RevMexAA (Serie de Conferencias), **36**, CD291–CD296 (2009)

MAGNETIC-DRIVEN ORBITAL EVOLUTION OF AN ACCRETING MILLISECOND PULSAR: WITNESSING THE BANQUET OF A HIDDEN BLACK WIDOW

L. Burderi,[1] T. Di Salvo,[2] A. Riggio,[1] A. Papitto,[3,4] and M. T. Menna[4]

RESUMEN

Se reporta sobre la evolución orbital del pulsar de milisegundos SAX J1808.4–3658. En particular, encontramos para esta fuente la primera estimación de la derivada del periodo orbital en un pulsar acretante de milisegundos, $\dot{P}_{\rm orb} = (3.40 \pm 0.12) \times 10^{-12}$ s/s, y una estimación más refinada del periodo orbital, $P_{\rm orb} = 7249.156499 \pm (1.2 \times 10^{-5})$ s. Esta derivada es positiva y es más alta por un orden de magnitud a la esperada por una evolución secular por pérdidas de momento angular causadas por radiación gravitacional bajo la hipótesis de conservación de transferencia de masa. Bajo la hipótesis de que la derivada medida del periodo orbital refleja la evolución secular del sistema, proponemos una explicación simple de este resultado suponiendo que durante los periodos de emisión de rayos X la fuente eyecta materia (y momento angular) desde el punto interno de Lagrange. La evolución orbital propuesta del sistema sugiere la existencia de una estrella compañera degenerada o completamente convectiva e indica que esta clase de fuentes son capaces de erosionar con efectividad a las estrellas compañeras, y por eso son viudas negras visibles en rayos X durante los episodios de acreción de masa transitorio.

ABSTRACT

We report here on the orbital evolution of the accreting millisecond pulsar SAX J1808.4–3658. In particular, we find for this source the first estimate of the orbital period derivative in an accreting millisecond pulsar, $\dot{P}_{\rm orb} = (3.40 \pm 0.12) \times 10^{-12}$ s/s, and a refined estimate of the orbital period, $P_{\rm orb} = 7249.156499 \pm (1.2 \times 10^{-5})$ s. This derivative is positive and is more than one order of magnitude higher than what is expected from secular evolution driven by angular momentum losses caused by gravitational radiation under the hypothesis of conservative mass transfer. In the hypothesis that the measured derivative of the orbital period reflects the secular evolution of the system, we propose a simple explanation of this puzzling result assuming that during X-ray quiescence the source is ejecting matter (and angular momentum) from the inner Lagrangian point. The proposed orbital evolution of the system suggests a degenerate or fully convective companion star and indicates that this kind of sources are capable to efficiently ablate the companion star, and therefore are black widows visible in X-rays during transient mass accretion episodes.

Key Words: stars: neutron stars — stars: magnetic fields — pulsars: general — pulsars: individual (SAX J1808.4–3658) — X-ray: binaries

1. INTRODUCTION

SAX J1808.4–3658 is the first one discovered among the eight known accreting millisecond pulsars (hereafter AMSPs), all of them transient X-ray sources, and is still the richest laboratory for timing studies of these sources. Although timing analysis have been now performed on most of the sources of this sample with interesting results (see Di Salvo et al. 2008 for a review and references therein), SAX J1808.4–3658 is the only known AMSP for which more than one outburst has been observed by the RXTE/PCA with high time resolution. In particular, the first outburst of this source observed by the RXTE/PCA was in April 1998, when coherent X-ray pulsations at ~ 2.5 ms and orbital period of 2 hrs (Wijnands & van der Klis 1998; Chakrabarty & Morgan 1998) were discovered. The source showed other X-ray outbursts in 2000 (when only the final part of the outburst could be observed, Wijnands et al. 2001), in 2002 (when kHz QPOs and quasi-coherent oscillations during type-I X-ray bursts were discovered, Wijnands et al. 2003; Chakrabarty et al.

