 0.027	111
$E_{B-V}(\text{Ferne90})$	$+0.053 \pm 0.005$	0.922 ± 0.011	157
$E_{B-V}(\text{Sasselov})$	$+0.029 \pm 0.013$	0.707 ± 0.096	5
$E_{B-V}(\text{Krockenberger})$	$+0.016 \pm 0.008$	0.849 ± 0.038	11
$E_{B-V}(\text{DWC})$	$+0.017 \pm 0.010$	0.988 ± 0.027	34
$E_{B-V}(\text{LS})$	$+0.051 \pm 0.014$	0.900 ± 0.021	20
$E_{B-V}(\text{LC})$	$+0.032 \pm 0.004$	1.000 ± 0.014	40

ibrated using the data inferred for cluster Cepheids in that era, or to how reddening was treated in the Strömgren system.

Γ -index reddenings (Kraft 1960; Spencer Jones 1989), which are also tied to the strength of the G band in stellar spectra, as measured photometrically, are presumably reddening-free as well. They were calibrated in the original studies using known reddenings for FG-type supergiants and cluster Cepheids. A comparison of our standard system reddenings with Γ -index color excesses is depicted in Figure 3 (bottom), where the Γ reddenings from Kraft (1960) and Spencer Jones (1989) were averaged in an attempt to reduce the rather sizeable scatter in each. A regression fit to the data is given in Table 3. The Γ -index reddenings display large scatter no matter how the data are combined. They were therefore omitted from possible inclusion in an augmented standard system.

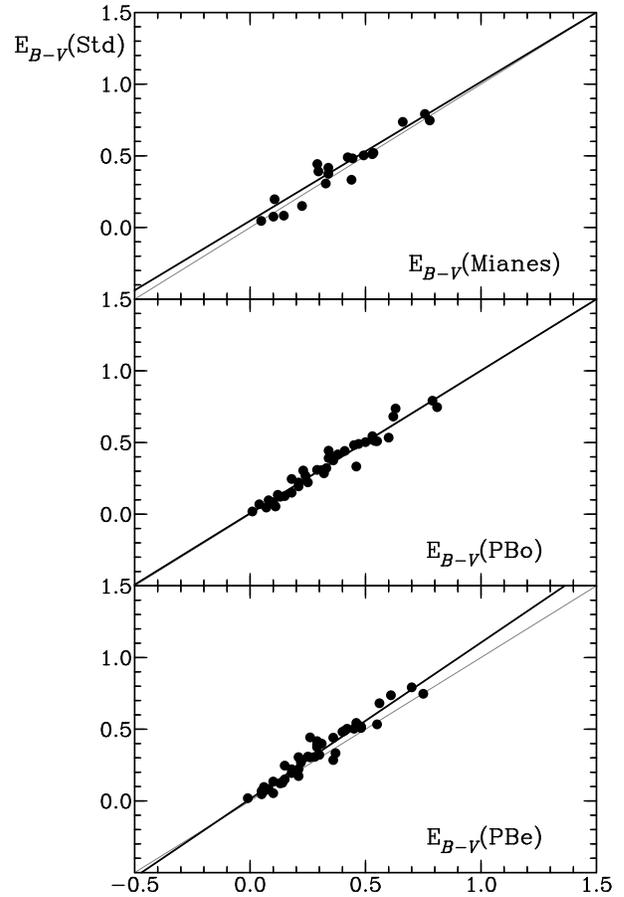


Fig. 4. A similar comparison to that of Figures 2 & 3 for color excess compilations by Mianes (1963, top), Parsons & Bouw (1971, middle), and Parsons & Bell (1975, bottom). Lines are as in Figure 2.

The comparisons so far have been to systems that should be independent of any period-color formulation. That is not always the case for other published reddening systems. Three comparisons with less well-defined systems are presented in Figure 4, and summarized in Table 3. All three systems are tied to Lick six-color ($UVBGRI$) photometry (Kron 1958; Stebbins & Kron 1964) of Cepheids by Mianes (1963) and the sources cited by Parsons (1971). Figure 4 (top) compares the reddenings of Mianes (1963) with the present standard system, the former being converted from $G-I$ reddenings to $B-V$ reddenings using the 1.89 conversion factor of Schneider (1969). The Mianes (1963) reddenings appear to agree overall with the color excesses of Table 2, but the scatter for individual Cepheids is too large to make them suitable for incorporation into the standard system.

