TABLE 2

DEDUCED PARAMETERS FOR CLUSTERS ANALYZED BY JHK_S PHOTOMETRY

Cluster	$E(J–H)$	$E(B–V)$	$J–M_J$	$V_0–M_V$	d (kpc)
Berkeley 44	0.295 ± 0.02	1.00 ± 0.07	12.1 ± 0.1	11.38 ± 0.24	1.89 ± 0.21
Turner 1a	0.15 ± 0.02	0.51 ± 0.07	9.7 ± 0.2	9.34 ± 0.29	0.74 ± 0.08
Turner 1b	0.29 ± 0.02	1.02 ± 0.07	13.2 ± 0.2	12.47 ± 0.29	3.12 ± 0.42
Collinder 419	0.39 ± 0.02	1.32 ± 0.07	11.7 ± 0.2	10.75 ± 0.28	1.42 ± 0.18

Turner et al. (1986). In each case the optical photometry was limited by the faintness of cluster members arising from large interstellar reddening, which is where JHK_s photometry has advantages over optical band photometry. The point to emphasize, however, is that JHK_s observations of cluster stars often provide relatively complete samples that can be used to infer reasonably accurate estimates for the all-important interstellar reddening toward the cluster, in most cases in more straightforward manner than what is often used for optical band $BVRI$ observations. Once the reddening is known, ZAMS fitting, or isochrone fitting in some cases, can then be used to derive the distance to the cluster. That is true despite the relatively low intrinsic precision of existing 2MASS data for most Galactic fields.

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