[1]Università degli Studi di Cagliari, Dipartimento di Fisica, SP Monserrato-Sestu, KM 0.7, 09042 Monserrato, Italy.
[2]Dipartimento di Scienze Fisiche ed Astronomiche, Università di Palermo, via Archirafi 36 - 90123 Palermo, Italy.
[3]Dipartimento di Fisica, Università degli Studi di Roma 'Tor Vergata', via della Ricerca Scientifica 1, 00133 Roma, Italy.
[4]Osservatorio Astronomico di Roma, Via Frascati 33, 00040 Monteporzio Catone (Roma), Italy.

2003), and again in 2005, approximately every two years (see Wijnands 2005 for a review).

Papitto et al. (2005) performed a temporal analysis of the outbursts of SAX J1808.4–3658 that occurred in 1998, 2000, and 2002, which resulted in improved orbital parameters of the system. The large uncertainty caused by the relatively limited temporal baseline made it impossible to derive an estimate of the orbital period derivative. In this paper we use all the four outbursts of SAX J1808.4–3658 observed by RXTE/PCA, that means a temporal baseline of more than 7 years, to derive an orbital period derivative, the first reported to date for an accreting millisecond pulsar. The value we find with high statistical significance is surprising, $(3.40 \pm 0.12) \times 10^{-12}$ s/s. Orbital evolution calculations in the hypothesis of conservative mass transfer show that the orbital period derivative caused by emission of Gravitational Radiation (GR) is given by $\dot{P}_{\rm orb} = -1.4 \times 10^{-13}\ m_1\ m_{2,\,0.1}\ m^{-1/3} P_{2h}^{-5/3}\ [(n - 1/3)/(n + 5/3 - 2q)]$ s/s (derived from Verbunt 1993; see also Rappaport et al. 1987), where m_1 and m are, respectively, the mass of the primary and the total mass in units of solar masses, $m_{2,\,0.1}$ is the mass of the secondary in units of 0.1 $M_\odot$, P_{2h} is the orbital period in units 2 h, $q = m_2/m_1$ is the mass ratio and where n is the index of the mass-radius relation of the secondary $R_2 \propto M_2^{\rm n}$. Therefore a positive orbital period derivative certainly indicates a mass-radius index $n < 1/3$, and therefore, most probably, a degenerate or fully convective companion star (see e.g. King 1988). However, the $\dot{P}_{\rm orb}$ we measure is an order of magnitude higher than what is expected from GR. In § 3 we propose a simple explanation of this result assuming that it reflects the secular evolution of the system. In particular, we hypothesize that during quiescence SAX J1808.4–3658 experiences a highly non-conservative mass outflow, in which a great quantity of mass is lost from the system with a relatively low specific angular momentum.

Indeed, SAX J1808.4–3658 and the other known AMSPs are all transient systems (accreting just for a very short fraction of the time), with small values of the mass function (implying small minimum mass for the secondary) and short orbital periods (less than a few hours). As predicted by King et al. (2005), these pulsars may be part of the population of the so-called black widows which are millisecond pulsars in binary systems believed to evaporate their companion star. Although in some of these systems radio eclipses have been observed, clearly demonstrating the presence of matter around the system, a direct proof of severe mass losses from these system has never been found to date. The orbital evolution of SAX J1808.4–3658 indicates that this X-ray transient millisecond pulsar indeed may expel mass from the system for most of the time with just short episodes of accretion observed as X-ray outbursts. We therefore propose that SAX J1808.4–3658 (and perhaps most AMSPs) is indeed black widow still eating the companion star.

2. TIMING ANALYSIS

In this paper we analyse RXTE public archive data of SAX J1808.4–3658 taken during the April 1998 (Obs. ID P30411), the February 2000 (Obs. ID P40035), the October 2002 (Obs. ID P70080), and the June 2005 (Obs. ID P91056 and Obs. ID P91418) outbursts, respectively. In particular, we analysed data from the PCA (Jahoda et al. 1996), which is composed of a set of five xenon proportional counters operating in the $2 - 60$ keV energy range with a total effective area of 6000 cm^2. For the timing analysis, we used event mode data with 64 energy channels and a 122 μs temporal resolution; the arrival times of all the events were referred to the solar system barycentre by using JPL DE-405 ephemerides along with spacecraft ephemerides. This task was performed with the faxbary tool, considering as the best estimate for the source coordinates the radio counterpart position, that has a 90% confidence radius of 0.4 arcsec, which is compatible with that of the optical counterpart (Rupen et al. 2002; Giles et al. 1999).