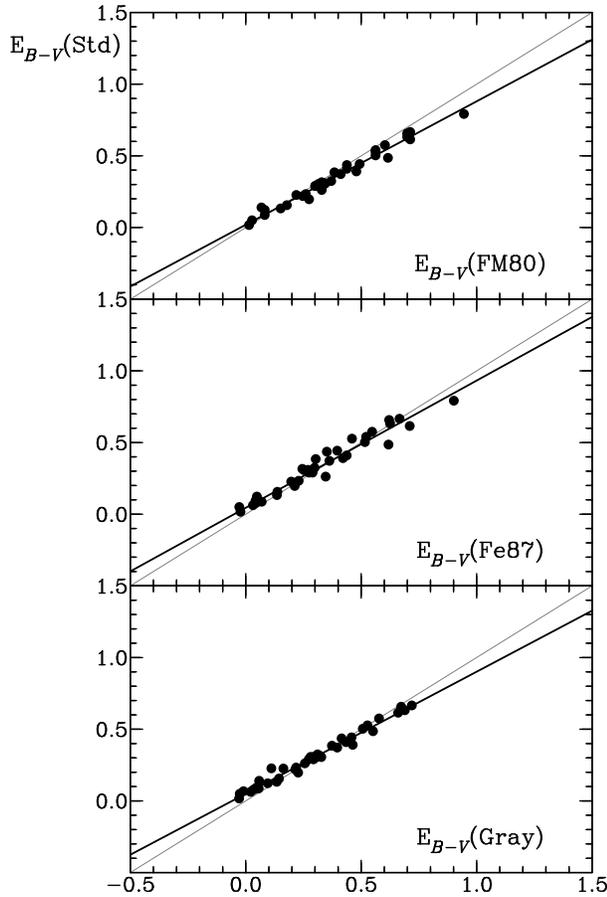


Fig. 5. A similar comparison to that of Figures 2 & 3 for color excess compilations by Feltz & McNamara (1980, top), Fernie (1987, middle), and Gray (1991, bottom). Lines are as in Figure 2.

Figure 4 (middle) is a comparison with the reddenings of Parsons & Bouw (1971), derived by means of model atmosphere and stellar evolutionary model fits to Lick six-color photometry of Cepheids. Figure 4 (bottom) presents similar results for the subsequent study by Parsons & Bell (1975) using a similar methodology with a Q parameter. So-called “reddening-free” Q parameters depend directly on the reddening law adopted, which, as noted in §2, is variable around the Galaxy. Reddenings tied to Q parameters are therefore inherently unreliable, and that applies here to the Parsons & Bell (1975) system linked to Lick six-color photometry.

The reddenings of Parsons & Bell (1975) and Parsons & Bouw (1971) have long been considered to be of high quality, although there is a negative offset of the former relative to the latter, presumably arising from the different methods of analysis. The generally small scatter in Figure 4 (middle) and the small un-

certainties in the regression fit (Table 3) confirm the relatively high quality of the Parsons & Bouw (1971) reddenings, despite uncertainties in how interstellar reddening was treated. The Cepheids in the Parsons & Bouw (1971) sample were therefore included in the standard system without adjustment, although the residual scatter is of concern. Similar statements do not apply to the Parsons & Bell (1975) reddenings, which display slightly larger scatter and a noticeable reddening-dependent trend. It was decided to omit them from inclusion in the standard system.

Figure 5 presents comparisons with the present standard system for three reddening sources on the Strömgren system, all derived from the observations of Feltz & McNamara (1980): the original color excesses of Feltz & McNamara (1980, top), those of Fernie (1987, middle), and those of Gray (1991, bottom). As was the case for the reddenings of Williams (1966), the Strömgren system reddenings E_{b-y} were converted to Johnson system reddenings E_{B-V} using the factor of 0.73 adopted by Turner et al. (1987). All three sources compared in Figure 5 and summarized in Table 3 appear to display reddening-dependent offsets that likely originate from the adoption of reddening-free indices for Strömgren photometry that are not strictly “reddening-free” according to the arguments presented above. Despite that, the scatter about the best-fitting relations is generally small, although that for the Gray (1991) reddenings is sufficiently tight to allow them to be amalgamated into the present standard system after adjustment for a zero-point offset of +0.05 and a scaling factor of 0.85. The actual adjustments were made using the Table 3 relationship.

Figure 6 and Table 3 present comparisons with the present system for three reddening sources tied to spectroscopic observations of Cepheids, either mixed spectroscopy/photometry as in the case of the color excesses by Tsarevsky & Yakimova (1970, Figure 6 (top)) or using photometric systems designed to make use of gravity and temperature-sensitive features in the spectra of yellow and red supergiants: the Washington system used by Harris (1981a,b, Figure 6 (middle)), with conversion from $E_{T_1-T_2}$ to E_{B-V} following Harris (1981a), and the DDO system used by Dean (1981, Figure 6 (bottom)). The Cepheid reddening system of Harris (1981a,b) was calibrated using a period-color relation tied to cluster members, so the large scatter in the resulting reddenings is accounted for by the fact that Cepheids of identical period lying in the instability strip can have appreciably different effective temperatures, hence colors. The large scatter for the red-

