TABLE 1

COEFFICIENTS OF THE POLYNOMIAL GIVEN BY EQUATION (3)

R (kpc)	$0-0.765$	$0.765-2.9$	$2.9-3.825$	$3.825-13$	>13
a_0	0.0	325.0912	329.8	−2346.0	230.6
a_1	3069.81	−248.1467	−250.1	2507.60391	...
a_2	−15809.8	231.87099	231.87099	−1024.068760	...
a_3	43980.1	−110.73531	−110.73531	224.562732	...
a_4	−68287.3	25.073006	25.073006	−28.4080026	...
a_5	54904.0	−2.110625	−2.110625	2.0697271	...
a_6	−17731.0	...	...	−0.080508084	...
a_7	...	...	...	0.00129348	...

TABLE 2

COEFFICIENTS OF EQUATION (4)

	a	b
U	0.39	1.31
V	0.40	1.10
W	0.53	0.94
Total	0.40	1.40

The recent Geneva-Copenhagen Survey of the Solar Neighborhood (cf. Nordström et al. 2004; Holmberg, Nordström, & Andersen 2007, 2009) has considerably improved the relations involving the ages, kinematics, and chemical composition of a large sample containing about 14000 F and G nearby stars. Using basically the original *Hipparcos* parallaxes, $uvby-\beta$ photometry and the Padova stellar evolution models, several basic relations were investigated. In particular, high correlations have been obtained between the velocity dispersions σ_U, σ_V, σ_W, and σ_T and the ages of the stars, which clearly show a smooth increase of the velocity dispersions (of the U, V, W components and the total velocity T) with time. From the calibration by Holmberg et al. (2009) these correlations can be approximately written as

$$\log \sigma = a \log t + b\,, \tag{4}$$

where the age t is expressed in Gyr and the constants a, b are given in Table 2. This approximation is valid in the age interval $0 < t(\mathrm{Gyr}) < 14$ with an estimated average uncertainty of about 25%. Method 1 consists of assuming that the discrepancy in the rotation velocity $\Delta\theta$ is due to the evolved status of the CSPN, so that we should expect a correlation between $\Delta\theta$ and the velocity dispersion, as given by equation (4). Since in this method we are using the rotation velocity, we have considered two possibilities, according to which the velocity discrepancy $\Delta\theta$ can be associated with (i) the V component of the total velocity (σ_V), or (ii) the total velocity (σ_T). Moreover, since we are adopting two distance scales and two theoretical rotation curves, we have 8 different age distributions for Method 1, characterized by the timescales t_1 to t_8, as explained in Table 3.

TABLE 3

PARAMETERS FOR METHOD 1

Distance	Rotation Curve	Dispersion	Age
Maciel	PN	σ_V	t_1
Maciel	PN	σ_T	t_2
Maciel	Clemens	σ_V	t_3
Maciel	Clemens	σ_T	t_4
Stanghellini	PN	σ_V	t_5
Stanghellini	PN	σ_T	t_6
Stanghellini	Clemens	σ_V	t_7
Stanghellini	Clemens	σ_T	t_8

2.2. *Method 2: The U, V, W velocity components*

Method 2 is also a kinematic method, and in principle more accurate than Method 1, as discussed in more detail in § 4. From the PN radial velocities and distances, we have estimated their proper motions both in right ascension and declination, μ_α and μ_δ. We have assumed that, on average, the tangential velocities are similar to the radial velocities, namely $V_t \simeq V_r$. In view of the large distances of the nebulae, this hypothesis in practice does not introduce any major uncertainties in the