TABLE 2

TIMES OF MID-ECLIPSES FOR V1898 CYG AND THE O-C RESIDUALS (SEE TEXT)

Minimum time (HJD-2400000)	Epoch	O-C(I)	O-C(II)	O-C(III)	Ref.
45960.6758	−3126	−0.0371	0 .0139	0.0004	1
45963.6986	−3124	−0.0406	0.0105	−0.0030	1
46010.6101	−3093	−0.0355	0.0151	0.0018	1
46013.6351	−3091	−0.0367	0.0138	0.0006	1
50690.6948	0	0.0000	0.0010	0.0010	3
52169.7772	977.5	0.0174	0.0027	0.0014	3
52185.6636	988	0.0161	0.0013	0.0000	3
52895.3220	1457	0.0259	0.0036	0.0007	2
52901.3740	1461	0.0255	0.0031	0.0002	2
52928.6107	1479	0.0262	0.0035	0.0005	3
53207.7802	1663.5	0.0269	0.0013	−0.0025	3
53226.6966	1676	0.0294	0.0036	−0.0003	3
53246.3663	1689	0.0287	0.0027	−0.0013	2
53270.5757	1705	0.0284	0.0021	−0.0020	3
54443.2559	2480	0.0483	0.0096	0.0011	4
54443.2565	2480	0.0489	0.0102	0.0017	4
54691.4070	2644	0.0494	0.0080	−0.0016	4
54691.4089	2644	0.0513	0.0099	0.0003	4
54691.4097	2644	0.0521	0.0107	0.0011	4

References: (1) Halbedel (1985), (2) Dallaporta & Munari (2006), (3) Caton & Smith (2005), (4) Brát et al. (2008).

elements O-C(II) with respect to the epoch numbers suggests an upward curved parabola. Therefore, a parabolic fit to the data was adopted, which gives,

$$\mathrm{MinI(HJD)} = 2\,450\,690.6938(8) + 1^d.5131260(2) \times E + 1.38(13)10^{-9} \times E^2 \,. \quad (2)$$

The standard mean errors in the last digits are given in parentheses. The coefficient of the quadratic term is positive, which indicates that the orbital period of the system is increasing with the epoch number. Such a quadratic ephemeris appears to a very good representation of the orbital period change of V1898 Cyg, as well as other interacting Algols. This quadratic behaviour of the O-C(II) residuals, plotted in the upper panel of Figure 3, is an indication of the secular period increase for the system. In the bottom panel of Figure 3 the O-C(III) residuals with respect to ephemeris (2) are plotted, which illustrates a good agreement between the timings and the new ephemeris. It is known that the classical Algols have an evolved less massive component which fills its Roche lobe and transfers its mass to the more massive primary star through the Lagrangian L_1 point of the system. The orbital period of the system is increasing at an average rate of $\dot{P}/P = 6.68(\pm0.63) \times 10^{-7}$ yr^{-1} which means that the orbital period has increased by about $0.38(\pm0.04)$ seconds in the last 24 years. The sum of the squares of residuals for the parabolic fit is 3.672×10^{-5} d^2. We limited times of mid-eclipses covering about 24 yrs. The time span of the observations is too short to reveal any abrupt period change caused by the fast mass transfer phenomenon. Of course, observations to be obtained in the coming years could indicate some hints about the nature of the orbital period change. Assuming a conservative mass transfer from the less massive component to the more massive primary star we estimate a transfer of $1.88(\pm0.17) \times 10^{-7}$ $M_\odot$ in a year.

5. ANALYSIS OF THE LIGHT CURVES

Three light curves V1898 Cyg have been obtained and published. The first photometric observations were obtained by Halbedel (1985) between July and November, 1985. The second photometric observations were made by CS from August 22, 2001 to Oc-