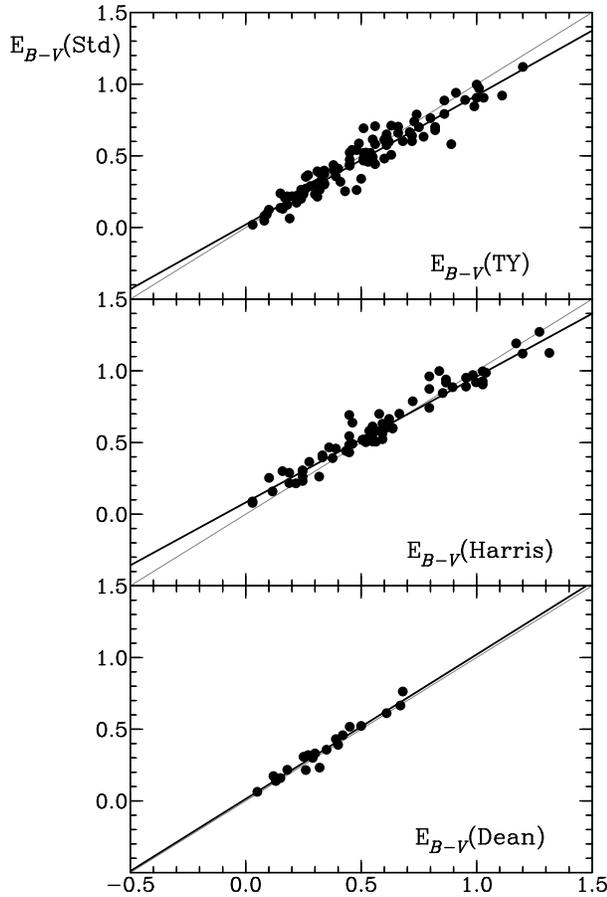


Fig. 6. A similar comparison to that of Figures 2 & 3 for color excess compilations by Tsarevsky & Yakimova (1970, top), Harris (1981a,b, middle), and Dean (1981, bottom). Lines are as in Figure 2.

denings of Tsarevsky & Yakimova (1970) may be of similar origin, given that they are tied to the spectroscopic inferences of Kraft, some of which are linked to Γ -index reddenings (Kraft 1960), previously indicated to exhibit large scatter (Figure 3). The DDO system reddenings of Dean (1981) are a remarkably close fit to the present system reddenings, however, and were therefore amalgamated into the latter as cited.

Three more reddening sources are compared with the present standard system in Figure 7, with regression fits summarized in Table 3. Figure 7 (top) shows a comparison with the reddenings of Pel (1978) on the Walraven photometric system, converted to Johnson system reddenings using the relationship adopted by Pel (1985). Figure 7 (middle) shows a comparison with the Geneva system reddenings of Bersier (1996), while Figure 7 (bottom) is a comparison with the mixed-source reddenings of Kiss &

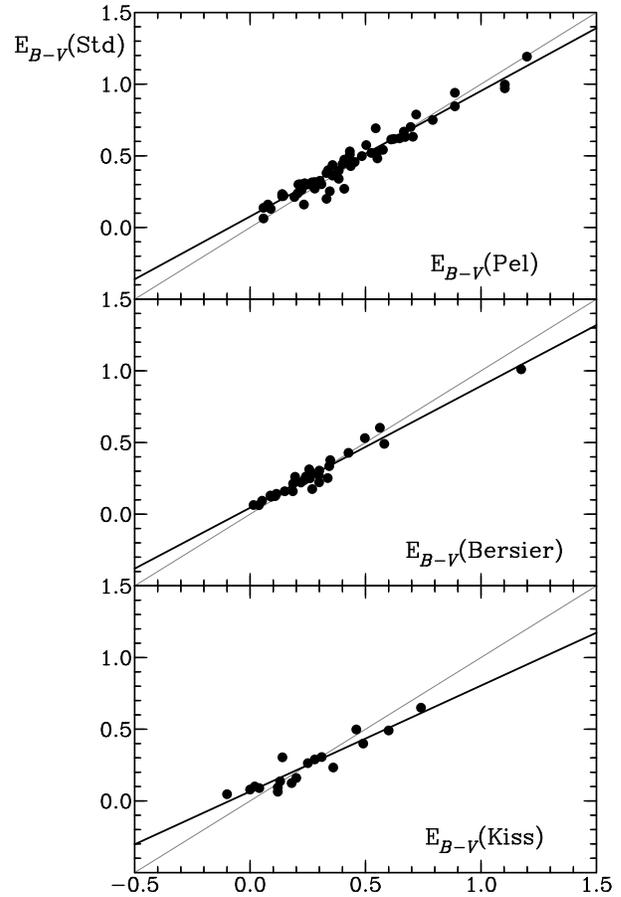


Fig. 7. A similar comparison to that of Figures 2 & 3 for color excess compilations by Pel (1978, top), Bersier (1996, middle), and Kiss & Szatmáry (1998, bottom). Lines are as in Figure 2.

Szatmáry (1998). The Walraven system reddenings (Pel 1978) do not match the present standard system, displaying large scatter as well as zero-point and scaling offsets. The reasons are unclear.

The Geneva system reddenings (Bersier 1996) are a much better match to the present system, but with sufficient scatter to make the choice of amalgamation into the standard set of this paper a difficult decision. Six Cepheids in particular do not fit the general trend, and again the reason is uncertain. Meanwhile, the final comparison shown in Figure 7 speaks to the unusual method of calibration used by Kiss & Szatmáry (1998) in their study. The large scatter, as well as the zero-point offset and large scaling term, indicate a very poor match to the present standard system. The color excesses inferred in the Kiss & Szatmáry (1998) study should not be used for studies of the intrinsic properties of Cepheids.

