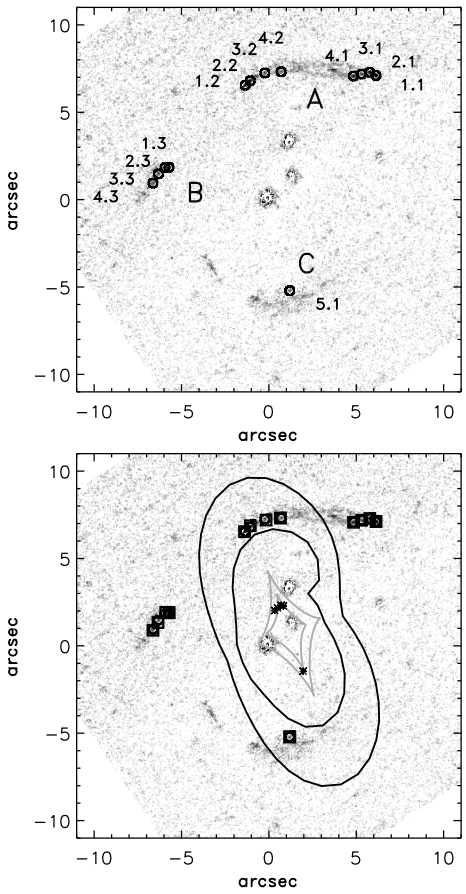


Fig. 1. Top panel. HST/ACS image (F475W) after the subtraction of the central galaxies of the group. The circles shows the positions of the conjugated images that constitutes the arc system AB and the location of arc C. Bottom panel. Tangential critical lines (black) and caustic lines (gray) for two different sources located at $z = 1.0$ and $z = 1.7$. The circles shows the measured positions of the images (input data for the model). The squares are the model-predicted image positions after optimization and the asterisks the positions in the source plane.

We use the dynamical information of the group in order to constrain the scale radius of the NFW profile (see the discussion in Verdugo et al. 2010). In the bottom panel of Figure 1 we show the model-predicted image positions after optimization as well as the observed positions. The remarkable agreement between both positions is quantified by a mean scatter in the image plane smaller than $0.04''$ and a $\chi^2_{\rm DOF} = 1.2$

## 4. COMPARING MASSES

The weak lensing analysis of this group was presented in the work of Limousin et al. (2009) and the dynamical analysis was done using spectroscopic

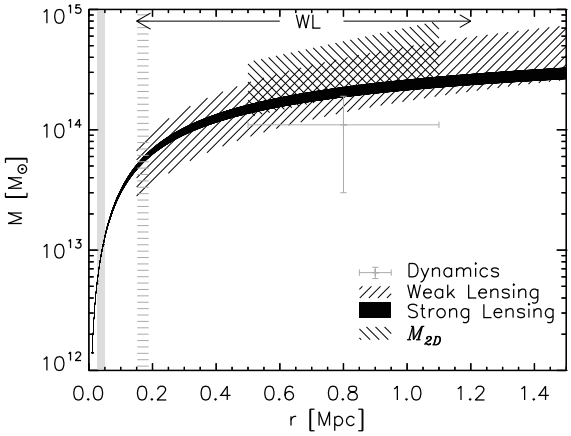


Fig. 2. The 2D projected mass as a function of the aperture radius measured from the center of the group. The $M_{\rm 2D}$ depict the mass of the profile calculated analytically. The point with error bars is the mass estimate from dynamics. The gray-shaded region shows the area where the arc systems lie and the gray region filled with horizontal lines shows the model-predicted scale radius.

data form VLT/FORS2 and KECK/LRIS (see Verdugo et al. 2010). In Figure 2 we compare the masses obtained with different methods: strong lensing, weak lensing, $M_{\rm 2D}$ derived analytically from the profile, and the virial mass calculated from the dynamics. As we place limits on the possible values of the scale radius, the strong lensing mass is still reliable up to this radius ($r_s = 170 \pm 18$ kpc). From Figure 2 we can note that the masses are in agreement within the errors

## 5. CONCLUSIONS

We have probed the mass of SL2S J02140-0535 on a large radial range using complementary techniques: strong lensing, weak lensing, and dynamics. Our model reproduces the observed features in this system and we shows the possibility of constraining large scale properties that could not be measured by strong lensing, such as the scale radius in the NFW profile.

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