could, in principle, lead to a quadrupole-like wind distribution, independently of the eccentricity of the orbit. Demonstrating that a quadrupolar wind distribution in a rotating star reproduces the observed line-profile variations requires the use of a model that allows for departures from spherical symmetry.

In conclusion, we find that the N V 4604 line-profile variability may be understood in terms of a variable wind structure which involves changes in the N V opacity. The mechanism causing this variability remains to be determined.

GK acknowledges support from PAPIIT/Universidad Nacional Autónoma de México grant and computing support from Ulises Amaya. AF acknowledges the sabbatical stay in the INAOE and support from Conacyt under the target financed project 121166 "Spectroscopy Variations of Wolf-Rayet Stars", and by the UNACAR grant 1320.

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