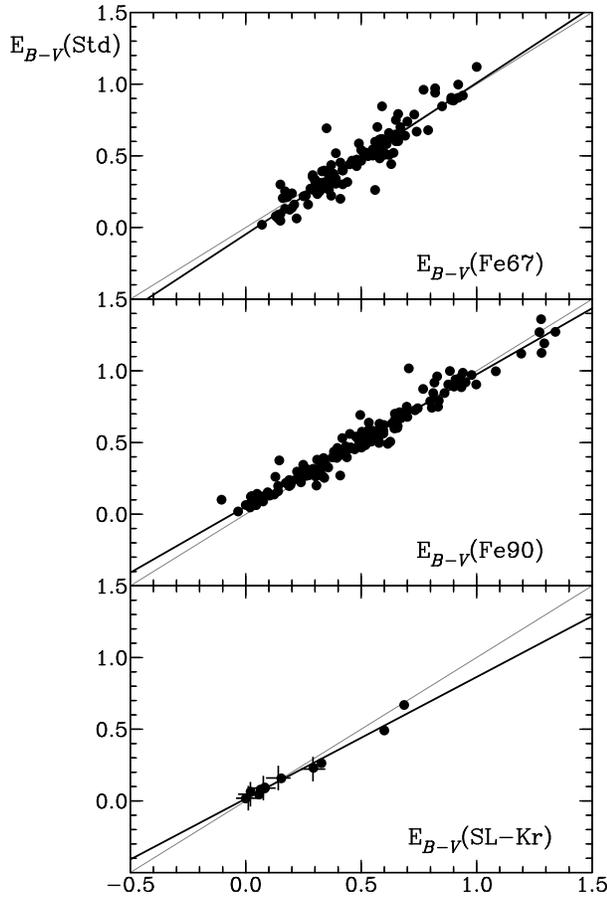


Fig. 8. A similar comparison to that of Figures 2 & 3 for color excess compilations by Fernie (1967, top), Fernie (1990a, photometric scale, middle), and Sasselov & Lester (1990, bottom, plus signs) and Krockenberger et al. (1998, bottom). Lines are as in Figure 2.

Figure 8 and Table 3 show comparisons with two large compilations of Cepheid reddenings: an older but useful study by Fernie (1967) presented in Figure 8 (top), and the oft-cited homogeneous collection of Fernie (1990a, photometric scale) presented in Figure 8 (middle). The bottom portion of the plot is a comparison with the reddenings of 5 Cepheids derived by Sasselov & Lester (1990) and 11 Cepheids (including the previous 5) derived by Krockenberger et al. (1998) using line ratios for infrared spectroscopic lines such as C I and Si I to infer effective temperatures and, hence, intrinsic colors. Both Fernie compilations went to some length to place reddenings derived in various studies on a homogeneous system, and both scales appear to be a reasonably good match with the system of space reddenings for Cepheids presented here. Reddening-dependent trends in the two compilations appear to be negli-

gible. But the scatter is large in both cases, more so in the case of the older study by Fernie (1967), so they cannot be adopted as published for amalgamation with the present standard system. Similar results apply to the cluster reddening scale of Fernie (1990a).

The comparison with the results of Sasselov & Lester (1990) and Krockenberger et al. (1998) in Figure 8 (bottom) is useful for indicating that the premise of using infrared spectroscopic line ratios for Cepheids to infer reddenings is quite sound, as also demonstrated by Kovtyukh et al. (2008) and Andrievsky et al. (2012). The reddenings for four of the five Cepheids studied by Sasselov & Lester (1990), plotted using plus signs, agree closely with the present results. The fifth Cepheid, T Mon, is somewhat discrepant with respect to most photometric reddening estimates for the star, but does follow the trend for other stars with reddenings derived in similar fashion (Krockenberger et al. 1998). The discussion by Sasselov & Lester (1990) notes the difference with respect to the Fernie (1990a) reddening scale for T Mon. The field of this Cepheid is not well-studied with regard to the reddening of nearby companions, so it may be possible to resolve the minor discrepancy for the star with additional observations of nearby companions to T Mon. Given the generally small scatter for the reddenings, the color excesses of Sasselov & Lester (1990) and Krockenberger et al. (1998) were included in the standard system of this paper using the adjustments cited in Table 3 for the Krockenberger et al. (1998) results.

Finally, Figure 9 and Table 3 present comparisons with reddening systems tied to near infrared photometry of Cepheids. The study by Dean, Warren & Cousins (1978) using Cape system *BVI* photometry generated Cepheid reddenings that are well-matched to the standard system presented here, as indicated in Figure 9 (top). They were therefore amalgamated into the present standard system with only minor concerns about the small discrepancies for some stars.