For each of the outbursts we derived a precise orbital solution using standard techniques (see e.g. Burderi et al. 2007; Papitto et al. 2007, and references therein). In particular, we firstly corrected the arrival times of all the events with the orbital solution given by Papitto et al. (2005). Then we looked for differential corrections to the adopted orbital parameters as described in the following. We epoch-folded time intervals each with a duration of about 720 s (1/10 of the orbital period) at the spin period of 2.49391975936 ms, and fitted each pulse profile obtained in this way with a sinusoid in order to derive the pulse arrival times or pulse phases. Note that the determination of the pulse phases is insensitive to the exact value of the spin period chosen to fold the light curves providing that this value is not very far from the true spin period of the pulsar. In order to choose a value of the spin period as close as possible to the true one we used the value above, that is in between the best-fit spin period reported by Chakrabarty & Morgan (1998) for the

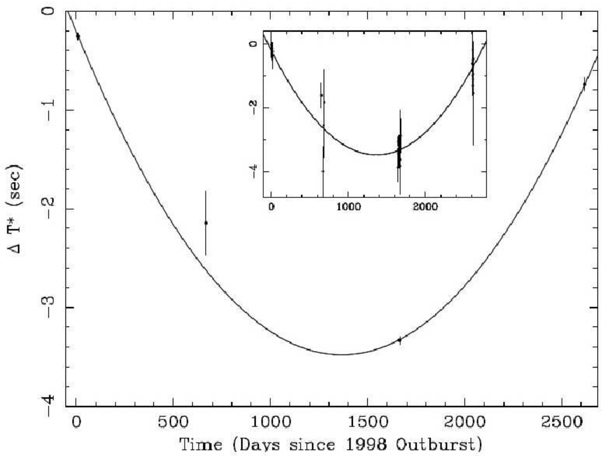

Fig. 1. Differential correction, ΔT^*, to the time of passage of the NS at the ascending node for each of the four outbursts analysed. In the inset we show the single measurements of ΔT^* obtained for each of the consecutive time intervals in which each outburst has been divided (see text). All the times are computed with respect to the beginning of the 1998 outburst, $T_0 = 50914.8099$ MJD.

1998 outburst and the best-fit spin period reported by Burderi et al. (2006) for the 2002 outburst.

We then looked for differential corrections to the adopted orbital parameters, which can be done by fitting the pulse phases as a function of time for each outburst. In general, any residual orbital modulation is superposed to a long-term variation of the phases, e.g. caused by a variation of the spin. However, as noted by Burderi et al. (2006) for the 2002 outburst, SAX J1808.4–3658 shows a very complex behaviour of the pulse phases with time, with phase shifts probably caused by variations of the pulse shape that are difficult to model and to interpret. To avoid any fitting of this complex long-term variation of the phases, we preferred to restrict the fitting of the differential corrections to the orbital parameters to intervals in which the long-term variation and/or shifts of phases is negligible. We therefore considered consecutive intervals with a duration of at least 4 orbital periods (depending on the statistics), and fitted the phases of each of these intervals with the formula for the differential corrections to the orbital parameters (see e.g. Deeter, Pravdo, & Boynton 1981 and equation (3) in Papitto et al. 2007). No significant corrections were found on the adopted values of the orbital period, $P_{\rm orb}$, and the the projected semi-major axis of the neutron star (NS) orbit, $a_1 \sin i/c$.

