TABLE 7

COMPARISON OF DERIVED AND SOLAR ABUNDANCES

	HR 3273		HR 8118		HR 8567		HR 8937		Sun
Species	$\log N/N_H$	n	$\log N/N_H$	n	$\log N/N_H$	n	$\log N/N_H$	n	$\log N/N_H$
HeI	-1.12 ± 0.06	5	-1.01 ± 0.01	3	-1.04 ± 0.05	5	-1.47 ± 0.04	3	-1.01
CII	-3.61 ± 0.11	2	-3.10 ± 0.02	2	-3.82 ± 0.27	3	-3.24 ± 0.03	2	-3.45
OI	-3.23	1	-3.71	1	-3.09	1			-3.13
MgI	-5.17	1	-4.96 ± 0.02	2	-5.14	1	-5.01 ± 0.01	2	-4.42
MgII	-5.11 ± 0.03	2	-4.54 ± 0.10	3	-4.91 ± 0.03	2	-4.73 ± 0.15	3	-4.42
AlII	-6.80	1	-5.47	1	-6.52	1			-5.53
SiII	-4.92 ± 0.22	8	-4.61 ± 0.17	6	-4.78 ± 0.20	7	-4.52 ± 0.18	8	-4.45
PII	-5.97	1					-5.37	1	-6.55
SII	-4.99 ± 0.14	7	-4.45 ± 0.19	5	-4.73 ± 0.18	7	-5.09 ± 0.09	3	-4.67
CaII	-5.57	1	-6.18	1	-5.35	1	-5.93	1	-5.64
ScII	-8.03	1	-7.02	1	-7.62	1			-8.83
TiII	-6.35 ± 0.24	34	-6.11 ± 0.22	34	-5.97 ± 0.23	32	-6.60 ± 0.23	26	-6.98
CrI			-5.20	1					-6.33
CrII	-5.53 ± 0.23	37	-5.19 ± 0.19	23	-5.34 ± 0.21	32	-5.73 ± 0.25	30	-6.33
MnI	-4.63 ± 0.22	7	-5.54 ± 0.18	3	-4.20 ± 0.26	14	-4.67 ± 0.19	11	-6.61
MnII	-4.68 ± 0.21	39	-5.26 ± 0.32	12	-4.25 ± 0.23	41	-4.46 ± 0.21	39	-6.61
FeI	-4.34 ± 0.02	2	-4.27 ± 0.35	19	-4.78 ± 0.24	3	-4.30 ± 0.40	18	-4.50
FeII	-4.75 ± 0.21	64	-4.31 ± 0.19	63	-4.91 ± 0.21	56	-4.29 ± 0.20	80	-4.50
FeIII	-4.56	1					-4.35	1	-4.50
NiII	-6.74 ± 0.30	2	-6.06 ± 0.21	2	-6.79 ± 0.25	2			-5.75
SrII			-8.43	1	-8.35	1	-6.30	1	-9.03
YII	-7.68 ± 0.26	5	-6.32 ± 0.24	9	-7.40 ± 0.16	8	-6.37 ± 0.20	11	-9.76
ZrII	-7.79 ± 0.05	2	-7.06 ± 0.16	13	-7.59 ± 0.32	2	-7.72 ± 0.14	6	-9.40
BaII			-9.17	1					-9.87
HgII	-5.25	1	-5.03	1	-5.58	1	-4.23	1	-10.83
$T_{\rm eff}$	12253		11381		12315		12088		
$\log g$	3.42		4.11		3.68		4.28		

(1996). We also present in Tables 8–11 the line by line abundance values derived for the sample of HgMn stars. The columns shown are the code of the species, the name of the element, multiplet number, wavelength of the line, $\log gf$ value, reference for the $\log gf$ value, equivalent width and abundance of the line. In Figure 1 we show the abundance anomalies relative to solar as a function of the atomic number. The HgMn stars are presented as crosses (HR 3273), empty triangles (HR 8118), empty circles (HR 8567) and empty squares (HR 8937). For comparison purposes we show the abundances of four other HgMn stars, depicted as filled circles (HR 4817), filled triangles (π Boo), filled hexagons (μ Lep) and filled squares (28 Her). These data are taken from Adelman & Pintado (1997). From the values presented in Table 7 and Figure 1, our sample of HgMn stars shows an abundance pattern similar to other HgMn stars. However there are some differences in particular cases.

The star HR 3273 shows a relatively low abundance of SiII (8 lines) compared to other HgMn stars. In the star HR 8567, the abundance of CrII (32 lines) is relatively high compared to other HgMn stars. The CrII lines $\lambda\lambda$4588, 4634 and 5237, for instance, are intense in the spectra. The Mn and Hg abundances are similar to other HgMn stars. The abundance of CII in HR 8118 and HR 8937 is also relatively high. However, this value is based on few lines, and assumes the LTE approximation. For star HR 8937, the abundance of SrII is high compared to other HgMn stars. However, this value is derived