TABLE 3

$uvby-\beta$ PHOTOELECTRIC PHOTOMETRY OF NGC 1647

WEBDA	V	$b-y$	m_1	c_1	$[m_1]$	$[c_1]$	β	MK	ST
2	7.465	0.996	0.510	0.245	0.828	0.045	2.893	K0	late
10	11.052	0.398	0.035	0.979	0.162	0.899	2.892	B7.5IV	B9
68	10.150	0.361	0.117	0.995	0.232	0.922	2.889		A3
71	10.671	0.312	0.054	0.941	0.153	0.878	2.864	A0	B9
34	10.593	0.257	0.061	0.900	0.143	0.848	2.851	B9	B9
42	9.667	0.246	0.077	0.912	0.155	0.862	2.850	B9.5IV	B9
29	10.674	0.259	0.062	0.967	0.144	0.915	2.838	A0 IV	B9
49	10.048	0.290	0.083	0.945	0.175	0.887	2.832	B8	B8
31	10.264	0.219	0.082	0.905	0.152	0.861	2.794	B8	B9
48	10.402	0.247	0.075	0.825	0.154	0.775	2.778	A0	B8
5	10.220	0.358	0.056	0.660	0.170	0.588	2.774	A0	B6
37	8.864	0.246	0.040	0.651	0.118	0.601	2.774	B9.5IV	B5
57	9.985	0.333	0.041	0.691	0.147	0.624	2.772	A0	B5
70	10.466	0.288	0.062	0.707	0.154	0.649	2.767	B9	B6
419	9.811	0.197	0.047	0.629	0.110	0.589	2.756		B4
51	9.943	0.303	0.080	0.766	0.177	0.705	2.753	B9 IV	B7
99	10.059	0.327	0.044	0.794	0.148	0.728	2.753	B9 IV	B7
54	10.075	0.290	0.053	0.654	0.145	0.596	2.746	A0 V	B5
22	9.067	0.199	0.042	0.739	0.105	0.699	2.742	B9 II	B6
59	9.698	0.400	0.008	0.773	0.136	0.693	2.738	A0	B7
3	10.026	0.290	0.055	0.656	0.147	0.598	2.736	B7.5IV	B4
94	9.667	0.216	0.047	0.666	0.116	0.622	2.726	B9 III	B5
55	10.246	0.341	0.040	0.632	0.149	0.563	2.724	B7 V	B4
66	10.246	0.350	0.017	0.913	0.129	0.843	2.721	B8	B8
233	10.136	0.254	0.092	0.571	0.173	0.520	2.708		B4
102	9.307	0.281	0.026	0.770	0.115	0.713	2.702	B9 III	B6
45	8.860	0.248	0.039	0.645	0.118	0.595	2.700	B7.5IV	B4
44	9.213	0.231	0.033	0.769	0.106	0.722	2.697	B8 IV	B8
15	8.561	0.304	0.020	0.644	0.117	0.583	2.620	B8 III	B4
65	9.561	0.364	0.153	0.34	0.269	0.267	2.611	G0	late
105	8.391	1.048	0.556	0.214	0.891	0.004	2.585	K2	late
1	5.963	0.727	0.694	0.280	0.926	0.134	2.562	K0	late
84	8.779	1.084	0.840	0.088	1.186	−0.128	2.553	K5	late
16	9.480	0.496	0.232	0.449	0.39	0.349	2.548	K0 V	late

early (B or early A) or late (late A or F) types; the later-class stars (later than G) were not considered in the analysis since no reddening determination calibration has yet been developed for such MS stars. Further analyses were done considering the numerical criteria of Lindroos (1980) for Strömgren photometry.

4. RESULTS

The application of the above mentioned numerical packages gave the results listed in Table 5 for NGC 1647 and Table 6 for NGC 1778. The following values are listed: the WEBDA identificator, the reddening, $E(b-y)$, the reddening-free indexes (in decreasing β values), the absolute magnitude, the distance modulus, the distance (in parsecs), the effective temperature and the surface gravity. When histograms of the distances are drawn (Figure 3a), one can see that 24 of the 35 observed stars in NGC 1647 lie around a distance of 564 pc corresponding to a distance modulus of 8.75 ± 0.09. However, when the histogram of distances to the stars is plotted for