On the other hand we found that the times of passage of the NS at the ascending node at the beginning of each outburst, T_N^*, were significantly dif-

TABLE 1

BEST FIT ORBITAL PARAMETERS FOR SAX J1808.4–3658

Parameter	Value
T_0^* (days)	50914.8784320(7)
$P_{\rm orb}$ (s)	7249.15650(1)
$\dot{P}_{\rm orb}$ (s/s)	$3.4(1) \times 10^{-12}$
χ^2/dof	2.2/1

The reference time at which the orbital period, $P_{\rm orb}$ and its derivative, $\dot{P}_{\rm orb}$, are referred to is the beginning of the 1998 outburst, that is $T_0 = 50914.8099$ MJD. Numbers in parentheses are the uncertainties in the last significant digit at 90% c.l.

ferent from their predicted values, $T_0^* + NP_{\rm orb}$ where T_0^* is the adopted time of ascending node passage at the beginning of the 1998 outburst and the integer N is the exact number of orbital cycles elapsed between two different ascending node passages, i.e. N is the integer part of $(T_N^* - T_0^*)/P_{\rm orb}$ under the assumption that $|T_N^* - T_0^*| < P_{\rm orb}$ that we have also verified *a posteriori*. We therefore fixed the values of $P_{\rm orb}$ and $a_1 \sin i/c$ and derived the differential corrections, ΔT^*, to the time of passage at the ascending node, obtaining a cluster of points for each outburst, which are plotted as a function of time in the inset of Figure 1. Since each of the points corresponding to an outburst may be considered like an independent estimate of the same quantity, we computed the weighted mean of the points corresponding to each outburst, obtaining the four points shown in Figure 1.

These points show a clear parabolic trend that we fitted to the formula:

$$\Delta T^* = \delta T_0^* + \delta P_{\rm orb} \times N + (1/2)\dot{P}_{\rm orb} P_{\rm orb} \times N^2 \,, \quad (1)$$

In this way we found the best fit values $T_0^* + \delta T_0^*$, $P_{\rm orb} + \delta P_{\rm orb}$ and $\dot{P}_{\rm orb}$ at $t = T_0$ shown in Table 1, with a $\chi^2 = 2.2$ (for 1 d.o.f.). This corresponds to a probability of 13% of obtaining a χ^2 larger than the χ^2 we found. Our result is therefore acceptable at more than 5% c.l., which is the standard confidence level to accept a model based on χ^2 (see e.g. Bevington & Robinson 2003). As a further check we fit with the same formula the points shown in the inset of Figure 1, obtaining, as expected, the same results.

Note that the orbital period of SAX J1808.4–3658 is now known with a precision of 1 over 10^9, that is one order of magnitude improvement with respect to the previous estimate by Papitto et al. (2005) and two orders of magnitude with respect to Chakrabarty & Morgan (1998). We also find a highly statistically significant derivative of the orbital period, which indicates that the orbital period in this system is increasing at a rate of $(3.40 \pm 0.12) \times 10^{-12}$ s/s, that is at least an order of magnitude higher than what is predicted by a conservative mass transfer driven by GR. In the next section we discuss a possible explanation of this result.

3. DISCUSSION AND CONCLUSIONS

We have performed a precise timing analysis of all the X-ray outbursts of the AMSP SAX J1808.4–3658 observed to date by the RXTE/PCA, covering more than 7 years in time. We divided each outburst in several intervals and found for each interval an orbital solution as differential corrections to previously published orbital parameters. The obtained times of passage of the NS at the ascending node were significantly different in different outbursts. We fitted these times with a parabolic function of time finding an improved orbital solution valid over more than 7 years time span. This solution includes a highly significant derivative of the orbital period, $\dot{P}_{\rm orb} = (3.40 \pm 0.12) \times 10^{-12}$ s/s. This is the first measure of the orbital period derivative for an accreting millisecond pulsar. However, this orbital period derivative is quite unexpected, since it is more than one order of magnitude higher than what is expected from conservative mass transfer driven by GR.

In order to compare the measured orbital period derivative with what is expected from the secular evolution of the system driven by GR, we have calculated the orbital period derivative and the mass transferred by the secondary star under the following assumptions. (i) Angular momentum losses are due to GR and are given by: $\dot{J}/J = -32G^3 M_1 M_2 (M_1 + M_2)/(5c^5 a^4)$, where G is the Gravitational constant, M_1 and M_2 the masses of the primary and of the secondary, respectively, c the speed of light, a the orbital separation, and where $J = M_1 M_2 [Ga/(M_1 + M_2)]^{1/2}$ is the binary angular momentum (see e.g. Landau & Lifschitz 1958; Verbunt 1993). (ii) For the secondary we have adopted a mass-radius relation $R_2 \propto M_2^n$. (iii) We have imposed that the radius of the secondary follows the evolution of the secondary Roche Lobe radius: $\dot{R}_{L2}/R_{L2} = \dot{R}_2/R_2$, where for the radius of the secondary Roche Lobe we have adopted the Paczynski (1971) approximation: $R_{L2} = 2/3^{4/3}[q/(1+q)]^{1/3} a$ that is valid for small mass ratios, $q = M_2/M_1 \leq 0.8$.

