TABLE 2

THE NUMBERS AND RELATIVE NUMBERS OF SY1S AND SY2S GALAXIES IN THREE BINS OF INCLINATION i

Type	VC-V					
i	$N(<30^\circ)$	$N(31^\circ-60^\circ)$	$N(>61^\circ)$	$N(<30^\circ)/N_t$	$N(31^\circ-60^\circ)/N_t$	$N(>61^\circ)/N_t$
Sy1	20	49	10	0.25	0.62	0.13
Sy2	35	103	50	0.19	0.55	0.26

shows that the distribution of inclinations is not uniform which is evidence of the rather prolate space configuration of spiral galaxies (Vincent & Ryden 2005).

In the absence of biases one would expect the distribution of inclinations of the Seyferts to be the same as that of the normal galaxies, and they should also have about the same relative numbers in the three intervals of inclinations considered. However, we found that the distribution of inclinations of Seyfert galaxies differs from that of spiral galaxies not containing active nuclei (Table 2). A comparison of Table 1 and Table 2 shows that the relative number of both types of Seyferts in almost face-on galaxies ($i \leq 30^\circ$) is higher than in the whole sample of spiral galaxies. The relative number of Seyferts in the second bin is about the same, as for spiral galaxies. At the same time, the relative number of galaxies in the third bin of highly inclined Seyfert host galaxies ($i \geq 61^\circ$) is smaller than that of spirals.

2.3. *The number of missing Seyfert galaxies*

We suggest that the higher abundance of Seyferts seen in almost face-on host spiral galaxies and their paucity in galaxies with very high inclination (third bin) is due to absorption of the emission of the Seyfert nuclei by the dust disc of their host galaxies. In order to estimate the number of missing Seyferts we assume that the relative number of their host galaxies with $i \leq 30^\circ$ ($b/a < 0.86$) must be equal to the corresponding value of spiral galaxies, i.e. about 0.125. Dividing the numbers of Sy1s and Sy2s with inclinations $i \leq 30^\circ$ to the deduced relative number $n_i \leq 30^\circ$ of spiral galaxies with inclinations in the same bin, we estimate the expected total number n_t of Sy1s and Sy2s correspondingly. Subtracting from the estimated total number n_t of Seyferts the corresponding observed number n_o, we find the number n_m of the missing Seyferts. The results of our accounting are presented in Table 3. In the sixth column the relative number n_m/n_t is given.

TABLE 3

NUMBER OF SY1S AND SY2S WITH INCLINATIONS OF THE HOST SPIRAL GALAXIES $i \leq 30^\circ$, THEIR ESTIMATED TOTAL NUMBER, NUMBER OF MISSING ONES AND RELATIVE NUMBER OF MISSING SEYFERTS

	$n_{i\leq 30^\circ}$	n_t	n_o	n_m	n_m/n_o
Sy1	20	160	79	81	1.02
Sy2	35	280	188	92	0.49

Table 3 shows that the number of missing Sy1s is about the same as the number of the known ones. The number of the missing Sy2s is about half of the known ones.

3. CONCLUSIONS

We compared the distribution of inclinations of the host galaxies of Sy1 and Sy2 galaxies from the VC-V catalog with that of the control sample of spiral galaxies not containing AGN from the Revised Shapley-Ames and the SSRS2 catalogs, and estimated the number of Seyferts in which the AGN is hidden by the absorbing dust in the disc of inclined galaxies. We find that about 100% of Sy1s and 50% of Sy2s are missing. The relatively smaller number of missing Sy2s might be expected, since the NLR of some of them are observed at large distances from the galactic plane, where absorption of the dust disc is smaller. The number of missing Seyferts is large and may influence the conclusions on the nature of these galaxies and on the mechanisms of their formation. The relative number of Seyferts among galaxies is thus higher than previously estimated. Taking into account the missing Seyferts increases the opening angle of the dust torus from 90° to 100°. When considering the influence of the environment in activating the Seyfert nuclei, it is necessary to be careful in compiling a control sample: the inclusion of large numbers of highly inclined galaxies in the