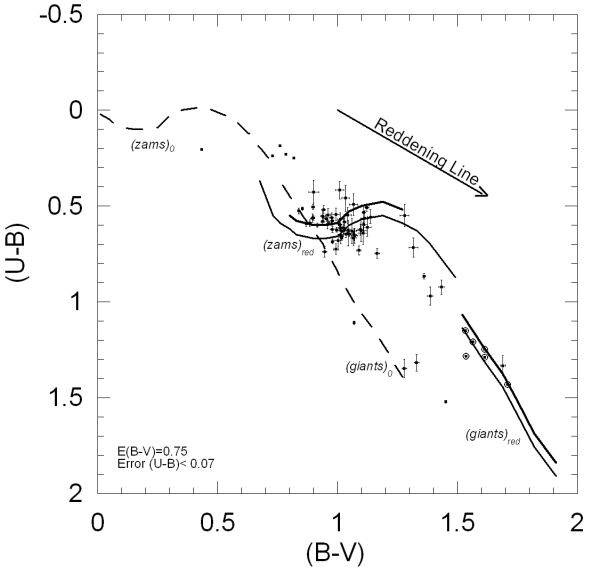

Fig. 8. The $(U-B,\ B-V)$ plot of Be 10. The symbols and the curves are the same as in Figure 2.

moduli, $(V_0 - M_V)$, and ages, A, found from these five CM diagrams and their respective isochrone fittings are given in Table 8.

In these CM diagrams a significant number of stars are seen extending to brighter magnitudes and red-ward from the fainter and redder observational limits of the main sequences, i.e., the stars extending red-ward and upward from $(V,\ B-V) \approx (18\overset{\mathrm{m}}{.}5, 1\overset{\mathrm{m}}{.}5)$ or $(V,\ R-I) \approx (18\overset{\mathrm{m}}{.}5, 0\overset{\mathrm{m}}{.}9)$ (cf. Figure 6). These are probably field red-giant stars contributed by the Galactic bulge, as suggested by the Galactic longitude and latitude of Ru 135, $\ell \simeq 16.4°$ and $b \simeq +6.2°$ (see Binney & Merrifield 1998; Stanek et al. 1996, Figures 3.5 and 2, respectively). The fact that Ru 135 lies near the direction of the Galactic central region also explains the significant number of brighter and bluer foreground stars seen in its CC and CM diagrams.

3.3. *Be 10*

In Figure 8 the loci of stars observed in the direction of Be 10 are shown in the $(U-B,\ B-V)$ diagram, together with the standard interstellar reddening vector and the two-color curve of SK82, shifted along this vector to procure the best fit to the data. From the fits along the $(B-V)$ and $(U-B)$ axes, $E(B-V) = 0\overset{\mathrm{m}}{.}75 \pm 0\overset{\mathrm{m}}{.}09$ and ([Fe/H], Z) $= (-0.49 \pm 0.02$ dex, $+0.006 \pm 0.0003)$ are found, following the procedures described in § 3 and § 3.1 (see Table 8 for partial and mean results). Again, the appropriate isochrones of MGBG have been computed online with this corresponding metallicity and are used below for further analyses of Be 10.

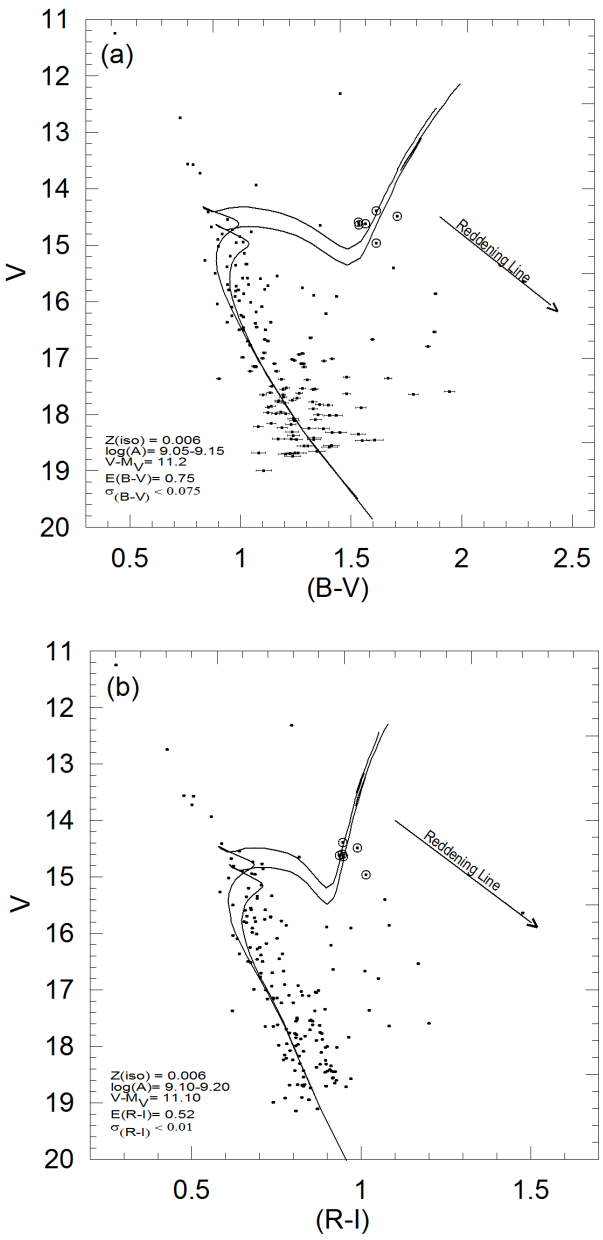

Fig. 9. The $(V,\ B-V)$ and $(V,\ R-I)$ diagrams for Be 10. Solid lines show the isochrones of MGBG interpolated to $Z = +0.006$. The larger open circles identify the RC candidates. See the text, and Tables 7 and 8, for the inferred values for reddening, metallicity, distance modulus, and age.

For Be 10, $DM = (V - 3.1 \cdot E(B-V) - M_V) = 11\overset{\mathrm{m}}{.}16 \pm 0\overset{\mathrm{m}}{.}06$, the distance, $d = 1.7 \pm 0.05$ kpc, the metallicity, $Z = +0.006 \pm 0.0003$, $\log(A) = 9.06 \pm 0.05$, and the age, $A = 1.08 \pm 0.08$ Gyr have