In these hypotheses we derived a simple expression for the orbital period derivative and the mass transfer rate in the extreme cases of totally conservative and totally non-conservative mass transfer (see e.g. Verbunt 1993; van Teeseling & King 1998; King et al. 2003; King et al. 2005):

$$\begin{aligned} \dot{P}_{\rm orb} &= -1.38 \times 10^{-12} \left[\frac{n - 1/3}{n + 5/3 - 2g}\right] \\ &\times\ m_1^{5/3} q (1+q)^{-1/3} P_{2h}^{-5/3} \ \text{s/s}\,, \qquad (2) \\ \dot{M} &= 4.0 \times 10^{-9} \left[\frac{1}{n + 5/3 - 2g}\right] \\ &\times\ m_1^{8/3} q^2 (1+q)^{-1/3} P_{2h}^{-8/3} \ M_\odot/\text{yr}\,, \ (3) \end{aligned}$$

where $g = q$ for totally conservative mass transfer, and $g = (\alpha + q/3)/(1 + q)$ for totally non-conservative mass transfer, and where $\alpha = l P_{\rm orb} (M_1 + M_2)^2/(2\pi a^2 M_1^2)$ is the specific angular momentum of the ejected matter, l, in units of the specific angular momentum of the secondary.

Comparing the measured orbital period derivative with equation 3 (assuming that the orbital period derivative we measure reflects the secular evolution of the system), we note that, in order to have $\dot{P}_{\rm orb} > 0$ we have to assume an index $n < 1/3$. In the case of SAX J1808.4–3658 the secondary mass is $m_2 \leq 0.14$ at 95% c.l. (Chakrabarty & Morgan 1998), and therefore the mass ratio $q \leq 0.1$ implies that for the totally conservative mass transfer case $(g = q)$, the $\dot{P}_{\rm orb}$ expected from GR must be of the order of 10^{-13} s/s, not compatible with what we measure for SAX J1808.4–3658. In other words, if we assume a conservative mass transfer for the system (that means that the mass transferred during outbursts is completely accreted by the NS, and during quiescence no or negligible mass is lost from the system), than it is impossible to explain the observed orbital period derivative with an evolution driven by GR.

On the other hand, if we assume that during X-ray quiescence the companion star is still overflowing its Roche Lobe but the transferred mass is not accreted onto the NS and is instead ejected from the system, we find a good agreement between the measured and expected orbital period derivative assuming that the matter leaves the system with the specific angular momentum at the inner Lagrangian point, $\alpha = [1 - (2/3^{4/3}) q^{1/3} (1+q)^{2/3}]^2$. Adopting the measured value $\dot{P}_{\rm orb} = 3.4 \times 10^{-12}$ s/s and the other parameters of SAX J1808.4–3658, equation 3 translates into a relation between m_1, m_2 and the mass-radius index n; this is plotted in Figure 2 (left panel) for different values of n going from 0 to $n = -1/3$. The constraint on m_1 vs. m_2 imposed by the mass function of the system is also plotted (the