The study by Laney & Stobie (1993) focused mainly on reddenings inferred from *JHK* observations of Cepheids, and was tied to a sequence of published reddenings for Galactic calibrators. The comparison is depicted in Figure 9 (middle). The most deviant data point corresponds to the Cepheid S Vul, which did not have a solid field reddening in the era when the study was completed. Otherwise, the data display both a zero-point offset and a reddening dependence that conflict with the present system of reddenings. Residual scatter in some cases

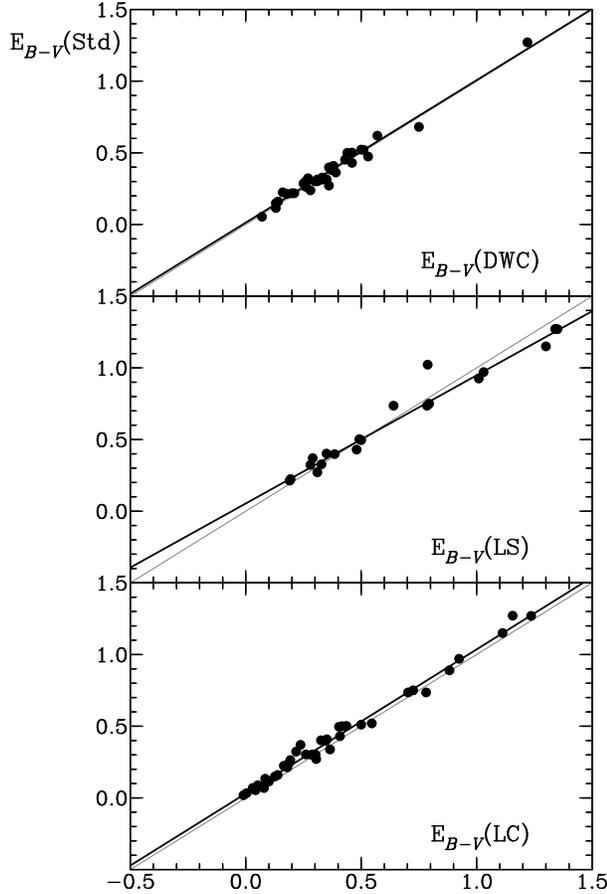


Fig. 9. A similar comparison to that of Figures 2 & 3 for color excess compilations by Dean et al. (1978, top), Laney & Stobie (1993, middle), and Laney & Caldwell (2007, bottom). Lines are as in Figure 2.

may be tied to other calibrators for which field reddenings were not well-established at the time of the Laney & Stobie (1993) study.

The study by Laney & Caldwell (2007) was intended to solidify the Cape system reddenings with reference to Cepheid field reddenings, and the generally small scatter in the comparison displayed in Figure 9 (bottom) implies a very good match to the standard system established here, except for a noticeable zero-point offset. The relationship of Table 3 found for the Laney & Caldwell (2007) reddenings was therefore used to adjust them to the present scale, in which they were subsequently included.

Reddenings have been derived by Schechter et al. (1992) for several Cepheids using *JHK* photometry, but there is no overlap of their sample, all southern hemisphere objects, with the present selection of calibrating Cepheids. It is therefore not possible to test the manner of deriving Cepheid redden-

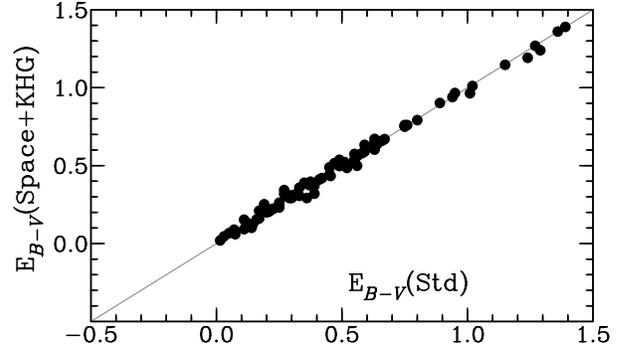


Fig. 10. A comparison of the combined space and *KHG* reddenings (*y*-axis) with the system of standard reddenings developed in this study (*x*-axis). The gray line represents coincident reddening scales.

ings from period-color relations using *JHK* colors, although presumably the smaller scatter in intrinsic colors for Cepheids observed at far infrared wavelengths should overcome most of the objections that would apply if the same methodology had adopted optical band colors such as *B-V*.

4. SUMMARY

This study presents a collection of space reddenings for 68 Cepheids, and uses the sample to test a variety of published reddening compilations for Cepheids. Initial testing was made relative to the *KHG* reddenings of Turner et al. (1987), since the Brigham Young University *KHG* index is essentially reddening independent. Similar comparisons with the spectroscopic reddenings of Kovtyukh et al. (2008) and Andrievsky et al. (2012) revealed small zero-point and reddening-dependent offsets relative to the system of space and *KHG* reddenings. Corrections for those, and use of the Turner et al. (1987) reddenings as published, permitted the creation of an expanded standard system comprising the best of the available reddenings adjusted to the space reddening and *KHG* reddening combination. The expanded standard system (Std) was used to make comparisons with more extensive systems of published reddenings, with the results presented in §3 and Figures 2–9.

