tem (FITS) data format[3]. This online survey sampler of the *Max-Planck-Institut für Radioastronomie* (MPIfR) near Bonn, Germany, allows users to pick a region of the sky and obtain images and data at different frequencies. The 2720-MHz Stockert survey has an angular resolution of 0°.35, the 1420-MHz Stockert survey (Reich & Reich 1986) 0°.59, the 820-MHz Dwingeloo survey (Berkhuijsen 1972) 1°.2, the 408-MHz all-sky survey (Haslam et al. 1982) 0°.85, and the 34.5-MHz Gauribidanur survey (Dwarakanath & Udaya Shankar 1990) 0°.7. The corresponding observations are given at the following rates (measured data) for both l and b: 1°/8 at 2720 MHz, 1°/4 at 1420 MHz, 1°/2 at 820 MHz, 1°/3 at 408 MHz and 1°/5 at 34.5 MHz. The effective sensitivities are about 5 mK T_b (T_b is an average brightness temperature), 50 mK, 0.20 K, 1.0 K and about 700 K, respectively.

FITS data format stores the multidimensional arrays and 2-dimensional tables containing rows and columns of scientific data sets. We extracted observed brightness temperatures from this data format into ASCII data files, and afterwards, using our programs in C and FORTRAN, we obtained the results presented in this paper.

2.2. *Method*

As there is a great influence of background radiation over the Cygnus loop area, it is very difficult to determine its area precisely. The area of this loop is enclosed with brightness temperature contours. The maps of a region in Cygnus, in new Galactic coordinates (l, b), with contours of the brightness temperatures T_b are plotted in Figures 1–5. In these figures, among all the brightness temperature contours, the most important are the outer contour and inner one, which represent the loop borders. The outer one (which corresponds to the minimum temperature of the loop) separates loop and background, while the inner one (corresponding to the maximum temperature) separates the loop and some superposed source. The interval of Galactic longitude of this loop is l=[76.5°, 71.5°], and of latitude it is b=[−10.5°, −7°]. We used the same method of calculation as given in Borka (2007) for Galactic radio loops I–VI, Borka, Milogradov-Turin, & Urošević (2008) for Loops V and VI and in Borka Jovanović & Urošević 2009b for the Monoceros loop. Our aim is to apply our method, which we developed for main Galactic loops I-VI and described in Borka (2007) and Borka et al. (2008), to smaller remnants and to show that it is applicable to almost all remnants. To confirm that our method works well, we also provide the area maps of Cygnus at the five frequencies (Figures 1–5) and some examples of the temperature profiles at 2720 MHz (Figure 6). In Figure 6 it should be noticed that the bottom panel shows only back-

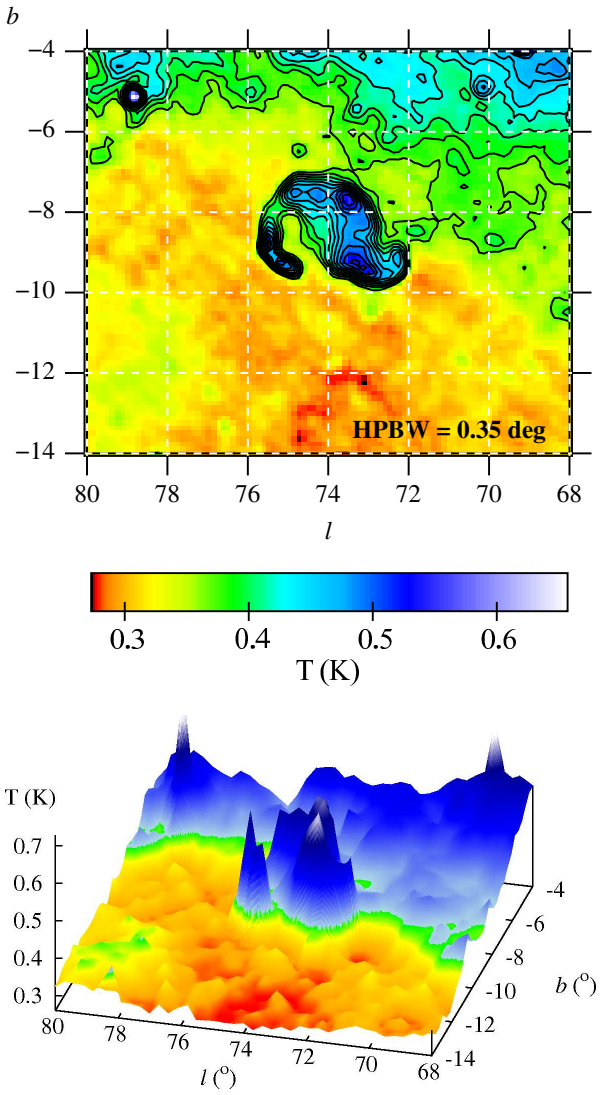

Fig. 1. Top: the 2720 MHz map of a region in Cygnus, in new Galactic coordinates (l, b), showing contours of brightness temperature. This radio loop has position: l = [76.5°, 71.5°]; b = [−10.5°, −7°]. The HPBW (Half Power Beam Width) for this frequency is 0°.35. The contours plotted represent the temperatures $T_{\rm min}$ and $T_{\rm max}$ from Table 1 and nine contours in between. The contours are plotted every 0.265 K, starting from the lowest temperature of 0.395 K and up to 0.66 K. The corresponding temperature scale is given (in K). Bottom: the 2720 MHz area map of Cygnus. The color figure can be viewed online.

[3]Available at the site `http://www.mpifr-bonn.mpg.de/survey.html`.