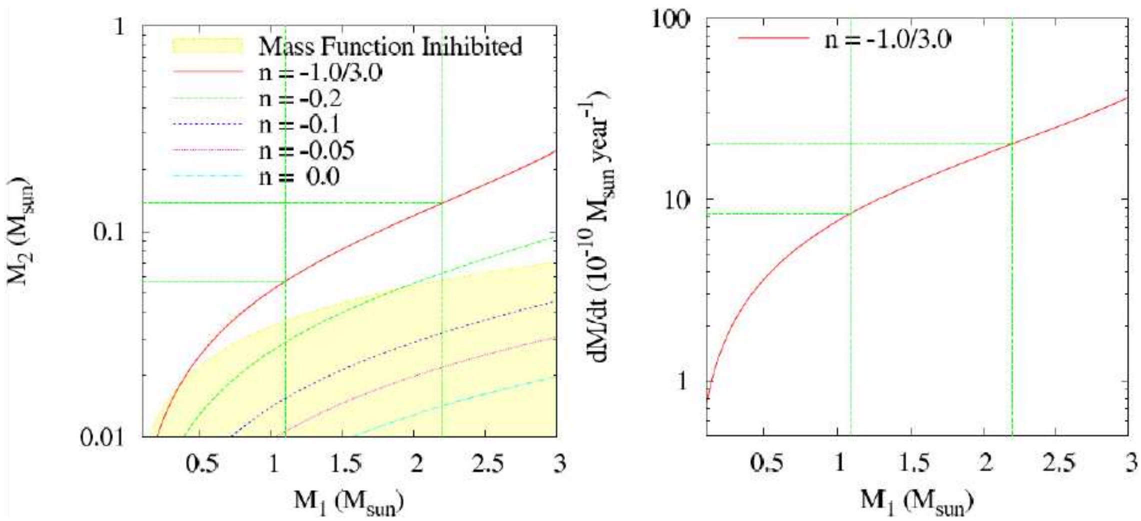

Fig. 2. Left: Companion star mass vs. NS mass in the hypothesis of totally non conservative mass transfer (with matter leaving the system with the specific angular momentum at the inner Lagrangian point) and assuming the $\dot{P}_{\rm orb}$ measured for SAX J1808.4–3658. Different curves correspond to different values of the mass-radius index n of the secondary. Horizontal lines indicate the limits for the secondary star mass corresponding to reasonable limits for the NS mass and to $n = -1/3$. Right: Mass rate outflowing the secondary Roche Lobe in the hypothesis of totally non conservative mass transfer (as above) and assuming $n = -1/3$.

shadowed region in the figure) and indicates that the most probable value for n is $-1/3$, which in turn indicates a degenerate or, most probably, a fully convective companion star.

In fact, in a system with orbital period less than 3 h, where the mass of the Roche-Lobe filling companion is below $0.2 - 0.3\ M_\odot$, the companion star becomes fully convective with a mass-radius hydrostatic equilibrium equation $R \propto M^{-1/3}$ (e.g. King 1988; Verbunt 1993). Also, for reasonable minimum, average and maximum values of the NS mass, 1.1, 1.56, and 2.2 $M_\odot$, respectively, we obtain the following values for the secondary mass: 0.053, 0.088, and 0.137 $M_\odot$, and the following values for the inclination of the system: 44°, 32°, and 26°, respectively.

Assuming therefore $n = -1/3$ we have plotted in Figure 2 (right panel) the corresponding non-conservative mass transfer rate as a function of m_1. We find that for $m_1 = 1.5$ the mass transfer rate must be of the order of $10^{-9}\ M_\odot$/yr, much higher than what is expected in a conservative GR driven mass transfer case. Note that in Burderi et al. (2006), the proposed interpretation of the complex observed behaviour of the pulse phase evolution during the 2002 outburst implied a mass accretion rate of $1.8\times10^{-9}\ M_\odot$/yr at the beginning of the outburst; we note that this value is in agreement with the value of $\dot{M}$ necessary to explain the observed orbital period derivative in the hypothesis of a strongly non-conservative mass transfer during quiescence. Actually, during the X-ray outbursts, the mass transfer is conservative since the transferred matter is accreted onto the NS. However, the accretion phase duty cycle, about 40 days / 2 years = 5%, is so small that the totally non-conservative scenario proposed above is substantially unchanged.

