TABLE 3

ASYMMETRY COMPARISON BETWEEN MODEL A AND OBSERVATIONS

Ion	Line (Å)	$v(\tau_{\max})/900$ km s^{-1}	Ramírez et al. (2005)[a]
S XVI	4.729	0.55 (R)	0.33±0.1 (B)
S XV	5.039	0.67 (R)	0.23±0.1 (B)
Si XIII	5.681	0.70 (R)	0.35±0.1 (B)
Si XIV	6.182	0.67 (R)	0.23±0.08 (B)
Si XIII	6.648	0.70 (R)	0.43±0.074 (B)
Mg XII	7.106	0.70 (R)	0.32±0.070 (B)
Mg XI	7.473	0.76 (R)	0.74±0.067 (R)
Fe XXIII	8.303	0.55 (R)	0.037±0.060 (B)
Mg XII	8.421	0.70 (R)	0.32±0.059 (B)
Mg XI	9.169	0.76 (R)	0.31±0.054 (B)
Ne X	9.708	0.76 (R)	0.45±0.052 (B)
Ne X	10.240	0.76 (R)	0.17±0.048 (B)
Ne IX	11.547	0.82 (R)	0.69±0.043 (R)
Fe XXII	11.770	0.55 (R)	0.22±0.042 (B)
Ne X	12.134	0.76 (R)	0.29±0.041 (B)
Fe XXI	12.284	0.62 (R)	0.42±0.040 (B)
Fe XVIII	14.373	0.70 (R)	0.27±0.035 (B)
Fe XVIII	14.534	0.70 (R)	0.51±0.034 (R)
O VIII	14.832	0.82 (R)	0.21±0.033 (B)
Fe XVII	15.015	0.70 (R)	0.76±0.033 (R)
O VIII	15.188	0.82 (R)	0.38±0.033 (B)
O VIII	16.006	0.82 (R)	0.26±0.030 (B)
O VII	17.200	0.90 (R)	0.29±0.029 (B)
O VII	17.396	0.90 (R)	0.30±0.029 (B)
O VII	17.768	0.90 (R)	0.32±0.028 (B)
O VII	18.627	0.90 (R)	0.32±0.027 (B)
O VIII	18.969	0.82 (R)	0.26±0.026 (B)

[a] $\bar{v}/v_1$, see definition in the text.

extended blue wings, as the flow moves outwards and acquires more velocity, in agreement with the observations. The second, which we call LIF (low ionization flow), characterized by $\log \xi_0 = 2$, is able to create the extended blue wing of the oxygen lines. This model has the fundamental ingredients that explain the asymmetry observed in the lines of NGC 3783 and, at the same time, gives us the bulk velocity of the flow.

In the third column of Table 5 we present the characterization of the lines coming from the two-outflow model. The HIF component is the major contributor to the formation of the high ionization lines, i.e. S XVI, Fe XXI, Fe XXII, Fe XXIII, because their optical depths peak close to $\log \xi_0$ (HIF), and as the velocity of the flow increases their ionization fraction decreases, forming their blue wings. Under this picture, only a few lines are characterized as blue, while the rest are red. Some intermediate-to-low-ionization ions have appreciable fractions in the LIF component, which broadens the lines and yields

TABLE 4

PARAMETERS OF THE TWO OUTFLOW MODEL (MODEL C)

Outflow HIF	Outflow LIF
$\log \xi_0 = 3.0$	$\log \xi_0 = 2.0$
$v_\infty = 1100$ km s^{-1}	$v_\infty = 2200$ km s^{-1}
$N_H = 10^{22}$ cm^{-2}	$N_H = 10^{21}$ cm^{-2}
$w_0 = 0.4$	$w_0 = 0.2$
Solar composition	Solar composition