(z = 0.0178). GALEX FUV and NUV images of this group reveal that there are 15 sources located in the vicinity of the intergalactic HI clouds of the compact group which extend to over 130 kpc away from the main galaxies (de Mello et al. 2008). Figure 1 shows the Gemini images taking during the mask preparation of two of the UV sources. A tidal dwarf galaxy (#4 in Figure 1) is located in the densest region of the HI tail, ∼61 kpc from the brightest group member. A second candidate (#3 in Figure 1) is located at ∼120 kpc away from the main galaxies. They resemble dwarf irregulars and are ∼7–8 kpc in diameter. Object #4 contains clumps of ∼200 pc which we interpreted as large star-forming regions. For comparison, the nearest massive star-forming region in our galaxy, the Carina nebula (NGC 3372), and the very large OB association, NGC 604 in M33 have diameters ∼150 pc (Zinnecker & Yorke 2007) and ∼140 pc (Maiz-Apellaniz et al. 2004), respectively. The Gemini spectra of both objects (Urrutia-Viscarra et al. 2011) indicate that they have metallicities higher than local dwarf galaxies suggesting that these star-forming regions are formed from pre-enriched material that was ejected from the host galaxies into the intragroup medium.

We have also searched for TDGs outside the merger remnant, NGC 2782 (34 Mpc). This galaxy shows a prominent tidal tail detected in HI, located to the western side of the object. By using GALEX FUV and NUV imaging and HI data we found seven UV sources located at distances greater than 26 kpc from the center of NGC 2782, and coinciding with its western HI tidal tail (Torres-Flores et al. 2011). These objects have ages of 8 to 188 Myr and masses ranging from $10^{5.2}$ to $10^{6.8}$ $M_{\odot}$. Using GEMINI/GMOS r-band images we resolved each UV source in several smaller systems showing that they are not as large as TDGs but resemble small stellar clusters. Using Gemini/GMOS spectroscopic data we confirm that they are at the same distance as NGC 2782. Two of them have super solar metallicities which could be explained if they were formed out of highly enriched gas which was once expelled from the center of the merging galaxies when the system collided, whereas an inflow of low abundance gas from the outskirts of the merging galaxies could explain the comparably lower metallicity observed close to the central region of NGC 2782. An additional possibility is that the tail has been a nursery of a few generations of young stellar systems which ultimately polluted this medium with metals, further enriching the already pre-enriched gas ejected to the tail when the galaxies collided.

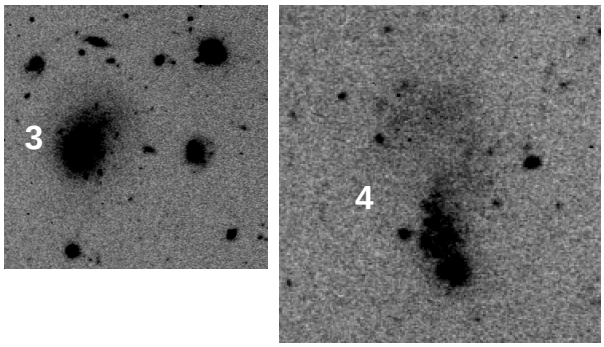

Fig. 1. Gemini/GMOS Image of the two new TDGs in the HCG100 HI tail.

The two systems we presented here host star formation outside the main galaxies. The less evolved system, HCG100, was able to form two tidal dwarf galaxies whereas the merger remnant formed less massive stellar clusters. The fate of these systems is unknown. However, they seem to be commonly found outside interacting systems. Small mass objects forming stars will be efficient pollutant of the intergalactic medium on large scales (Ferrara 2008). Therefore, these systems may play an important role in the cosmic metal enrichment. Moreover, the high-redshift Lyman Break Galaxies resemble small blue compact dwarf galaxies (Filkenstein et al. 2010) and TDGs might be their counterparts in the local universe.

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