The question to answer is why the accretion is inhibited during X-ray quiescence while the companion star is transferring mass at a high rate? We propose that the answer has to be find in the radiation pressure of the magneto-dipole rotator emission, with a mechanism that is similar to what is proposed to explain the behaviour of the so-called black widow pulsars (see e.g. King et al. 2003, 2005; see also Burderi et al. 2001; Burderi, D'Antona, & Burgay 2002). Indeed, the possibility that the magneto-dipole emission is active in SAX J1808.4–3658 during X-ray quiescence has been used by Burderi et al. (2003) to explain the optical counterpart of the source, which is observed to be overluminous during quiescence. In conclusion, we propose that SAX J1808.4–3658 during quiescence behaves exactly as a black widow pulsar, efficiently ablating its companion star. When (or if) the pressure of the outflowing matter becomes sufficiently high to temporarily overcome the radiation pressure of the magneto-dipole rotator, the source experiences a transient mass accretion episode, resulting in an X-ray outburst. We propose therefore that we are witnessing the behaviour of a hidden black widow eating its companion during X-ray outbursts and ablating it during quiescence.

This work was supported by the Ministero della Istruzione, della Università e della Ricerca (MIUR), national program PRIN2005 2005024090_004.

REFERENCES

Bevington, P. R., & Robinson, D. K. 2003, Data Reduction and Error Analysis for the Physical Sciences (3rd ed.; Boston: McGraw-Hill)
Burderi, L., D'Antona, F., & Burgay, M. 2002, ApJ, 574, 325
Burderi, L., Di Salvo, T., D'Antona, F., Robba, N. R., & Testa, V. 2003, A&A, 404, L43
Burderi, L., Di Salvo, T., Menna, M. T., Riggio, A., & Papitto, A. 2006, ApJ, 653, L133
Burderi, L., et al. 2001, ApJ, 560, L71
Burderi, L., et al. 2007, ApJ, 657, 961
Campana, S., et al. 2002, ApJ, 575, L15
Chakrabarty, D., & Morgan, E. H. 1998, Nature, 394, 346
Chakrabarty, D., Morgan, E. H., Muno, M. P., Galloway, D. K., Wijnands, R., van der Klis, M., & Markwardt, C. B. 2003, Nature, 424, 42
Di Salvo, T., Burderi, L., Riggio, A., Papitto A., Menna M. T. 2008, in AIP Conf. Proc. 1054, Cool Discs, Hhot Flows: The Varying Faces of Accreting Compact Objects, ed. M. Axelsson (Melville: AIP), 173
Deeter, J. E., Boynton, P. E., & Pravdo, S. H. 1981, ApJ, 247, 1003
Giles, A. B., Hill, K. M., & Greenhill, J. G. 1999, MNRAS, 304, 47
Jahoda, K., et al. 1996, Proc. SPIE, 2808, 59
King, A. R. 1988, QJRAS, 29, 1
King, A. R., Beer, M. E., Rolfe, D. J., Schenker, K., & Skipp, J. M. 2005, MNRAS, 358, 1501
King, A. R., Davies, M. B., & Beer, M. E. 2003, MNRAS, 345, 678
Landau, L. D., & Lifshitz, E. M. 1975, The Classical Theory of Fields (4th ed.; Oxford: Pergamon Press)
Paczynski, B. 1971, ARA&A, 9, 183
Papitto, A., Di Salvo, T., Burderi, L., Menna, M. T., Lavagetto, G., & Riggio, A. 2007, MNRAS, 375, 971
Papitto, A., Menna, M. T., Burderi, L., Di Salvo, T., D'Antona, F., & Robba, N. R. 2005, ApJ, 621, L113
Rappaport, S., Ma, C. P., Joss, P. C., & Nelson, L. A. 1987, ApJ, 322, 842
Rupen, M. P., Dhawan, V., Mioduszewski, A. J., Stappers, B. W., & Gaensler, B. M. 2002, IAU Circ. 7997, 2
van Teeseling, A., & King, A. R. 1998, A&A, 338, 957
Verbunt, F. 1993, ARA&A, 31, 93
Wijnands, R. 2005, in Pulsars New Research (New York: Nova Science Publishers)
Wijnands, R., Mendez, M., Markwardt, C., van der Klis, M., Chakrabarty, D., & Morgan, E. 2001, ApJ, 560, 892
Wijnands, R., & van der Klis, M. 1998, Nature, 394, 344
Wijnands, R., van der Klis, M., Homan, J., Chakrabarty, D., Markwardt, C. B., & Morgan, E. H. 2003, Nature, 424, 44