TABLE 1

THE NUMBER OF SPIRAL GALAXIES GALAXIES IN THREE BINS OF INCLINATION i

	$i \leq 30^\circ$	$31^\circ < i < 60^\circ$	$i > 61^\circ$	$i \leq 30^\circ$	$31^\circ < i < 60^\circ$	$i > 61^\circ$
Type	V-CV			SSRS2		
n	56	231	157	163	703	557
$n/n_{\rm total}$	0.13	0.52	0.35	0.12	0.49	0.39

ronments. However, contradictory results have also been obtained. Petrosian (1982), Dahari (1985), Laurikainen & Salo (1995), Dultzin-Hacyan et al. (1999), Storchi-Bergmann et al. (2001), Chatzichristou (2002), Sánchez & Gonzaléz-Serrano (2003), Kouluoridis et al. (2006) showed that Seyferts are found in more dense environments. Meanwhile, Viriani, De Robertis, & van Dalfsen (2000), Schmitt (2001, and references therein) did not find a relation between nuclear activity and environment. The discrepancy may be explained by the inclusion of galaxies with hidden Seyfert nuclei in the control samples.

The first attempt to estimate the number of missing Seyferts was made by Tovmassian (2001). In this paper we reconsider the problem, and estimate the number of missing Seyferts by the study of the distribution of inclinations of spiral galaxies in which Seyfert nuclei mainly reside. We take into account that, as Kinney et al. (2000) showed, the distribution of the angle β between the direction of the radio jet (or the dust torus axis) and the galaxy rotation axis is homogenous in the $0^\circ - 90^\circ$ range. This means that Seyferts should be observed in host galaxies with any inclination angle, if there are no biases in the catalogs. Kirhakos & Steiner (1990) suggested detecting the missing edge-on Seyferts on the basis of the IR and X-ray emission of galaxies. The presence of the broad line region (BLR) emission in some galaxies was detected by observation of the polarized emission scattered by particles located far from the nucleus of a galaxy. Such galaxies are called hidden Sy1s (Sy1h).

2. RESULTS AND DISCUSSION

2.1. *The data*

For our study we used the Sy1 and Sy2 galaxies from the 13-th Edition of the Catalogue of Quasars and Active Nuclei (Véron-Cetty & Véron 2010, hereafter VC-V) with magnitudes $V < 16^m$) and $z < 0.035$. Seyferts hosted in spiral galaxies of morphological types S0/a and later were included in the study. Morphological types of host galaxies of Seyferts were taken from the NED (The NASA/IPAC Extragalactic Database). In accordance with Meurs & Wilson (1984), Osterbrock & Shaw (1988), we joined galaxies with Sy1.2 and Sy1.5 types into the Sy1 category, and Sy1.8s and Sy1.9s into the Sy2 category. The sample of Sy1s includes galaxies of Sy1n type, and the sample of Sy2s includes galaxies classified as Sy1h (hidden Sy1). By cross-checking with the data in the NED we excluded from the studied sample those galaxies whose classification in the NED differed from those in the VC-V catalog. The compiled sample contains 79 Sy1s and 188 Sy2s.

For comparison we compiled two lists of spiral galaxies using the Revised Shapley-Ames catalog of bright galaxies (Sandage & Tammann 1981) and the Southern Sky Redshift Survey (SSRS2) (da Costa et al. 1998). From the latter catalog we used galaxies with radial velocities $V_r < 9000$ km s^{-1} and magnitudes $V < 15.5^m$. From both catalogs we selected spiral galaxies of morphological types from S0/a to Scd. The morphological types were taken from the NED. Galaxies with active galactic nuclei, such as Seyferts, LINERs or those having powerful radio emission were excluded. We did not include in the sample galaxies with any peculiar morphology or members of dense doubles.The compiled list from the Revised Shapley-Ames catalog contains 444 galaxies with measured a and b axes given in the NED. The corresponding list from the SSRS2 catalog consists of 1423 galaxies.

2.2. *Analysis and Results*

We calculated the inclinations i of the galaxies using their observed axial ratio b/a, and binned them in three intervals of inclination $i \leq 30^\circ$, $31^\circ < i < 60^\circ$ and $i > 61^\circ$. In Table 1 we present the numbers of spiral galaxies of the compiled lists from the Revised Shapley-Ames (1st row) and the SSRS2 (3rd row) catalogs in these intervals of inclination. The relative numbers for both catalogs in corresponding bins are given in Rows 2 and 4 of Table 1. Table 1