The comparative best of the other reddening systems sampled, once adjusted for zero-point and scaling effects, are those of Kovtyukh et al. (2008), Andrievsky et al. (2012), Storm et al. (2011), Parsons & Bouw (1971), Gray (1991), Dean (1981), Sasselov & Lester (1990), Krockenberger et al. (1998), Dean et al. (1978), and Laney & Caldwell (2007). They are combined with the space reddenings of Table 2

and the *KHG* reddenings of Turner et al. (1987) into the standard system of reddenings summarized in Table 4, with the result for VZ Pup modified as discussed earlier. A comparison of those reddenings with the original system of space and *KHG* reddenings is presented in Figure 10. There is clearly some residual scatter in the comparison, indicating the existence of uncertainties for some stars. Note, in particular, the derived standard reddenings of $E_{B-V} = 0.32$ and 0.37 for QZ Nor and V340 Nor, respectively, relative to the common space reddenings of 0.39 adopted in §2. It implies the necessity for additional work on problem objects to resolve the small residual differences. The average standard deviation for the combined reddenings of Cepheids with several independent estimates is ± 0.028 , although for individual objects the values range from ± 0.002 to ± 0.093 . Readers should therefore use caution when using the reddenings of Table 4 to deduce intrinsic properties for specific Cepheids.

The periods used for the Cepheids in Table 4 are either the observed periods, if considered to originate from pulsation in the fundamental mode, or the observed or inferred fundamental mode period P_0 for double mode pulsators or Cepheids, such as QZ Nor, SU Cas, and V1334 Cyg, considered to be pulsating in the first overtone mode with period P_1 . Decisions about pulsation mode may be incorrect in some instances.

5. DISCUSSION

The original intent of the present study was to demonstrate that it is possible to generate a system of Cepheid color excesses that is solidly tied to the available sources of space reddening for a significant sample of stars, including in the compilation color excesses derived from reddening-free indices as a means of augmenting the new standard system (Std). An earlier version of such a compilation was described by Turner (2001), where it was noted that some generalities that could be inferred from such a compilation regarding the Cepheid instability strip differed from those described by Fernie (1990b). Similar conclusions can be made from the Cepheid sample of Table 4. Figure 11, for example, is an empirical mapping of the colors of Cepheids as a function of pulsation period, as derived from the color excess summary, and is very similar to Figure 2 of Turner (2001). Note, for example, that the color width of the instability strip changes very little as a function of pulsation period, contrary to the conclusions of Fernie (1990b).

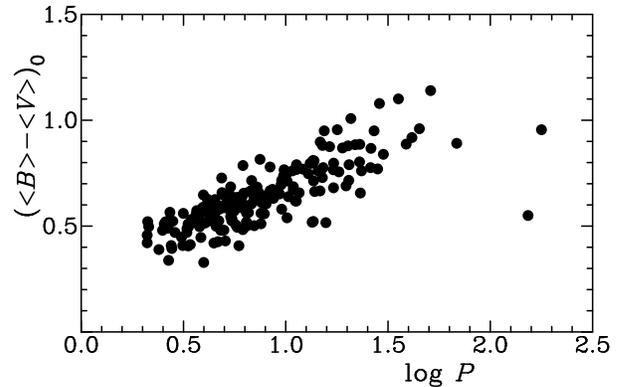


Fig. 11. Derived intrinsic $(\langle B \rangle - \langle V \rangle)_0$ colors for the Cepheids in Table 4 plotted as a function of the logarithm of the pulsational period P_0 .

The methodology described by Turner & Burke (2002) and Turner et al. (2010, see §1) can be used with the Table 4 data to map the instability strip in terms of effective temperature and luminosity, as done by Turner (2012b). Such results are not repeated here, although it is noted that they support the general conclusions made above. There is less scatter of Cepheids in the instability strip than is sometimes believed, although caution is needed when interpreting results for individual Cepheids. For example, the most deviant data in Figure 11 all relate to stars of special interest. A few are likely Type II Cepheids, for example, and the two stars of longest period are V810 Cen and HD 18391, respectively. The former appears to be double-mode Cepheid-like variable, the latter a small-amplitude double-mode pulsator, both lying blueward of the Cepheid instability strip.

A much smaller reddening is cited for HD 18391 by Arellano Ferro & Parrao (1990), who tabulated, but did not use, a value of $E_{B-V} = 0.56$ by Schmidt (1984) for the reddening of the double cluster h and χ Persei. That value is clearly much too small according to the study of HD 18391's surroundings made by Turner et al. (2009a), and would lead to an intrinsic $B-V$ color for the supergiant almost a half magnitude redder than expected for a G0 Ia star lying off the blue edge of the instability strip (see Turner et al. 2009a). Supergiants in the field around the double cluster also display strong evidence for differential reddening, with E_{B-V} ranging from 0.21 to 1.00, or larger, in this section of the Perseus spiral arm. The reddening in the core regions of h and χ Persei is therefore not representative of their outlying regions.

