TABLE 1

MIR PHOTOMETRY FOR EXTENDED GLIMPSE 3D PLANETARY NEBULAE

PN G	NAME	CAT	3.6 mag	4.5 mag	5.8 mag	8.0 mag
002.8–02.2	Pe 2-12	ACK	8.59	8.13	6.44	4.75
003.5–02.4	IC 4673	ACK	9.47	7.66	8.45	5.64
003.6–02.3	M 2-26	ACK	10.3	9.54	9.95	7.27
003.9–02.3	M 1-35	ACK	9.06	8.02	7.85	5.17
004.2–02.5	PHR1805-2631	M1	11.68	10.43	11.09	9.64
009.4–01.2	PHR1811-2123	M1	9.27	9.1	9.28	7.88
009.7–01.1	PHR1811-2102	M1	11.9	11.35	10.73	9.53
009.8–01.1	PHR1811-2100	M1	10.11	9.28	8.79	7.75
010.8–01.8	NGC 6578	ACK	7.81	6.99	7.43	5.06
025.9–02.1	Pe 1-15	ACK	11.17	9.83	10.29	7.51
330.6–02.1	He 2-153	ACK	9.63	8.76	8.57	6.86
335.4–01.9	PHR1637-4957	M1	9.83	9.1	9.22	7.86
336.2+01.9	Pe 1- 6	ACK	10.39	9.19	9.78	6.91
344.2–01.2	H 1- 6	ACK	9.36	8.02	7.65	6.01
344.4+01.8	MPA1654-4041	M2	11.2	10.41	11.21	9.2
345.2–01.2	H 1- 7	ACK	6.67	6.18	4.79	2.81
350.4+02.0	PHR1712-3543	M1	10.92	9.75	10.14	8.2
356.0–01.4	PPA1741-3302	M1	11.29	10.63	11.09	8.42
356.5–02.3	M 1-27	ACK	8.09	7.36	6.67	4.39
356.5–01.8	PPA1744-3252	M1	10.5	10.15	10.82	8.82
357.3–02.0	PPA1747-3215	M1	9.77	9.65	8.62	7.2
357.5–02.4	PPA1749-3216	M1	10.19	9.87	10.61	8.45
359.1–01.7	M 1-29	ACK	8.19	6.99	6.93	4.57
359.4–03.4	H 2-33	ACK	11.3	10.17	10.6	7.57

where A is a constant, $n = 1$ corresponds to the lowest (i.e. the outermost) contour level, B is the background, and C is the logarithmic separation between contours. Contour parameters A, B and C are indicated in the relevant figure captions.

We have additionally produced profiles through the sources after the elimination of background emission (particularly strong in the 5.8 and 8.0 μm IRAC channels), and the removal of linear spatial gradients. These results were then processed to indicate the variation of 3.6 μm/4.5 μm, 5.8 μm/4.5 μm and 8.0 μm/4.5 μm ratios with distance from the nebular nuclei. The rationale behind this is based on the fact that many PNe possess strong polycyclic aromatic hydrocarbon (PAH) emission bands at 3.3, 6.2, 7.7 and 8.6 μm, located in the 3.6, 5.8 and 8.0 μm IRAC bands. Given that the 4.5 μm band is usually dominated by the bremsstrahlung continuum, as well as by a variety of molecular and ionic transitions, such ratios enable us to gain an insight into the distribution of PAHs within the nebular shells (see e.g., Phillips & Ramos-Larios 2008a,b).

Contour mapping of the 3.6 μm/4.5 μm, 5.8 μm/4.5 μm and 8.0 μm/4.5 μm flux ratios was undertaken by estimating levels of background emission, removing these from the images, and subsequently setting values at < 3 rms noise levels to zero. The maps were then ratiod on a pixel-by-pixel basis, and the results contoured using standard IRAF programs. Contour levels are given through $R_n = D10^{(n-1)E}$, where the parameters (D, E) are again detailed in the figure captions.

The IRAC fluxes of extended sources are open to errors deriving from instrumental scattering, and require to be corrected by factors of 0.91 at 3.6 μm, 0.94 at 4.5 μm, 0.66–0.73 at 5.8 μm (we take a mean value of 0.70), and 0.74 at 8.0 μm – although this correction also depends upon the surface brightness