TABLE 4
REDDENING SUMMARY

Cepheid	$\log P$	E_{B-V}	$(B-V)_0$	Cepheid	$\log P$	E_{B-V}	$(B-V)_0$	Cepheid	$\log P$	E_{B-V}	$(B-V)_0$	Cepheid	$\log P$	E_{B-V}	$(B-V)_0$
T Ant	0.7707	0.300	0.407	CEa Cas	0.7111	0.508	0.613	BG Lac	0.7269	0.322	0.647	HW Pup	1.1289	0.724	0.518
U Aql	0.8466	0.397	0.629	CEb Cas	0.6512	0.497	0.534	V473 Lyr	0.3225	0.126	0.457	LR Pup	0.5226	0.420	0.510
SZ Aql	1.2340	0.614	0.797	CF Cas	0.6880	0.537	0.658	T Mon	1.4318	0.223	0.950	MY Pup	0.7555	0.142	0.506
TT Aql	1.1385	0.485	0.808	CG Cas	0.6401	0.640	0.580	SV Mon	1.1828	0.267	0.776	V620 Pup	0.4126	0.640	0.520
FF Aql	0.6504	0.231	0.524	CH Cas	1.1786	0.920	0.729	TW Mon	0.8511	0.701	0.636	S Sge	0.9234	0.135	0.667
FM Aql	0.7863	0.633	0.638	CY Cas	1.1577	0.905	0.761	TX Mon	0.9396	0.482	0.631	GY Sge	1.7081	1.147	1.140
FN Aql	0.9769	0.521	0.695	DD Cas	0.9918	0.465	0.766	TZ Mon	0.8709	0.473	0.655	U Sgr	0.8290	0.434	0.666
V496 Aql	0.8330	0.442	0.715	DL Cas	0.9031	0.523	0.630	UY Mon	0.3799	0.153	0.389	W Sgr	0.8805	0.133	0.612
V600 Aql	0.8597	0.845	0.684	FM Cas	0.7641	0.353	0.635	WW Mon	0.6686	0.707	0.426	X Sgr	0.8459	0.252	0.502
V733 Aql	0.7909	0.140	0.786	FO Cas	0.8324	0.760	0.605	XX Mon	0.7369	0.631	0.531	Y Sgr	0.7614	0.222	0.635
V1162 Aql	0.7305	0.196	0.685	IO Cas	0.7485	0.590	0.572	AA Mon	0.5953	0.742	0.589	VY Sgr	1.1322	1.192	0.798
V1344 Aql	0.8738	0.567	0.815	V379 Cas	0.6341	0.621	0.525	AC Mon	0.9039	0.575	0.605	WZ Sgr	1.3395	0.499	0.884
η Aql	0.8559	0.155	0.634	V636 Cas	0.9231	0.603	0.779	CU Mon	0.6728	0.874	0.525	XX Sgr	0.8078	0.586	0.562
V340 Ara	1.3183	0.563	1.008	V Cen	0.7399	0.305	0.576	CV Mon	0.7307	0.760	0.545	YZ Sgr	0.9802	0.319	0.716
RT Aur	0.5715	0.081	0.510	V810 Cen	2.1844	0.250	0.550	EE Mon	0.6821	0.533	0.480	AP Sgr	0.7040	0.217	0.587
RX Aur	1.0654	0.301	0.636	CP Cep	1.2519	0.679	0.955	EK Mon	0.5975	0.542	0.647	AV Sgr	1.1879	1.125	0.950
SY Aur	1.0062	0.466	0.539	CR Cep	0.7947	0.739	0.655	FG Mon	0.6529	0.646	0.541	AY Sgr	0.8175	0.940	0.564
YZ Aur	1.2599	0.602	0.756	IR Cep	0.3251	0.395	0.521	FI Mon	0.5169	0.628	0.488	BB Sgr	0.8220	0.308	0.662
AN Aur	1.0124	0.581	0.639	V351 Cep	0.4481	0.462	0.523	V465 Mon	0.4335	0.175	0.565	V350 Sgr	0.7122	0.318	0.588
BK Aur	0.9032	0.423	0.636	δ Cep	0.7297	0.089	0.572	V504 Mon	0.4431	0.610	0.394	V1882 Sgr	0.4332	0.640	0.530
CY Aur	1.1414	0.918	0.663	BG Cru	0.5241	0.122	0.496	V508 Mon	0.6163	0.252	0.627	RV Sco	0.7826	0.376	0.587
ER Aur	1.1958	0.639	0.516	X Cyg	1.2145	0.262	0.875	V526 Mon	0.4273	0.249	0.338	RY Sco	1.3078	0.788	0.716
EW Aur	0.4248	0.580	0.490	SU Cyg	0.5850	0.125	0.446	S Nor	0.9892	0.210	0.738	KQ Sco	1.4577	0.927	1.079
V335 Aur	0.5331	0.720	0.412	SZ Cyg	1.1793	0.598	0.881	TW Nor	1.0329	1.240	0.761	V500 Sco	0.9693	0.622	0.652
RX Cam	0.8983	0.556	0.640	TX Cyg	1.1676	1.120	0.666	QZ Nor	0.7333	0.320	0.585	Z Sct	1.1106	0.582	0.747
TV Cam	0.7239	0.506	0.629	VX Cyg	1.3039	0.845	0.879	V340 Nor	1.0526	0.371	0.790	RU Sct	1.2945	0.966	0.690
AB Cam	0.7625	0.700	0.494	VY Cyg	0.8953	0.659	0.558	Y Oph	1.2336	0.671	0.681	SS Sct	0.5648	0.381	0.573
AD Cam	1.0516	0.952	0.618	VZ Cyg	0.6870	0.292	0.585	BF Oph	0.6094	0.289	0.591	UZ Sct	1.1686	0.997	0.897
OX Cam	0.7050	0.750	0.430	BZ Cyg	1.0061	0.886	0.713	RS Ori	0.8789	0.436	0.511	EV Sct	0.4901	0.710	0.447
RY CMa	0.6701	0.269	0.577	CD Cyg	1.2323	0.502	0.766	GQ Ori	0.9353	0.302	0.673	V367 Sct	0.7989	1.269	0.589
TW CMa	0.8448	0.340	0.631	DT Cyg	0.3978	0.060	0.480	SV Per	1.0465	0.397	0.619	ST Tau	0.6058	0.335	0.514
VZ CMa	0.6484	0.560	0.420	MW Cyg	0.7749	0.711	0.631	UX Per	0.6595	0.545	0.499	SZ Tau	0.6515	0.308	0.539
XZ CMa	0.4079	0.310	0.490	V386 Cyg	0.7208	0.905	0.588	UY Per	0.7296	0.902	0.612	AE Tau	0.5907	0.598	0.519
YZ CMa	0.4992	0.560	0.560	V402 Cyg	0.6400	0.393	0.612	VX Per	1.0370	0.481	0.671	EF Tau	0.5376	0.388	0.531
CN CMa	0.4059	0.630	0.510	V532 Cyg	0.5164	0.564	0.470	AS Per	0.6966	0.704	0.596	EU Tau	0.3227	0.261	0.421
U Car	1.5885	0.295	0.887	V924 Cyg	0.7460	0.264	0.578	AW Per	0.8105	0.540	0.514	α UMi	0.5988	0.020	0.328
SX Car	0.6866	0.200	0.727	V1154 Cyg	0.6925	0.265	0.648	BM Per	1.3608	0.987	0.803	T Vel	0.6665	0.331	0.597
VY Car	1.2773	0.292	0.869	V1334 Cyg	0.5228	0.101	0.406	HQ Per	0.9364	0.566	0.664	RY Vel	1.4493	0.609	0.770
WZ Car	1.3620	0.270	0.886	V1726 Cyg	0.6270	0.358	0.554	HZ Per	1.0523	1.360	0.767	RZ Vel	1.3096	0.331	0.790
GT Car	1.1193	0.660	0.796	TX Del	0.7900	0.253	0.483	MM Per	0.6147	0.505	0.590	SW Vel	1.3700	0.390	0.761
ℓ Car	1.5507	0.161	1.101	β Dor	0.9931	0.067	0.733	OT Per	1.4165	1.390	0.867	SX Vel	0.9800	0.308	0.580
RW Cas	1.1701	0.441	0.760	W Gem	0.8984	0.306	0.607	V440 Per	0.8791	0.313	0.561	CS Vel	0.7712	0.754	0.600
RY Cas	1.0841	0.602	0.769	AD Gem	0.5784	0.197	0.497	X Pup	1.4143	0.451	0.776	S Vul	1.8355	0.999	0.891
SU Cas	0.4401	0.293	0.408	DX Gem	0.4966	0.531	0.408	RS Pup	1.6169	0.519	0.918	T Vul	0.6469	0.092	0.544
SW Cas	0.7357	0.541	0.539	ζ Gem	1.0065	0.046	0.751	VZ Pup	1.3649	0.505	0.656	U Vul	0.9026	0.650	0.626
SY Cas	0.8097	0.474	0.501	V Lac	0.6975	0.392	0.481	AD Pup	1.1334	0.253	0.810	X Vul	0.8007	0.792	0.598
SZ Cas	1.1346	0.961	0.522	X Lac	0.7360	0.375	0.525	AQ Pup	1.4769	0.516	0.839	SV Vul	1.6533	0.489	0.960
TU Cas	0.3303	0.119	0.497	Y Lac	0.6359	0.205	0.526	BD Pup	0.5931	0.670	0.565	HD 18391	2.2500	1.011	0.955
XY Cas	0.6534	0.518	0.625	Z Lac	1.0369	0.411	0.686	BE Pup	0.4581	0.640	0.470				
BD Cas	0.5624	1.017	0.557	RR Lac	0.8073	0.365	0.519	BN Pup	1.1359	0.457	0.713				

A useful extension of the present study would be to include light amplitude as an additional parameter with the Table 4 data plotted in Figure 11. Light amplitude has previously been demonstrated to be related to a Cepheid's location in the instability strip, as inferred from its rate of stellar evolution (see Turner et al. 2006a). An additional mapping of light amplitude within the strip using unreddened colors would help to confirm such findings. But that

is left for a later investigation. The results of the present study are most usefully summarized by the compilations of Tables 2 and 4.

As noted earlier, several investigations of Cepheid field reddenings through membership in open clusters are as yet incomplete. It might also be possible to augment the sample of Cepheid reddenings by applying the transformation relations of Table 3 to other Cepheids in the cited studies that

are not coincident with the main sample of this paper. A larger sample of semi-empirical reddenings for classical Cepheids is therefore within relatively easy reach.

The present study owes much to the efforts of many different research collaborators over the author's lifetime, for which he is eternally grateful. The referee, Armando Arellano Ferro, also did much to inspire a thorough reworking of the original results to place them on sounder footing and make them more useful to a broader community. The study is much better for that input.

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